## Comparison of nitric oxide measurements in the mesosphere and lower thermosphere from ACE-FTS, MIPAS, SCIAMACHY, and SMR

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Received: 8 October 2014 - Accepted: 23 November 2014 - Published: 18 December 2014
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Published by Copernicus Publications on behalf of the European Geosciences Union.

## AMTD

7, 12735-12794, 2014

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## Abstract

We compare the nitric oxide measurements in the mesosphere and lower thermosphere ( 60 to 150 km ) from four instruments: ACE-FTS, MIPAS, SCIAMACHY, and SMR. We use the daily zonal mean data in that altitude range for the years 2004-2010 5 (ACE-FTS), 2005-2012 (MIPAS), 2008-2012 (SCIAMACHY), and 2003-2012 (SMR).

We first compare the data qualitatively with respect to the morphology, focussing on the major features, and then compare the time series directly and quantitatively. In three geographical regions, we compare the vertical density profiles on coincident measurement days. Since none of the instruments delivers continuous daily measurements in this altitude region, we carried out a multi-linear regression analysis. This regression analysis considers annual and semi-annual variability in form of harmonic terms and inter-annual variability by responding linearly to the solar Lyman- $\alpha$ radiation index and the geomagnetic Kp index. This analysis helps to find similarities and differences in the individual data sets with respect to the inter-annual variations caused by geomagnetic and solar variability.

We find that the data sets are consistent and that they only disagree on minor aspects. SMR and ACE-FTS deliver the longest time series in the mesosphere and they both agree remarkably well. The shorter time series from MIPAS and SCIAMACHY also agree with them where they overlap. The data agree within ten to twenty percent when the number densities are large, but they can differ by 50 to $100 \%$ in some cases.

## 1 Introduction

Climate models aim to predict the trend of Earth's climate considering the composition of the atmosphere. This composition is influenced by a number of factors, including anthropogenic emissions and solar variability. To disentangle these effects, the evaluation of the solar influence is important. Solar particles and soft solar X-rays produce nitric oxide (NO) in the mesosphere and lower thermosphere (MLT, 50-150 km) (Barth

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et al., 2003). Thus, the NO content in this atmospheric region indicates how much solar activity impacts the atmospheric composition and how important solar variability is in climate models.

To relate atmospheric composition changes to solar activity, global NO measure5 ments over long time periods deliver important information. This data is provided by satellite instruments using different measurement methods. The consistency of these measurements is crucial for using the results for further work, for example to validate climate models and to find climate relevant forcing parameters. We compare the daily zonal mean NO number densities from four space-borne instruments: the Michelson
10 Interferometer for Passive Atmospheric Sounding (MIPAS, infrared limb emission) and the SCanning Imaging Absorption spectroMeter for Atmospheric CHartographY (SCIAMACHY, UV-vis-NIR limb and nadir emission) on Envisat, the Sub-Millimetre Radiometer (SMR, sub-mm limb emission) on the Odin satellite, and the Atmospheric Chemistry Experiment-Fourier Transform Spectrometer (ACE-FTS, infrared solar occultation) on 15 SCISAT. The aim of this comparison is to answer the question whether zonal mean data sets from different instruments consistently constrain the nitric oxide in the MLT.

The instruments are introduced in Sect. 2. In a first comparison step, we compare zonal daily mean distributions in an altitude range from 75 to 115 km (Sect. 3). In that section we analyse the time-latitude morphology of the NO number densities at different times of solar activity, from low (2008/09) to moderate (2010/11). In Sect. 4 we directly compare the time series of NO number densities at selected geographic locations. This more quantitative comparison emphasizes differences which cannot be detected in the morphological comparison. We further compare daily mean vertical profiles in order to detect differences that may result from the different retrieval strategies (Sect. 5). In Sect. 6, we extend the comparison towards derived diagnostic quantities of particular scientific interest, for example regression coefficients for the contribution of Lyman- $\alpha$ and the geomagnetic Kp index. In the Conclusions (Sect. 7) we evaluate the consistency of the analysed data sets.

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## 2 Instruments

### 2.1 Envisat/MIPAS

Two of the instruments considered here, MIPAS and SCIAMACHY, are on board the now defunct European research satellite Envisat. This satellite has been orbiting on 5 a sun-synchronous orbit at 800 km and at equator crossing times 10:00/22:00 since 2002. Communication to the satellite was lost in April 2012, which is therefore the latest date for which MIPAS and SCIAMACHY data are available.

The Michelson Interferometer for Passive Atmospheric Sounding (MIPAS) is an infrared Fourier transform spectrometer. It has a spectral range from 4.15 to $14.6 \mu \mathrm{~m}$
$10 \quad\left(685-2410 \mathrm{~cm}^{-1}\right)$, and the spectral resolution was $0.0625 \mathrm{~cm}^{-1}$ during the relevant time period. The instrument has an instantaneous field of view of 3 km in the vertical and 30 km in the horizontal direction.

MIPAS measured atmospheric emissions in a limb observation geometry (Fischer et al. 2008). The instrument provided two special observation modes dedicated 5 to the middle atmosphere (MA, 18-102 km) and upper atmosphere (UA, 42172 km ) (Raspollini et al., 2013). After a test phase of twelve measurement days from January 2005 to October 2007, about one day every three months, the measurements using these modes were scheduled regularly on two days every 10 days of nominal mode measurements ( $7-72 \mathrm{~km}$ ) since November 2007. In this study, we use only the day side half-orbit (downleg, am) MIPAS-UA measurements, which amount to 199 measurement days with about 500 scans per day from 21 January 2005 until 30 March 2012.

The NO data used here were produced using the MIPAS data processor developed at the Institute of Meteorology and Climate Research (IMK) in cooperation with the Instituto de Astrofísica de Andalucía (IAA) (von Clarmann et al., 2003). The retrieval is based on constrained multi-parameter non-linear least squares fitting of observed to calculated radiance spectra. The retrieval processor was extended to apply to non-local thermodynamic equilibrium (Non-LTE) emissions (Funke et al., 2001), which is partic-

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ularly important for NO retrievals. The vibrational, rotational, and spin level populations under Non-LTE were modelled by the Generic RAdiative traNsfer AnD non-LTE population Algorithm (GRANADA) (Funke et al., 2012) during each iteration of the retrieval.

NO in the altitude region of interest ( $70-120 \mathrm{~km}$ ) is derived from the fundamental NO band emission at $5.3 \mu \mathrm{~m}$. The retrieval of NO in the middle atmosphere (up to $\sim 100 \mathrm{~km}$ ) is described in detail in Funke et al. (2005). The NO data used here (Versions V5r_NO_520 (MA) and V5r_NO_620 (UA)) were measured from January 2005 to April 2012, when MIPAS used a slightly degraded spectral resolution. The data version described in Funke et al. (2005) has been substantially improved as summarised in Funke et al. (2014), and the present version includes these improvements.

The single profile vertical resolution of NO in the $70-100 \mathrm{~km}$ region is $15-20 \mathrm{~km}$ and the single profile precision ranges from $\sim 50 \%$ at 70 km to $\sim 30 \%$ at 100 km . In general, MIPAS has low sensitivity to NO outside of the polar winter region (latitudes greater than $50^{\circ}$ ) at altitudes of $70-90 \mathrm{~km}$. Systematic errors in the NO VMR in this region amount to around $10 \%$.

The retrieval of NO in the upper atmosphere ( $100-170 \mathrm{~km}$ ) is described in detail in Bermejo-Pantaleón et al. (2011). Temperature and nitric oxide mixing ratio profiles are retrieved jointly in this region. The typical single measurement precision of NO in this region is $10-30 \%$ for high geomagnetic activity, increasing to $20-50 \%$ for low geomagnetic activity. The vertical resolution is $5-10 \mathrm{~km}$ for high geomagnetic activity and degrades to $10-20 \mathrm{~km}$ for low geomagnetic activity. For extra-polar and low geomagnetic activity, a potential systematic bias in the nighttime NO version V4O_NO_611 profiles was identified (Bermejo-Pantaleón et al., 2011). This bias is caused by smoothing errors and was corrected in the current version V5r_NO_620 by using an appropriate a priori NO profile for nighttime conditions. The retrieved NO abundances in the thermosphere depend strongly on the assumed atomic oxygen above 120 km , which was taken from the NRL-MSISE-00 model. Other systematic errors as the uncertainties in the atomic nitrogen and the propensity for spin-conserving collisions, which largely

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control thermospheric spin non-LTE distributions, lead to minor systematic errors of 5-10\%.

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every 2 s which corresponds to a vertical spacing of 2-6 km (depending on beta angle). The vertical resolution of ACE-FTS is about 3 km based on the instrument field-of-view of 1.25 mrad .

ACE-FTS retrievals of NO use 36 microwindows between 5.18 and $5.43 \mu \mathrm{~m}, 19$ of 5 which are focused on the higher altitude ranges. The non-linear least squares fitting scheme used for ACE-FTS is described in Boone et al. (2005). Here we use version 3.0 of the retrievals which extend from about 6 to 107 km (Boone et al., 2013). The previous version of the ACE-FTS NO product (version $2.2+$ updates) has been compared with HALOE profiles where differences of up to $+10 \%$ are found between 93 and 105 km (Kerzenmacher et al., 2008). Compared to version 2.2+updates, version 3.0 NO profiles are lower by about $2 \%$ above 40 km (Waymark et al., 2013). The ACEFTS profiles are provided in two formats: on the measurement tangent grid and interpolated onto a uniform 1 km grid (using piecewise quadratic technique). In this study, we use the sunrise data on the uniform 1 km grid. 580 km and varying equator crossing times between 6 and 7 a.m. and p.m. (LT). These parameters are slightly changing with time due to the drifting orbit. Odin was initially a joint astronomy and aeronomy mission and before 2007, the observation time was equally divided between the two disciplines. The satellite is now entirely dedicated to atmospheric measurements.
The Sub-Millimetre Radiometer (SMR) is one of the instruments on board this satellite. It is a limb emission sounder measuring globally a variety of trace gases and the temperature in the whole middle atmosphere. SMR uses four sub-millimetre channels ( $486.1-503.9 \mathrm{GHz}, 541.0-558.0 \mathrm{GHz}, 547.0-564.0 \mathrm{GHz}, 563.0-581.4 \mathrm{GHz}$ ) and one

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millimetre wave channel ( $118.25-119.25 \mathrm{GHz}$ ) (Merino et al., 2002). The observation of different species requires channel switching. Nitric oxide is retrieved from the observation of thermal emission lines in a band centred around 551.7 GHz (Urban et al., 2007; Sheese et al., 2013). In the case of SMR, $\log (V M R)$ is retrieved, see Sheese et al. (2013) for resolution, uncertainties and more details. Pérot et al. (2014) also describes the characteristics of the SMR NO measurements. We use the data version 2.1 in this study and SMR contributes NO observations on 301 days from 7 October 2003 to 25 December 2012, with about 600 scans per day.

### 2.5 Data preparation

The individual measurements of nitric oxide ( NO ) from each instrument were averaged to daily zonal mean values binned into $5^{\circ}$ latitude bins. In the vertical direction, no additional interpolation was done for the MIPAS and SCIAMACHY data and the fixed retrieval altitude grids were used. The SMR data were retrieved on the tangent altitude grid, profile-by-profile, interpolated onto a fine altitude grid, followed by averaging in $5^{\circ}$ latitude bins on each level. As mentioned in Sect. 2.3, the ACE-FTS data on the interpolated 1 km grid were used in this study.

To prepare this data, volume mixing ratios had to be converted to number densities. The ACE-FTS VMR were converted to number density using the density from the simultaneously retrieved temperatures (Boone et al., 2005). The MIPAS data, which ${ }_{20}$ were retrieved as $\log (\mathrm{VMR})$, were also converted using the simultaneously retrieved temperatures (Funke et al., 2005; Bermejo-Pantaleón et al., 2011). The SMR data were also retrieved as log(VMR) using ECMWF analyses and CIRA temperature fields above 70 km (Barnett and Corney, 1985). The mixing ratios were converted to number densities using the same temperatures. Note that we only use the day time (a.m.) data versions from MIPAS and SMR and the sunrise data from ACE-FTS in this study.

The averaged data were additionally filtered by the sensitivity of the instruments. The ACE-FTS measurements have been filtered based on the Data Issues List (https: //databace.scisat.ca/validation/data_issues.php) to remove profiles marked as "do not

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use". The MIPAS and SCIAMACHY data are filtered based on the average diagonal element of the averaging kernel matrix. The thresholds are 0.003 for MIPAS and 0.01 for SCIAMACHY. The SMR data are filtered using the measurement response which is the sum of the rows of the averaging kernel matrix. All zonal mean boxes with a mean 5 measurement response below 0.75 were not considered.

## 3 Zonal mean data

Figures 1 and 2 show the zonal mean data at altitudes of 85 and 105 km from all instruments: ACE-FTS, MIPAS, SCIAMACHY, and SMR. The zonal mean data at additional altitudes from 75 to 115 km are shown in Figs. A1 to A5.

These figures give an overview of the available data set. ACE-FTS, as a solar occultation instrument, has only limited geographical coverage. MIPAS and SMR have limited sensitivity at altitudes below 85 km , in particular at middle and low latitudes. Additionally, MIPAS data from 75 to 100 km is at present only available since July 2008. The SCIAMACHY data are restricted to daytime measurements which adversely affects the number of measurements, in particular at high latitudes in the polar winter. SMR data are relatively sparse before 2007 when the Odin astronomy mission ended and more measurement days were dedicated to NO observations since then.

The zonal mean data of all four instruments are consistent with respect to the annual variation of the NO density in the MLT region. Throughout the latitude range, the number densities are low at times of low solar activity, 2008 and 2009, and increase with growing solar activity, 2010 and 2011. The NO density increases most in the polar regions and at higher altitudes. Between 95 and 115 km , the density increases also at lower latitudes, in particular after 2011.

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## 4 Time series

To put the comparisons from Sect. 3 onto more quantitative grounds, we directly compare the NO density time series in four different $5^{\circ}$ latitude bins and at altitudes from 75 to 115 km . These bins are located at high northern ( $67.5^{\circ} \mathrm{N}$ ) and high southern 5 latitudes $\left(67.5^{\circ} \mathrm{S}\right)$, at middle latitudes $\left(32.5^{\circ} \mathrm{S}\right)$, and at low latitudes $\left(2.5^{\circ} \mathrm{N}\right)$ near the equator.

### 4.1 Northern Hemisphere

We first compare the values in the Northern Hemisphere at $67.5^{\circ} \mathrm{N}$. Figure 3 shows the time series at two example altitudes, 85 km (upper mesosphere) and 105 km (lower thermosphere). Figure B1 shows the time series at more altitudes from 75 to 115 km .

The results are consistent throughout the altitude range, the largest values are observed between 95 and 105 km and smaller values below and above these altitudes. Only SMR and ACE-FTS provide data below 100 km for the years 2004 to 2008. Above 100 km , MIPAS contributes some data points from 2005 onwards, and SCIAMACHY data are available only from mid-2008.

ACE-FTS, SMR, and MIPAS show that the NO number density is correlated with solar activity. From 2004 to 2007, a period of moderate solar activity, the number densities were generally larger than in 2008/09, when solar activity was low. The NO density increases then again after 2009 with the onset of the next solar cycle. This correlation is visible at all altitudes and it is particularly strong in the main production region from 95 to 105 km . Unfortunately, the SCIAMACHY data set is too short to show the same correlation. The SCIAMACHY number densities are always on the low side compared to the other instruments. This is less pronounced at 75 km but is clearly visible at altitudes of 85 km and above.

In addition to the overall correlation of the NO densities with the long-term solar activity, the seasonal cycle is clearly visible in the data from all instruments. This annual variation is more pronounced at 85 km , but it is also visible at 105 km .

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### 4.2 Southern Hemisphere

Figure 4 shows the NO number density at $67.5^{\circ} \mathrm{S}$ at 85 and 105 km . The time series comparison at more altitudes from 75 to 115 km is shown in Fig. B2. The values agree at the same level as in the Northern Hemisphere and the density responds similarly to solar activity.

The annual cycle is also visible in all data sets, as it is in the Northern Hemisphere. Again, this cycle is more pronounced at 85 km , but is also visible at 105 km . The SCIAMACHY data above 95 km is low compared to the other three instruments, but is still within the error range.

### 4.3 Middle and Iow latitudes

Figure 5 shows the comparisons at middle latitudes $\left(32.5^{\circ} \mathrm{S}\right)$ at 85 and 105 km . The data at all altitudes are shown in Fig. B3. Similarly, the data at $2.5^{\circ} \mathrm{N}$ are compared in Fig. 6 ( 85 and 105 km ) and in the same Appendix in Fig. B4 (all altitudes).

In general, the magnitude of the NO number density in these regions is smaller than at higher latitudes by a factor of 5 to 10, in particular at polar winter, as also discussed in Bermejo-Pantaleón et al. (2011). Above 95 km, the SCIAMACHY measurements are low compared to the other instruments but they are well within the variations of the data.

At altitudes above 95 km , the number densities follow the solar cycle activity. They decline at the beginning and increase again at the end of the investigated period. A distinct annual cycle of the NO density is not clearly identifiable at these latitudes. This is in contrast to the time series at higher latitudes and the result of different production mechanisms depending on latitude. It is already known that at high latitudes and under auroral conditions, the production of NO is larger than that at equatorial latitudes. The differences between the polar summer and polar winter regions come from larger photochemical losses in the summer, leading to lower NO densities (Bermejo-Pantaleón et al., 2011).

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The time series from all instruments are consistent in all regions we compared, in particular considering the sometimes large error bars (equal to the $95 \%$ confidence interval of the daily zonal mean). Larger differences indicate short-term variations that are measured by one instrument on a particular day when none of the other instruments observed NO in the upper atmosphere.

## 5 Vertical profiles

We obtain the most direct comparison of the individual results by comparing the vertical density profiles. Discrepancies in the profiles provide insight into the characteristic strengths and weaknesses of the individual instruments and their NO retrieval. However, we focus here on daily zonal mean data and apart from SCIAMACHY and MIPAS on the same satellite, the local solar times of the measurements differ substantially. In addition, the different distributions of the individual measurement or tangent points make comparing vertical profiles difficult.

To obtain reliable statistics about the profiles, we first calculate the difference profile for each coincident $5^{\circ}$ latitude bin of two instruments, that is on days when both instruments provide data in the same bin. These differences are calculated on a common altitude grid, the data are interpolated if necessary. We then take the median of all these difference profiles in three geographical regions, $90-50^{\circ} \mathrm{S}, 50^{\circ} \mathrm{S}-50^{\circ} \mathrm{N}$, and $50-90^{\circ} \mathrm{N}$.

### 5.1 SCIAMACHY

First we compare the SCIAMACHY data to the other instruments because it provides the most regular data throughout the altitude and latitude range (see Figs. 1 and 2). Figure 7 shows the median profile of the relative number density differences between the SCIAMACHY measurements ( $n_{\mathrm{NO}}^{\mathrm{SCIAMACHY}}$ ) and the other instruments ( $\left.n_{\mathrm{NO}}^{\text {other }}\right)$, averaged over days coincident with other observations. Shown are the results in the

southern polar region ( $90-50^{\circ} \mathrm{S}$ ), at middle and low latitudes outside the polar regions $\left(50^{\circ} \mathrm{S}-50^{\circ} \mathrm{N}\right)$, and in the northern polar region ( $50-90^{\circ} \mathrm{N}$ ).

MIPAS and SCIAMACHY share the same satellite and therefore they performed the most congruent measurements. In addition, their limb scans were scheduled to measure the mesosphere and lower thermosphere during the same orbits once a month. Therefore, we get the best statistics from this pair of instruments.

SCIAMACHY has fewer coincident days with SMR than with MIPAS. We also have to consider the different local times of the measurements, which are 10:00 for SCIAMACHY and between 06:00 and 07:00 for SMR (at the equator). This timing of the SMR measurements makes them susceptible to the NO diurnal cycle and may lead to systematic differences in the measured number densities, in particular in the lower mesosphere. The coincidences with ACE-FTS amount to only about ten to twenty usable profiles. Since ACE-FTS measures at sunrise and sunset primarily at higher latitudes, the NO diurnal cycle also affects the retrieved number densities.

The patterns of the MIPAS to SCIAMACHY differences in Fig. 7 are similar in all three regions. The magnitude of the difference varies slightly with the region, it is smallest at middle and low latitudes and at high northern latitudes. Between 120 and 140 km , the MIPAS and SCIAMACHY number densities agree well in all three regions, the SCIAMACHY NO densities differ only slightly from the MIPAS measured densities. In the main production region from 100 to 120 km , SCIAMACHY significantly underestimates the NO number densities compared to MIPAS measurements by about $50 \%$ at high southern latitudes and by about $40 \%$ at middle and low latitudes and at high northern latitudes. Between 70 and 95 km in the northern polar region, the SCIAMACHY NO number densities are larger than the measurements from ACE-FTS and MIPAS. The SCIAMACHY data are consistent with the other instruments at these altitudes in the southern polar region and at middle and low latitudes.

SCIAMACHY consistently measures lower number densities than SMR by 10 to $20 \%$. The data agree from 80 to 100 km in the northern polar region. At middle and low latitudes, the SCIAMACHY and SMR densities agree at 95 and at 110 km .

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The SCIAMACHY data agree well with the ACE-FTS data in the Northern Hemisphere. In the Southern Hemisphere, the SCIAMACHY measurements are smaller than the ACE-FTS number densities between 90 and 105 km . Both number densities are consistent below 90 km considering the statistical error. At middle and low latitudes

### 5.2 Other instruments

Figure 8 shows the median of the relative differences comparing MIPAS data to the other instruments. In all three regions, MIPAS measures the largest NO number densities between 100 and 120 km , between 80 and $120 \%$ larger than measured by the other instruments. As seen in the previous section, MIPAS and SCIAMACHY data agree in all three latitude regions at altitudes between 120 and 140 km . In the northern polar region and at middle and low latitudes, both results agree also between 80 and 95 km . In the southern polar region, the MIPAS measurements are about 50 to $100 \%$ larger than the SCIAMACHY measurements from 70 to 95 km .

Between 80 and 100 km in the southern polar region, the MIPAS and the ACE-FTS data agree well. Below 80 and above 100 km in this region, the MIPAS number densities are larger by 50 to $100 \%$. At middle and low latitudes MIPAS and ACE-FTS have only a few coincident measurement days and even less comparable data points when considering the instruments' sensitivity. In the upper usable altitude region, between 95 and 105 km , MIPAS and ACE-FTS are consistent. In the lower altitude region from 65 to 70 km , MIPAS number densities are larger by 50 to $80 \%$. In the northern polar region, the MIPAS number densities are also larger than the ACE-FTS measurements in the same altitude region. They are smaller than the ACE-FTS number densities between 85 and 100 km .

Compared to SMR, the MIPAS NO number densities are significantly smaller in the southern polar region from 80 to 100 km by about 40 to $50 \%$. In the northern polar region from 80 to 100 km , the MIPAS number densities are about 30 to $50 \%$ smaller than the SMR data. Above and below, MIPAS and SMR agree within the statistical

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error. Both number densities agree well at middle and low latitudes between 90 and 100 km .

Figure 9 shows the median of the relative profile differences comparing the SMR data to the other measurements. In all three regions, the SMR number densities are consistently larger than the SCIAMACHY data above 100 km . We observe the largest differences in the southern polar region, up to $80 \%$. However, both number densities agree in the northern polar region from 80 to 100 km . At middle and low latitudes, the difference between SMR and SCIAMACHY vary between 10 and $40 \%$, reaching larger but insignificant values between 85 and 90 km .

Compared to MIPAS, the SMR number densities are significantly smaller in all three regions from 100 to 120 km by 20 to $50 \%$. From 85 to 100 km , the SMR number densities are consistently larger than the MIPAS data, in the southern polar region between 50 and $80 \%$, and in the northern polar region up to $120 \%$ but with a large uncertainty. At middle and low latitudes, both measurements are consistent between 90 and 105 km .

Compared to ACE-FTS, the SMR number densities are substantially larger at high southern latitudes above 100 km , differing by $100 \%$ at 105 km . ACE-FTS and SMR data agree well between 80 and 100 km . Below 80 km in that region, the SMR number densities differ from the ACE-FTS results between 50 and $80 \%$ but with a large statistical uncertainty. Both NO number densities are comparable at middle and low latitudes over the whole altitude range considering the statistics. The differences in the northern polar region behave similarly to the results at high southern latitudes, the maximum deviation is $150 \%$ at 105 km . Here, both data sets agree well between 90 and 100 km . The SMR number densities are up to $60 \%$ larger than the ACE-FTS data below 90 km in that region.

Figure 10 shows the median of the relative profile differences of the ACE-FTS NO number densities to the other instruments. It summarises the above discussion from ACE-FTS' point of view, showing the best agreement with MIPAS and SMR in the southern polar region from 85 to 100 km , and with SCIAMACHY and SMR in the north-

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ern polar region from 80 to 100 km . The ACE-FTS number densities are substantially larger than the SCIAMACHY measurements in the southern polar region between 90 and 105 km . They also differ significantly from the MIPAS data in the northern polar region between 85 and 100 km . At middle and low latitudes, the ACE-FTS measurements are smaller than all other number densities from 70 to 100 km by 40 to $100 \%$.

## 6 Multi-linear regression analysis

The direct quantitative comparison of the NO data of the four instruments is difficult for several reasons. Coincident measurements are sparse and the local times of the individual measurements differ substantially between some of the instruments. The solar UV radiation influences the NO density annually because of the different solar inclination, and inter-annually due to its varying intensity during the eleven-year solar cycle. Then, the NO density can vary substantially from day to day depending on particle precipitation rates, for example at times of high geomagnetic activity.

All instruments scanned the MLT region only on particular single days, for example SCIAMACHY data are only available for the later part of the time period. This makes capturing all variations of NO in the upper atmosphere difficult. To overcome these shortcomings, we carry out a multi-linear regression analysis of the zonal mean data.

We construct a simple transfer function for the NO number density $n_{\mathrm{NO}}$ for our multi${ }_{20}$ instrument intercomparison on non-coincident days. This function includes an offset, annual and semi-annual harmonic terms, and accounts for the solar cycle activity proportional to the UV Lyman- $\alpha$ flux. It also includes a term for daily NO variations caused by geomagnetic activity using the Kp index.

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The NO number density $n_{\mathrm{NO}}$ at altitude $z$, latitude $\phi$, and time $t$ is described by:

$$
\begin{align*}
n_{\mathrm{NO}}^{\text {model }}(\phi, z, t)= & a(\phi, z)+b(\phi, z) \cdot \operatorname{Ly} \alpha(t)+c(\phi, z) \cdot \mathrm{Kp}(t) \\
& +\sum_{n=1}^{2}\left[d_{n}(\phi, z) \cos (n \omega t)+e_{n}(\phi, z) \sin (n \omega t)\right] \tag{1}
\end{align*}
$$

In the regression model Eq. (1) $a$ is the constant offset, $d_{1}$ and $e_{1}$ are the annual, between 0 and $6 \times 10^{8} \mathrm{~cm}^{-3}$. At higher altitudes the residuals vary between
$+2 \times 10^{8} \mathrm{~cm}^{-3}$ with the number densities varying between 0 and $8 \times 10^{8} \mathrm{~cm}^{-3}$.

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Figure 15 shows the relative mean residuals (RMR) of the individual measurements to the composite fit. The RMR are defined as

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RMR := $\frac{\sum_{i}\left(n_{\mathrm{NO}}^{\text {meas }}\left(t_{i}\right)-n_{\mathrm{NO}}^{\text {model }}\left(t_{i}\right)\right)}{\sum_{i} n_{\mathrm{NO}}^{\text {model }}\left(t_{i}\right)}$,
where $n_{\mathrm{NO}}^{\text {meas }}$ and $n_{\mathrm{NO}}^{\text {model }}$ are the measured and modelled NO number densities. Here, 5 modelled data means applying the coefficients from the composite fit to the days $t_{i}$ of the individual measurements. The marked boxes in Fig. 15 indicate $95 \%$ significance determined by the $F$ test (Brook and Arnold, 1985; Neter et al., 1996). The ACE-FTS number densities are larger than the composite by about 10 to $30 \%$ where the fit is significant. The same is true for the MIPAS data, except between 70 and 105 km at middle to high northern latitudes. There, the MIPAS number densities are smaller by about the same amount. The SCIAMACHY number densities are smaller than the composite fit by about 10 to $30 \%$ above 90 km throughout the latitude range. They are slightly larger below 90 km , and substantially so between 75 and 85 km at $77.5^{\circ} \mathrm{S}$. The SMR number densities are larger than the composite fit by about 10 to $20 \%$ almost everywhere, except at the highest altitudes at middle and low latitudes.

We next analyse the coefficients of the individual regressions to compare the instruments' responses to the harmonic cycles and in particular to Lyman- $\alpha$ and Kp changes. Figures 16 and 17 show the cosine coefficients $d_{1}$ and $d_{2}$ of the annual and semiannual cycle. Values with a larger than $95 \%$ significance, calculated from $t$ statistics, are marked with crosses. We find that these coefficients have about the same values for all instruments. The MIPAS, SCIAMACHY, and SMR $d_{1}$ coefficients agree remarkably well. The SMR coefficients take high values below 80 km , but these values are not significant. The ACE-FTS coefficients, however, differ substantially at latitudes larger than $60^{\circ}$. The $d_{2}$ coefficients agree not as strongly, but they are also one order of mag-

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nitude smaller in general. The sine coefficients $e_{1}$ and $e_{2}$ are not shown, but they agree on a similar level with a slightly smaller magnitude than their cosine counterparts.

Figures 18 and 19 show the coefficients $b$ and $c$ of the Lyman- $\alpha$ UV index and the Kp geomagnetic index. The latitude-altitude patterns of both coefficients are similar for all instruments. The SMR and SCIAMACHY Lyman- $\alpha$ coefficients $b$ are enhanced at all latitudes in a band from 95 to 110 and 115 km . This band of larger values is less pronounced in the ACE-FTS and MIPAS data. The coefficients are also consistently enhanced at polar latitudes from 80 to 100 km in the ACE-FTS, MIPAS, and SMR data. The coefficients derived from SCIAMACHY data are increased only in the northern polar region.

We observe enhanced Kp coefficients $c$ in the main production regions north and south with all instruments. The magnitude of these coefficients is smallest in the SCIAMACHY data because polar night measurements were rare and only performed during a period of low solar activity. ACE-FTS also provides only few data points at the beginning and the end of the polar night, see Figs. 1 and 2. But in contrast to SCIAMACHY, ACE-FTS measured also during a period of higher solar activity which explains the larger Kp coefficients. The patterns, however, are consistent with the data from the other instruments.

## 7 Conclusions

In this study, we compared the measurements from four instruments, three limb sounders and one solar occultation instrument, using different spectral ranges: infrared, sub-millimetre waves, and ultra-violet. Despite these different methods and accompanying different retrieval strategies, the nitric oxide daily zonal mean densities of all four instruments are consistent during the comparison time period. However, the instruments have different altitude resolutions. But, as single selected profiles show, the structures of the NO density in the MLT region are larger than the instruments' altitude

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resolutions. Therefore, even though these resolutions differ, this comparison is feasible without applying additional corrections to the data, such as averaging kernels.

We find that the NO number density time series agree well, almost all data points lie within the error bars (equal to the $95 \%$ confidence interval) of the daily zonal mean 5 values. The remaining differences can be attributed to the different MLT measurement schedules and latitude-time coverage of the instruments. For example, SCIAMACHY provides only day-time measurements and therefore less data at high latitudes at polar winter, where the other instruments observe enhanced NO values. This biases the SCIAMACHY daily zonal mean NO number densities to lower values compared to the

The medians of vertical profile differences in three geographic regions ( $90-50^{\circ} \mathrm{S}$, $50^{\circ} \mathrm{S}-50^{\circ} \mathrm{N}$, and $50-90^{\circ} \mathrm{N}$ ) show that the NO number densities agree in general within $\pm 50 \%$. However, we found larger differences up to $100 \%$ in regions where the coincident statistics are very low, see, for example, Fig. 10. We also observed that MIPAS measured substantially larger densities than all other instruments between 100 and 120 km . But we identified no other systematic feature between the four instruments.
We set up a simple transfer function for the NO number density $n_{\mathrm{NO}}$ on noncoincident days in Sect. 6. This multi-linear regression model Eq. (1) describes all time series well, accounting for annual and semi-annual variations and for long-term solar and short-term geomagnetic variations using the Lyman- $\alpha$ UV index and the Kp index. We should note that this method involves jointly fitting data sets with different local times, geolocations, and temporal frequencies. It may therefore introduce a bias in the final result of the fit. However, our aim is not to provide a model for NO in the MLT, but to use this regression fit as a transfer function to check the consistency of temporally mismatched observations.

We also performed the regression analysis on an individual instrument basis and obtained consistent coefficients in the important altitude region. In particular, we found consistent responses to the estimators related to solar and geomagnetic variability.

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These estimators can be further used as an empirical model of NO in the middle atmosphere, in particular useful for climate models.

NO is an important proxy for the influence of solar activity on Earth's atmosphere. The mesosphere and lower thermosphere is the main production region of NO from solar particles. Continuous measurements of NO in this region are rare and with the loss of the Envisat satellite in April 2012, two important instruments are missing. With the end of SCIAMACHY and MIPAS measurements, only ACE-FTS and SMR provide continuous NO measurements between 80 and 120 km . From these two, only SMR delivers global data. ACE-FTS, however, offers only limited latitudinal coverage since
10 it observes solar occultations, scanning the atmosphere at sunrise and sunset. There are still more satellite instruments measuring NO in the middle atmosphere: OSIRIS, SABER, and SOFIE. However, OSIRIS measures only between 85 and 100 km and SABER only above 100 km . SOFIE is another solar occultation instrument and therefore covers also only a limited latitude range, similar to ACE-FTS.
${ }_{15}$ Acknowledgements. S. Bender and M. Sinnhuber thank the Helmholtz-society for funding this project under the grant number VH-NG-624. The IAA team (M. López-Puertas and B. Funke) was supported by the Spanish MINECO under grant AYA2011-23552 and EC FEDER funds. The SCIAMACHY project was funded by German Aerospace (DLR), the Dutch Space Agency, SNO, and the Belgium ministry. ESA funded the Envisat project. The University of Bremen
${ }_{20}$ as Principal Investigator has led the scientific support and development of SCIAMACHY and the scientific exploitation of its data products. The Atmospheric Chemistry Experiment (ACE), also known as SCISAT, is a Canadian-led mission mainly supported by the Canadian Space Agency and the Natural Sciences and Engineering Research Council of Canada. Odin is a Swedish-led satellite project funded jointly by Sweden (SNSB), Canada (CSA),
25 Finland (TEKES), France (CNES) and the third party mission program of the European Space Agency (ESA). The provision of MIPAS level-1b data by ESA is gratefully acknowledged. We acknowledge support by Deutsche Forschungsgemeinschaft and Open Access Publishing Fund of Karlsruhe Institute of Technology.
${ }_{30}$ The service charges for this open access publication

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Table 1. Instrument overview.

|  | SCISAT | Odin |  | Envisat |
| :--- | :--- | :--- | :--- | :---: |
|  | ACE-FTS | SMR | MIPAS | SCIAMACHY |
| orbit |  | sun-synchronous |  | sun-synchronous |
| altitude | 650 km | 580 km | 790 km |  |
| equatorial crossing time | variable | $06: 00-07: 00 / 18: 00-19: 00$ | 10:00/22:00 |  |
| measurement type | solar occultation | limb | limb | limb |
| NO wavelength | infrared | sub-mm | infrared | ultra-violet |
| MLT measurement days | 1941 | 301 | 199 | 78 |
| MLT time period | $2004-2010$ | $2003-$ present | $2005-2012$ | $2008-2012$ |
| data version | 3.0 | 2.1 | V40 (2005-2009) | 2.0 |
|  |  |  | V5R (2010-2012) |  |

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Figure 2. NO zonal mean time series at 105 km , ACE-FTS, MIPAS, SCIAMACHY, and SMR (from top to bottom).

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Figure 6. NO time series comparison as in Fig. 3 for $2.5^{\circ} \mathrm{N}$.

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SCIAMACHY


Figure 7. NO vertical profile comparison of the SCIAMACHY NO number density ( $n_{\mathrm{NO}}^{\mathrm{SCIAMACHY}}$ ) to the other instruments ( $\left.n_{\mathrm{NO}}^{\text {other }}\right)$. Shown is the median of the relative differences ( $n_{\mathrm{NO}}^{\mathrm{SClAMACHY}}-$ $\left.n_{\mathrm{NO}}^{\text {other }}\right) / n_{\mathrm{NO}}^{\text {other }}$ averaged over days coincident with other observations. The panels show the results in the southern polar region (left), at middle and low latitudes (middle), and in the northern polar region (right). The error bars indicate the $95 \%$ confidence interval using the median absolute deviation ( $3 \times \mathrm{MAD}$ ). The numbers in parentheses denote the number of coincident profiles in these three regions from south (left) to north (right).

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Figure 8. NO vertical profile comparison of MIPAS NO data $n_{N O}^{\text {MIPAS }}$. For a detailed description see Fig. 7.

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Figure 9. NO vertical profile comparison of the SMR data $n_{\mathrm{NO}}^{\mathrm{SMR}}$. See Fig. 7 for details.

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Figure 10. NO vertical profile comparison of the ACE-FTS data $n_{N O}^{A C E-F T S}$. See Fig. 7 for details.

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Figure 11. NO time series regression results at $67.5^{\circ} \mathrm{N}, 85 \mathrm{~km}$. The upper panel shows the individual time series and the composite regression fit using all data simultaneously. The lower panel shows the residuals.

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Figure 13. NO time series regression results at $67.5^{\circ} \mathrm{S}, 85 \mathrm{~km}$ as in Fig. 11.

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Figure 15. NO mean residuals of the individual measurements to the composite fit. The marked boxes indicate a larger than $95 \%$ significance determined using the $F$ test of the regression fit.

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Figure 17. NO regression coefficient $d_{2}$ of the semi-annual cycle (cosine part) as in Fig. 16.


Figure 18. NO regression coefficient $b$ of the long-term solar UV variations from the Lyman- $\alpha$ index as in Fig. 16.

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Figure 19. NO regression coefficient $c$ of the short-term geomagnetic variations from the Kp index as in Fig. 16.


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Figure C4. NO time series regression results at $2.5^{\circ} \mathrm{N}$, for $75,85,95,105$, and 115 km (from top to bottom).

