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Abstract

The Gimballed Limb Observer for Radiance Imaging of the Atmosphere (GLORIA) is an imaging limb emission sounder operating in the thermal infrared region. It is designed to provide measurements of the Upper Troposphere/Lower Stratosphere with high spatial and high spectral resolution. The instrument consists of an imaging Fourier transform spectrometer integrated in a gimbal. The assembly can be mounted in the belly pod of the German high altitude and long range research aircraft HALO and in instrument bays of the Russian M55 Geophysica. Measurements are made predominantly in two distinct modes: the chemistry mode emphasises chemical analysis with high spectral resolution, the dynamics mode focuses on dynamical processes of the atmosphere with very high spatial resolution. In addition the instrument allows tomographic analyses of air volumes. The first measurement campaigns have shown compliance with key performance and operational requirements.

1 Introduction

The Upper Troposphere/Lower Stratosphere (UTLS) region plays a crucial role in the climate system. Surface temperature is sensitive to changes in the dynamic structure and composition of this region (Solomon et al., 2010; Riese et al., 2012). The goal of the Gimballed Limb Observer for Radiance Imaging of the Atmosphere (GLORIA) project is to achieve substantially improved spatial resolution and spatial coverage in observing the UTLS by airborne remote sensing. Scientific applications for the improved spatial performance range from questions of atmospheric dynamics to the understanding of chemical fine structures (Riese et al., 2014). GLORIA also serves as a demonstrator for the potential of this technique for observations from space.

Limb emission sounding in the thermal infrared region is a mature technique that has been used from airborne and spaceborne platforms for earth observation since more than three decades. Radiometers like LIMS (Gille and Russell, 1984) and HIRDLS

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(Gille and Barnett, 1996), Fourier transform spectrometers (FTS) like SIRIS, MIPAS and TES (Kunde et al., 1987; Oelhaf et al., 1994; Beer et al., 2001; Fischer et al., 2008), and dispersive spectrometers like CRISTA (Offermann et al., 1999) have been employed. The majority of these instruments scans the atmospheric limb in vertical direction with the aid of a mirror system. TES is designed to observe a part of the atmospheric limb simultaneously with a linear detector array with 16 elements. Larger linear and two-dimensional infrared detector arrays have become available in the last two decades and were used in a nadir sounding operational grating spectrometer (Aumann et al., 2003) and in a prototype imaging FTS nadir sounder for atmospheric research (Elwell et al., 2006). Development of an operational imaging FTS nadir sounder has commenced (Rodriguez et al., 2009). In parallel, commercial imaging FTS for military and security applications (Vallieres et al., 2005; Sabbah et al., 2012) based on large focal plane arrays have been realised in recent years. However none of these instruments has the spectral coverage, sensitivity and spectral resolution that are needed for an imaging limb sounder for atmospheric research. Accordingly a development of an instrument optimised for the requirements of limb sounding from an airborne platform had to be undertaken.

Section 2 discusses the requirements and Sect. 3 outlines the main design trade-offs. A condensed description of the instrument design follows. Section 4 shows essential characterisation results and discusses to what extent the most important requirements can be achieved.

2 Instrument requirements

It is challenging to get a comprehensive picture of the UTLS due to the complex chemistry with a large number of involved species and the dynamical environment that is characterised by small structures. Filaments in the UTLS, while wide-stretched in one horizontal direction, typically only extend on the order of 100 km in the other horizontal direction and a few hundred meters vertically (Bacmeister et al., 1999; Pan et al., 2009;

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Weigel et al., 2012; Barre et al., 2012; Ungermann et al., 2013). Therefore, GLORIA was designed for two main measurement tasks: the first one, called chemistry mode (CM), focuses on high spectral resolution and thus maximises the number of retrieved trace species with sufficient spatial sampling for the aforementioned structures. The second one, called dynamics mode (DM), has lower spectral resolution but provides further increased horizontal resolution and the ability for three dimensional sounding for a limited number of species. The geometric and spectral requirements of the two modes are summarised in Tables 1 and 2. Future scientific applications request instrument flexibility allowing other compromises between spectral and spatial resolution and the ability for nadir viewing.

The first group of requirements (Table 1) deals with observation geometry, observation coverage and spatial and temporal sampling of the measurements. The nature of the requirements dictated the principal design decisions for the instrument. The quest for high vertical resolution directly leads to the concept of a limb sounder. The demanding requirements concerning simultaneous vertical coverage and on temporal sampling compel to use an imaging optical system with very high throughput. The goal to perform tomographic measurements could only be achieved with angular agility of the instrument in yaw direction. The use of a research aircraft as carrier with its inevitable pitch, roll and yaw movements during flight made a three axes stabilisation of the field of view (FOV) during measurement recording indispensable.

The specific values for the above mentioned requirements (Table 1) are motivated by the following considerations: scientific interest in the free troposphere requires to set the value for the lower boundary of the vertical coverage to approximately 4 km altitude. Viewing by up to 0.8° above the horizon provides some information on trace gas columns amounts above the aircraft (Woiwode et al., 2012).

Infrared limb sounding is capable of providing vertical resolutions in the 300 to 400 m range (Ungermann, 2011; Ungermann et al., 2011). Consequently, the vertical sampling should be in the few hundred meter range at the tangent point to be compliant with this resolution. The retrieval study of Ungermann et al. (2011) furthermore indicates

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that the horizontal FOV size should not be much larger than 10 km at a tangent altitude of 10 km to be compatible with the tomographic retrieval approach. In addition, finer sampling in horizontal direction should be possible for cloud identification and analysis.

The realised values of 1.9 arcmin vertical and horizontal pixel size reflect the properties of the detector array and the chosen optical system. The typical horizontal sample size of 1.5° mentioned in Table 1 represents the integration over a row of 48 pixels. These values correspond to a sample size of 140 m vertically and about 6.7 km horizontally at a tangent altitude of 10 km seen from an observer altitude of 15 km.

The study also showed that the yaw angle should be adjustable between 45° and 135°.

The second group of requirements (Table 2) deals with the spectrometer itself: the goal of covering at least emission bands of carbon dioxide, ozone, water vapour, nitric acid, some chlorofluorocarbons, nitrous oxide, chlorine nitrate and methane in the thermal infrared region leads to a broad spectral coverage requirement of the instrument. Chlorine nitrate determines the spectral cut-on to approximately 780 cm⁻¹, water vapour and methane force to measure up to 1400 cm⁻¹. The spectral sampling requirement of 0.0625 cm⁻¹ for the chemistry mode is based on experience with the MIPAS instruments (von Clarmann et al., 2009; Friedl-Vallon et al., 2004), the requirement of 0.625 cm⁻¹ for the dynamics mode is based on experience with the CRISTA instrument (Riese et al., 1997, 1999) and on studies performed for a satellite borne imaging FTS (ESA, 2012). The sensitivity requirements of target (threshold) 5 nW(cm² srcm⁻¹)⁻¹ (10 nW(cm² srcm⁻¹)⁻¹) for dynamics mode and 15 nW(cm² srcm⁻¹)⁻¹ (30 nW(cm² srcm⁻¹)⁻¹) for the chemistry mode are derived from the comparable performances of the precursor instruments. Experience with these precursor instruments also has proven that the radiometric and spectral accuracy requirements of 1 % target/2 % threshold and 10 ppm respectively are necessary and sufficient.

A third group of requirements, not mentioned in the tables, is of a technical nature. The instrument needs to operate on aircraft that are able to reach the lower

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stratosphere. GLORIA is primarily designed for use on the German high altitude and long range research aircraft HALO, which is based on a commercial Gulfstream G550 business jet (see Fig. 1). However, the possible use of the instrument on other aircraft, notably on the Russian M55 Geophysica research aircraft, has to be provided for. Accordingly environmental requirements for temperature inside the instrument compartment down to -50°C and for ambient pressure down to 70 hPa have to be taken into account. The instrument will be exposed to vibrations excited by the aircraft turbines and by aerodynamics. Precise quantification of excitation levels is difficult before flight, since the instrument itself, its mounting position and the aerodynamic fairing influence the vibration environment. Finally, the instrument has to be certified for operation on the aircraft. In particular the certification for the HALO is demanding, since HALO in the long term shall be able to operate as civil aircraft without flight restrictions imposed by the instrumentation.

3 Instrument concept

3.1 Overview

The GLORIA spectrometer is a cooled imaging Fourier transform spectrometer (IFTS) with a large cryogenic HgCdTe detector array for detection of the infrared (IR) radiation. The spectrometer is mounted in a gimbal that allows agility in pitch, roll and yaw direction. The instrument is designed to be mounted outside of the aircraft cabin. Calibration sources are two large area blackbodies and a “deep space” view for radiometric and spectral calibration. A graphical overview of the instrument is given in Fig. 2 and a picture of the instrument mounted to the HALO aircraft is depicted in Fig. 3. Table 3 provides a summary of key properties.

An IFTS was chosen as the baseline spectrometer concept to benefit from the massive total throughput advantage of nearly two orders of magnitude that this kind of instrument offers in comparison to a conventional FTS. This advantage is essential

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in achieving the required spatial and temporal sampling. Any scanning system, FTS or grating, would need an entrance aperture impossible to realise on an aircraft to achieve the same throughput. Further arguments for the imaging FTS are its flexibility in spectral and spatial resolution and the non-availability of infrared detector arrays that support a high spectral resolution dispersive system.

Although state of the art IR detector systems still suffer from significant electronic noise, photon noise in our case substantially contributes to the noise performance. Furthermore high photon background limits integration time since the full well capacity of the integration capacitors in the read-out integrated circuit is finite. This limitation has direct impact on the sensitivity. Thus a reduction of the optics temperature from room temperature to 220 K leads to a signal to noise advantage of roughly a factor of three for a spectrometer using the infrared detector arrays available to the project (see Sect. 3.2). This fact led us to the design of an actively cooled spectrometer.

As mentioned in the requirement section, stabilisation of the FOV in altitude and orientation is essential. This can either be achieved by mounting the spectrometer in a gimbal or by a fore-optics with a set of controlled scanning mirrors. Two aspects led to the adoption of a gimbal mount: an imaging system has an extended FOV, accordingly not only elevation and azimuth but also image rotation needs to be stabilised. The combination of image rotation device and azimuth and elevation scanner would have led to a complex optical set-up with many optical elements, long optical path and subsequently large optical elements. It is noted that a downside of the decision for the gimbal is the very strict limitation on size and mass of the spectrometer.

On HALO the whole instrument is mounted outside the aircraft in the unpressurised belly pod to avoid a large IR window in the pressurised aircraft cabin. Such a window would have been outside the calibration path and it would have caused severe flight certification problems. On the M55 Geophysica there is no choice: all instruments have to be mounted in unpressurised compartments. An insulating window is implemented in the spectrometer enclosure to avoid water contamination and reduce aeroacoustic excitation within the spectrometer.

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The agility of the gimballed yaw frame allows for inflight calibration by turning the spectrometer towards the apertures of the two large area blackbody radiation sources. Thus, the whole optical chain including the spectrometer window is in the calibration path. Due to the high spectral resolution of the spectrometer, a posteriori spectral calibration using atmospheric lines is possible.

3.2 Spectrometer

The GLORIA spectrometer needs to be as compact and rigid as possible. Size and mass are limited within the gimballed frame, the cooled volume must be kept small. The vibration environment of an aircraft demands a sturdy design with as few optical elements as possible. The optical set-up of the interferometer is straightforward and simple (see Fig. 4). The radiation enters through an insulating double germanium window with wedged plates. Amplitude splitting is performed by a dielectric beam splitter on a potassium chloride substrate. Two Zerodur cube corners – one fixed and one linearly moveable – act as retro reflectors. After recombination the radiation is focused on the focal plane array with an achromatic infrared doublet (Friedl-Vallon et al., 2014). The interferometer needs to be enclosed since it is cooled and has hygroscopic components in the infrared optics. The insulating window solution is compact and allows simple testing of the cooled instrument on ground under ambient conditions.

Since the task of imaging the limb scene on the detector can be solved without a telescope in front of the interferometer, a telescope was omitted to save complexity, volume and mass. A transfer optics between interferometer and detector was avoided to reduce the number of optical components and the spatial envelope. The downside of a larger beam diameter in the interferometer was an acceptable trade-off given the benefits. A linear interferometer set-up was chosen since it was considered as the most compact and mechanically robust solution. The spectral resolution requirement does not impose a movement of both retro reflectors. The moving cube corner is mounted on a dove tail slide that is driven by a lead screw with a direct drive motor controlled by an angular high resolution optical shaft encoder. The approach leads to a system

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with high resonance frequencies (Piesch et al., 2014). However the slide construction requires regular servicing. The interferometer provides a maximum optical path difference (OPD) of up to 9 cm. Two-sided interferograms with 0.8 cm or with 8 cm maximum OPD are recorded in dynamics and chemistry mode, respectively.

5 The quadrature reference system for the optical path determination is based on a laser diode in the visible spectral domain whose temperature is controlled with a thermoelectric element. The spectral stability of the diode is enhanced by use of volume holographic gratings. The design was chosen for its compactness and its ability to operate within the cooled spectrometer at low temperatures.

10 The infrared radiation is focused on the detector array with an air-spaced aspheric Ge-ZnSe doublet with a focal length of approximately 72.2 mm at an operating temperature of 218 K. Since the total thermal operating range of the instrument is too large for an athermal design, a motorised thread allows focusing the optics according to operation temperature.

15 The infrared detector is an HgCdTe large focal plane array (LFPA) from the German supplier AIM GmbH and has 256×256 photovoltaic detector elements with a pitch of $40 \mu\text{m}$. The read out integrated circuit (ROIC) is operated in stare-while-scan mode. It is read out with a 10 MHz clock and 8 taps and allows selection of sub frames. Typical frame rates are in the 3–6 kHz range, depending on the size of the chosen sub
20 frame. The long-wave cut-off of the HgCdTe material of roughly 780 cm^{-1} , depending on operating temperature. The detector operation temperature of 50 K is maintained with a Stirling cooler.

25 The interferometer itself is actively cooled. Liquid carbon dioxide is sprayed into the cooler tank and carbon dioxide snow evolves. The carbon dioxide snow keeps the instrument at its operating temperature range of 200 K to 220 K for approximately 24 h (Piesch et al., 2014). The operating temperature has been chosen as an optimal compromise between technical complexity and improved noise performance. A lower operating temperature would have required evacuation of the instrument leading to increased mass and volume. Furthermore it would not have led to an additional

significant increase of signal to noise performance, since full integration time efficiency of the detection system can be reached at the chosen operating temperature.

3.3 Calibration

The radiometric calibration system consists of two identical large area high-precision blackbody radiation sources, which are independently stabilised at two different temperatures with the help of thermo-electric coolers (Olschewski et al., 2013). The achievable temperatures of the blackbodies depend on the ambient temperature and pressure with the colder of the two reaching at least 10 K below ambient. The blackbodies are cavities with aluminium pyramids on the baseplate. The emissivity of the blackbodies has been required to be better than 0.997 and the temperature uncertainty to be better than 0.1 K in order to reach the radiometric calibration uncertainty target of 1 % (Monte et al., 2014). In addition, the gimballed frame allows for directing the centre of the FOV to 10° upwards at the 90° position of the yaw gimbal. Scene spectra taken at this position can be used as proxy for deep space spectra. They serve as complementary calibration input for direct offset estimation after removal of residual atmospheric emission lines.

Spectral calibration is achieved by the analysis of the spectral position of selected CO₂ lines in the range from 940 to 972 cm⁻¹. The scene spectra of the upward measurements are best suited for this purpose, because they are usually free of clouds over the whole FOV.

3.4 Gimballed frame and pointing

As scientifically requested, the gimbal mount provides agility in all three axes. For all gimbal angles set to 0° the gimbal coordinate system is congruent with the aircraft system. When looking across track, the gimbal pitch frame compensates aircraft roll movements. The gimbal yaw level is free to move from -20° to 140° relative to the aircraft longitudinal axis. Thus it allows view of the atmospheric limb in a range from

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45° to 135° and access to the two blackbodies that are mounted at gimbal yaw angles of about -10° and 20° . Depending on the carrier aircraft and mounting position of the instrument, angular views above 128° may be obstructed by the aircraft wing and accordingly will not be used during measurements. The gimbal roll level allows compensation for aircraft movements that would otherwise rotate the FOV between 0° and 10° . The gimbal pitch level is agile from 14° for “deep space” calibration measurements down to -98° for nadir observation.

The pointing accuracy requirements can only be fulfilled with high precision sensors. Measurements from an inertial measurement unit (IMU) with laser gyroscopes on the gimbal yaw frame and a 3-axes microelectromechanical (MEMS) gyroscope on the pitch frame as well as information from a GPS and inductive angular encoders are fused to provide both fast attitude changes and stabilisation of the FOV. The navigation and gimbal control system was developed in collaboration with iMAR Navigation GmbH, Germany. All gimbal levels are powered with direct drive motors and independent drive controller units (Maucher et al., 2014). An additional camera operating in the visible spectral domain is mounted on the interferometer. This camera covers a wide FOV thereby completely enclosing the FOV of the spectrometer. For most of the interferograms correlated images of the scene or video sequences are taken. The information provided by this camera can be used for cloud identification and for pointing quality analysis.

3.5 Data acquisition concept

The general data acquisition concept is straightforward and simple, but technically demanding. All scientific raw data is stored on a fast solid state disk array in the aircraft. No attempt is made to process a significant part of the data during flight since adequate telemetry bandwidth to send the data to ground is not available on civil aircraft. Most important goals of the electronics design were to provide high bandwidth storage capabilities and to allow semi-automatic instrument control by supplying meaningful housekeeping and data quality parameters. A modular scalable design allowed for

parallel development of the different functional units and facilitates repair in case of malfunction of units (Neubert et al., 2014).

The command and control communication infrastructure for the main functional units interferometer electronics, pointing control, the different power control units, additional devices like the visual cloud camera and the central control computer is realised by Ethernet. Communication is based on a custom made slim and flexible software architecture and protocol (Kretschmer et al., 2014). On ground, operators can connect to the instrument via a wireless interface. During flight, a low bandwidth connection with an Iridium link provides limited access to command and control of the instrument. On HALO an optional in-flight operator can be directly connected to the instrument local area network (LAN). The operator can use the Iridium link for low-bandwidth text based communication with the ground crew.

The IR data is digitised with 14 bit analog-digital converters in the detector front end electronics, each IR image is tagged with a time-stamp issued from a 80 MHz clock within the interferometer data acquisition electronics and then sent to storage via the Camera-Link interface. This additional high speed link, dedicated to the transfer of the measured interferograms, guarantees continuous data throughput and avoids congestion of the command and control LAN. The reference laser zero-crossings are detected with high temporal resolution and the time-stamps of these events, issued from the same clock as for the IR images, are sent to storage via LAN.

3.6 Data processing

IR interferograms are reconstructed post flight for each pixel on a common spatially equidistant sampling grid using Brault's interpolation method (Brault, 1996). Spectral calibration parameters are derived in a pre-processing from atmospheric signatures. Apparent image distance, the position of the optical axis on the array and a linear spectral scaling factor are determined. These parameters feed a simple model that is used in the reconstruction process of the interferograms (Kleinert et al., 2014). The first processing step also includes a non-linearity correction based on a common correction

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curve for all pixels derived from laboratory characterisation of the detector module (Sha, 2013). The interferograms are then Fourier transformed into complex spectra.

Three different datasets are available for radiometric calibration: a “deep space” view with the instrument looking 10° upwards and two blackbody views with different temperature. Since the instrument optics is thermally drifting with a rate of approximately 1 K h^{-1} , calibration measurements are performed every 30 to 45 min. The “deep space” measurements are contaminated with atmospheric lines that are removed by subtraction of forward calculated spectra using climatological parameterisation and a preliminary estimate of the spectrometer gain. In the next step, the spectrometer gain is determined from the cold blackbody measurement and the offset estimate. The spectra are calibrated in the complex domain following the approach introduced by Revercomb et al. (1988). The resulting hypercubes of calibrated data are an interim product with full spatial and spectral resolution. They can be used for image analysis, e.g. cloud determination.

After application of a bad pixel mask that eliminates pixels with radiometric gain errors larger than 1.5%, the calibrated data is spatially co-added to row averages. The final data product is a column of 128 calibrated spectra per measurement and is fed into different retrieval codes (Ungermann et al., 2010; Woiwode et al., 2012; Kaufmann et al., 2014), depending on the scientific focus of the measurement.

4 Instrument performance

The first flights of GLORIA took place within the ESA Sounder Campaign (ESSenCe) on the M55 Geophysica from Kiruna in northern Sweden in December 2011. Two successful technological flights into the Arctic were performed. In late summer 2012, GLORIA was embedded in the Transport and Composition in the UT/LMS (TACTS) and Earth System Model Validation (ESMVal) campaigns. Scientific flights over Europe and the Atlantic were performed with the HALO aircraft from Oberpfaffenhofen in southern Germany. In addition, HALO circled Africa and performed an excursion towards

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Antarctica. GLORIA encountered various climatological conditions ranging from mid-latitude summer to tropics and Antarctic winter. Thirteen scientific flights with more than 100 flight hours in total were carried out. GLORIA delivered data from all but one flight. It was operated in all scientific operation modes including tomography mode.

The following section concentrates on the verification of the compliance to the overarching scientific requirements from flight data. Detailed characterisation and performance analysis on subsystem level is described in the companion articles concerning the spectrometer, pointing system and processing (Friedl-Vallon et al., 2014; Piesch et al., 2014; Maucher et al., 2014; Kleinert et al., 2014).

Figure 5 depicts an overview of the result of the dynamics mode temperature retrieval for the flight over the North Atlantic of 26 September 2012. Good quality temperature information is available from ECMWF. We used the analysis model with a horizontal resolution of T1279 and interpolated the model data linearly in space and time to the location of the measurement points, i.e. to the tangent points. The comparison of the retrieved data with this interpolated ECMWF data provides an opportunity for an overall quality check on product level.

Compliance with geometrical instrument requirements i.e. the vertical extension and the correct placement of the FOV at the horizon can be verified with this data set. Figure 6 illustrates the agreement of the retrieval result with the ECMWF field. The retrieval has been performed without using the model data for the construction of the a priori information. The retrieval for a nominal FOV of 4.08° as derived from engineering data and focal length determination during spectral calibration is compared to a retrieval where the FOV was assumed to be 1 % larger and smaller. The nominal FOV shows the smallest average difference between retrieved temperature and ECMWF field. A negative bias of approximately 0.5 K is seen over most of the altitude range. Due to the limb observing geometry, the largest sensitivity to a FOV scaling error is seen in the lower altitudes. The results indicate that an erroneous reconstruction of the retrieval grid is detectable. This inherently verifies other geometrical requirements such as sampling size in vertical and horizontal direction. The mean pointing offset for

this particular flight is determined to approximately 0.05° by minimising the variance of the deviation.

Vertical resolution of the instrument can be directly deduced from IR images showing cloud structures. Figure 7 shows an IR image of a subvisible cirrus cloud taken during a flight over the North Atlantic Ocean on 26 September 2012. It illustrates the in-flight imaging quality of the IR system. The observed cirrus layer is partly as thin as one pixel row. Vertical resolution of the retrieval can be derived from retrieval diagnostics. Figure 8 depicts the achieved vertical resolution for the trace gas HNO_3 and demonstrates the agreement with the scientific requirement. The retrieval diagnostics in chemistry mode indicates a vertical resolution of the product of 280 to 400 m. The difference in mixing ratio and vertical resolution base on the retrieval characteristics of the used codes, in particular on selection of the spectral windows and on the regularisation scheme.

Compliance with yaw agility and temporal sampling requirements for the tomography mode is illustrated in Fig. 9. The yaw and pitch movement of the gimbal during one tomography sequence is shown. The yaw angle is decremented from 128° to 44° in 4° steps. Every measurement point represents one column of 128 spectra, 42 columns are recorded in 84 s. The pitch angle variation over the measurement sequence is caused by the compensation of the pitch and roll movements of the aircraft. The result of a successful tomographic retrieval is shown by Riese et al. (2014).

Pointing stability is characterised using IR image analysis methods combined with the evaluation of attitude sensor data. The analysis of image motion in the IR data leads to pointing information with high temporal resolution due to the detector frame rate of 6.3 kHz. Two different image analysis methods have been implemented and are currently tested. Housekeeping data of the navigation and gimbal control unit and the 3-axes MEMS gyro is used to complement this analysis. Results can be found in Maucher et al. (2014).

Spectral coverage and sensitivity of the instrument can be deduced from Fig. 10. The given values for dynamics mode are derived from the variance of a sequence of

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blackbody measurements during flight and can be seen as a conservative estimate for the performance during atmospheric measurements. The values indicate that GLORIA is closer to target than to threshold requirements over most of the spectral range. The spectral cut-off of the detector is rather soft and the region from 780 to 900 cm^{-1} is already in the cut-off slope leading to a small non-compliance in the wavenumber region below 800 cm^{-1} . Even better sensitivity was expected from analysis of individual pixels, but spatial coaddition to row averages does not fully improve the noise values as expected, hinting towards a spatially correlated electronic noise source. The Noise Equivalent Spectral Radiance (NESR) analysis in CM mode in Fig. 10 is performed with a “deep space” measurement. In this mode, the instrument complies with target requirements over most of the spectral range. Here the noise determination bases on the statistics over the 128 rows. A detailed comparison with expected performance and an in depth analysis of the noise contributors can be found in Friedl-Vallon et al. (2014).

The consistently good quality of temperature retrievals with an overall bias of approximately 0.5 K (see Fig. 6) is an indicator for a sufficient radiometric performance of the instrument. A detailed analysis using redundant in-flight calibration measurements is ongoing. First results are presented in Kleinert et al. (2014) and indicate a radiometric accuracy of better than 1.5 % over most of the spectral range.

Spectral accuracy requirements are not very stringent and can be met with moderate effort. Compliance is regularly checked in the retrieval process.

5 Conclusions and outlook

GLORIA has been deployed in three different scientific campaigns (ESSenCe, 2011; TACTS and ESMVal in 2012) on two different aircraft (Geophysica, HALO) in various climatological conditions. The instrument has demonstrated the capability of the infrared limb-imaging technique (Riese et al., 2005; Friedl-Vallon et al., 2006) to provide atmospheric parameters like temperature and trace gas fields with very high spatial resolution. This allows to address some of the most pressing scientific questions regarding

physical and chemical processes in the UTLs as outlined in the accompanying paper of Riese et al. (2014).

The initial instrument concept has proven to be adequate and the experience made in the first year of operation does not force to make substantial changes to the overall set-up or to essential subsystems. Some areas for further improvement have been identified, though. In particular spatially correlated noise sources in the detection chain like EMC and electronic noise will be addressed to improve sensitivity and radiometric accuracy over the full spectral range. Non-linearity characteristics of the HgCdTe detector arrays will be in depth characterised to allow optimal operation and ease processing. Upgrades of some electronics units will improve performance, reliability and operability of the instrument. In order to reduce acoustic and vibration excitation of the instrument the opening in the belly pod of the HALO aircraft will be modified to improve aerodynamics.

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We thank the TACTS and ESMVal coordinators for organising these extraordinary successful scientific campaigns and the whole team of DLR flight operations at Oberpfaffenhofen for excellent HALO flights, patience and trusting support. Last but not least we want to thank the certification team of DLR for steering us safely through the cliffs of the certification process.

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**Table 1.** Instrument specifications – geometry.

Translation from angular values to distance is made for an observer altitude of 15 km and a tangent altitude of 10 km.

Requirement	Value
Vertical spatial coverage	
Lower boundary	−3.3° below local horizon = typ. 4 km altitude
Upper boundary	0.8° to local horizon
Vertical sampling	Required: < 3 arcmin Realised: 1.9 arcmin = 140 m
Horizontal spatial coverage	Typically 1.5° = 6.7 km
Yaw pointing range	45° to 135°
Temporal sampling	Required: < 3 s / < 30 s for DM/CM Parametrisation during TACTS/ESMval campaigns: 1.2 s / 12 s interferogram length for DM/CM 2 s / 12.8 s temporal sampling for DM/CM
Pointing stability (vertical)	0.7 arcmin = 50 m (1σ)

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Table 2. Instrument requirements – spectrometry, (given as target/threshold, where applicable).

Requirement	Value
Spectral coverage	780 cm ⁻¹ – 1400 cm ⁻¹
Spectral sampling	< 0.625 cm ⁻¹ / < 0.0625 cm ⁻¹ for DM/CM
Sensitivity (NESR)	< 5 nW (cm ² sr cm ⁻¹) ⁻¹ /10 nW (cm ² sr cm ⁻¹) ⁻¹ for DM < 15 nW (cm ² sr cm ⁻¹) ⁻¹ /30 nW (cm ² sr cm ⁻¹) ⁻¹ for CM
Radiometric gain accuracy	1%/2%
Spectral accuracy	10 ppm

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Table 3. Basic instrument properties.

Property	Value
Mass of instrument	212 kg
Mass of control computer	28 kg
Overall length	1610 mm
Overall width	975 mm
Overall height	675 mm
Typical total power consumption	ca. 800 W
Typical raw data rate	600 Mbit s ⁻¹

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Table 4. Basic Spectrometer Properties

Property	Value
Mass of spectrometer	48 kg
Operating temperature of spectrometer	200–220 K
Entrance window diameter	100 mm
Max. optical path difference	90 mm double sided
Camera focal length	72.2 mm at 218 K
Aperture ratio	$f/2$
Detector type	256 ² HgCdTe-LFPA by AIM GmbH
Pixel pitch	40 μm
Operating temperature of detector	50 K

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Fig. 1. HALO fully equipped during take-off. The image was taken during the preparations for the TACTS campaign. GLORIA is mounted in the belly pod below the aircraft.

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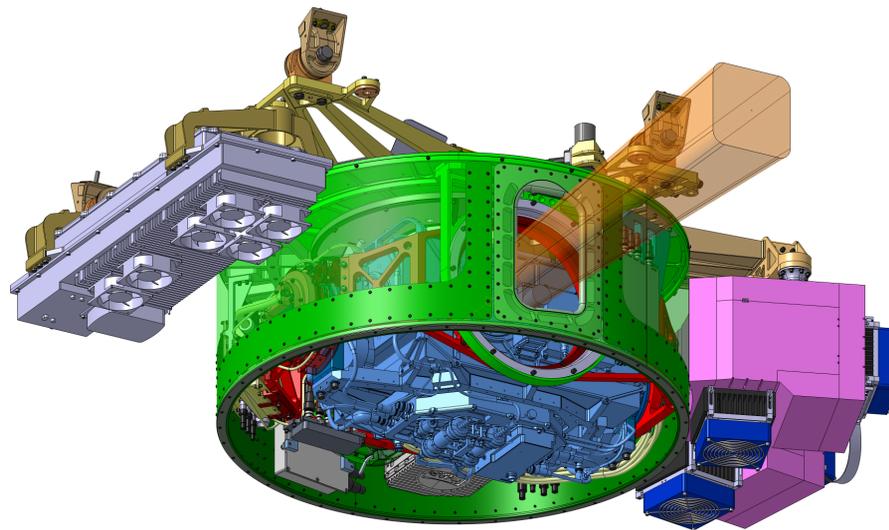


Fig. 2. CAD rendering of the GLORIA instrument. The different gimbals are colour coded: blue indicates the pitch frame including spectrometer, red the roll frame, green the yaw frame and gold the mounting frame to the aircraft. The transparent orange cone represents the field of view of the instrument. The grey box at the left hand side is the power electronics, the pink boxes at the right hand side are the blackbodies for radiometric calibration.

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Fig. 3. GLORIA mounted to the belly of the HALO aircraft.

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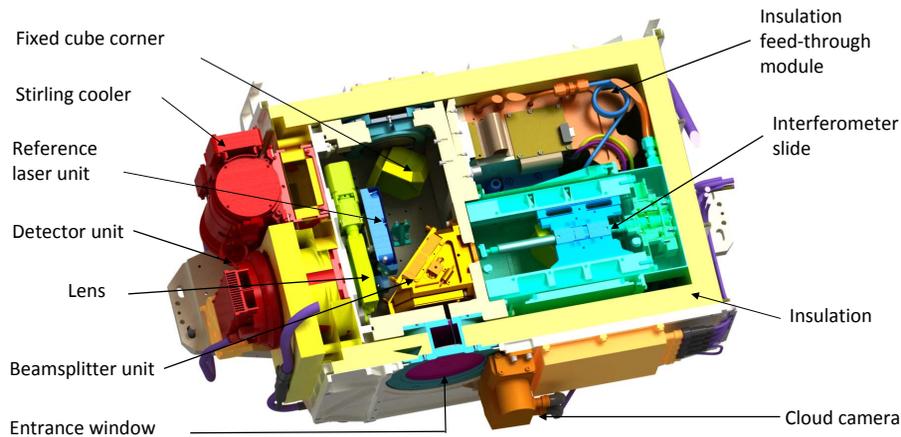


Fig. 4. CAD rendering of the GLORIA spectrometer including insulation. The instrument is cut in the horizontal plane to show the optical setup and the slide construction.

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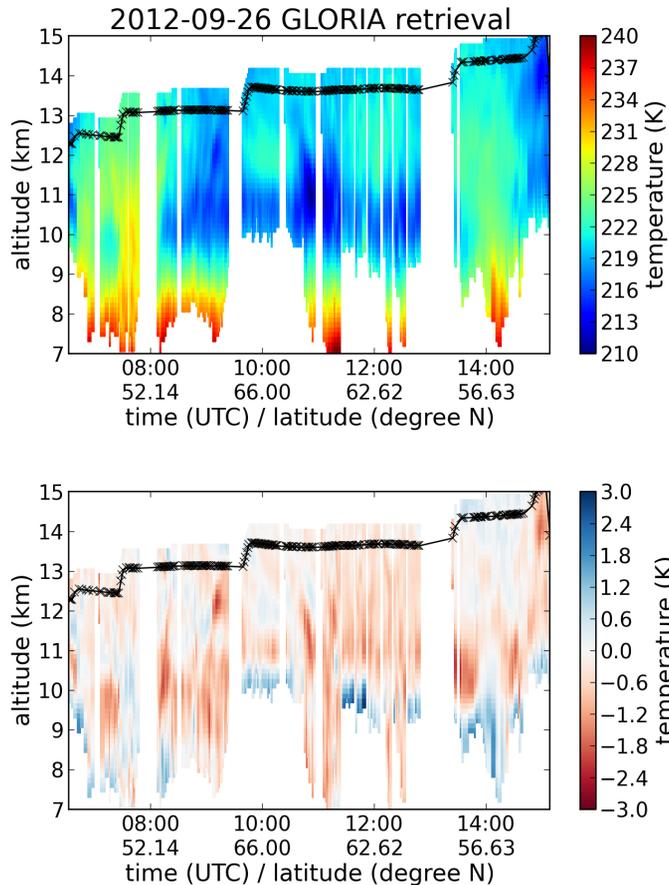


Fig. 5. Upper panel: retrieval result for temperature for the flight of 26 September 2012 above the North Atlantic Ocean. The flight altitude of HALO is indicated by black crosses. The lower boundary of the retrieval is determined by clouds. Lower Panel: temperature difference between the retrieval result and the ECMWF analysis.

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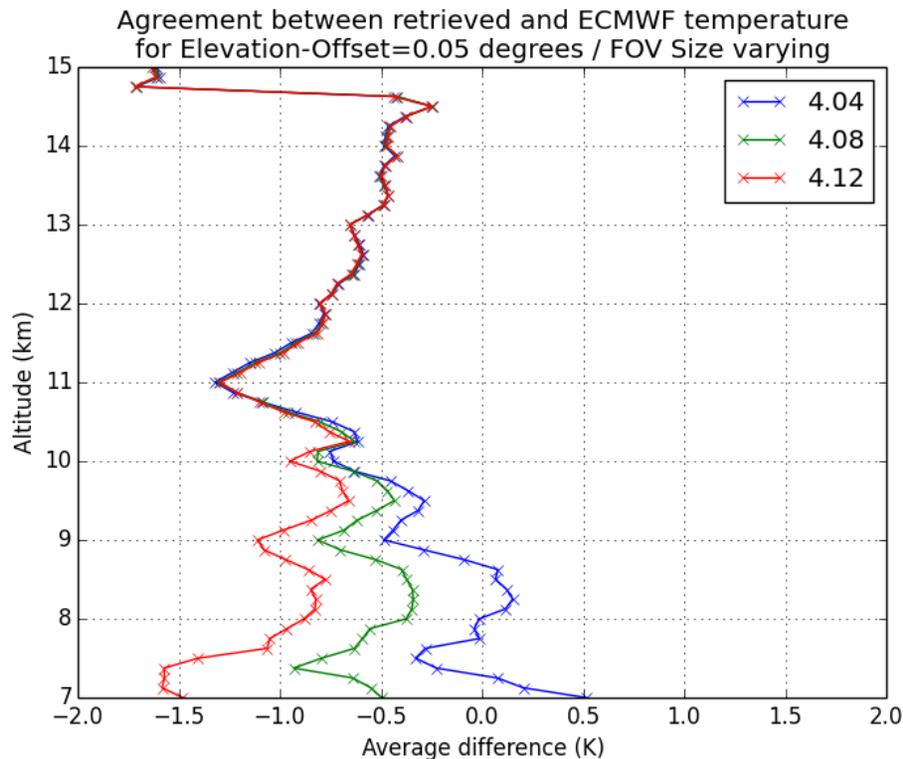


Fig. 6. Deviations of temperature retrieval results for different FOV sizes. Average difference between ECMWF analysis and temperature retrieval is shown for the flight of 26 September 2012. The blue line indicates the average difference for an assumed vertical FOV of 4.04° , the red line for an assumed vertical FOV of 4.08° and the green line for an assumed vertical FOV of 4.12° .

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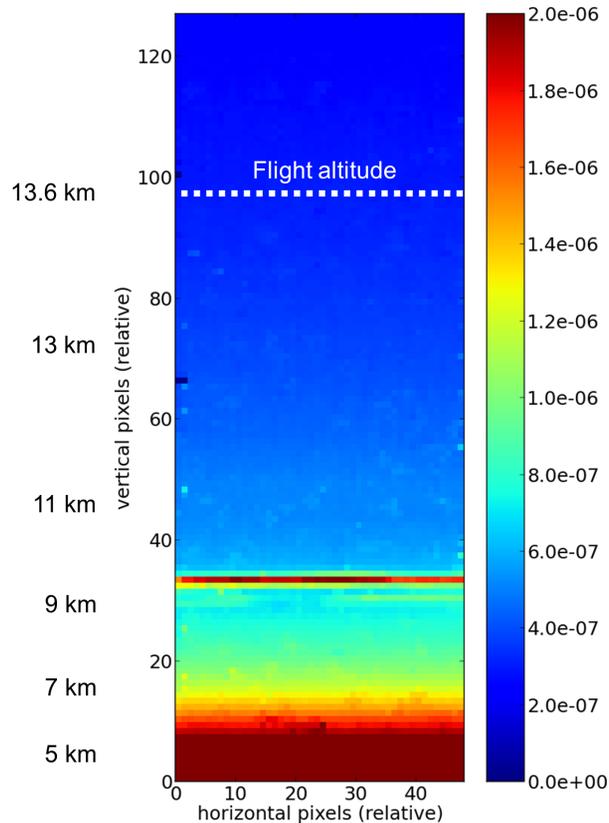


Fig. 7. Radiance image taken on 26 September 2012 over the North Atlantic Ocean. The image shows mean radiance in $\text{W}(\text{cm}^2 \text{sr cm}^{-1})^{-1}$ between 800 cm^{-1} and 950 cm^{-1} . Approximate altitudes of tangent points are indicated by the scale on the left hand side. A thin Cirrus cloud layer can be seen at an altitude of approximately 9.5 km. The white dotted horizontal line indicates the flight altitude of 13.6 km.

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F. Friedl-Vallon et al.

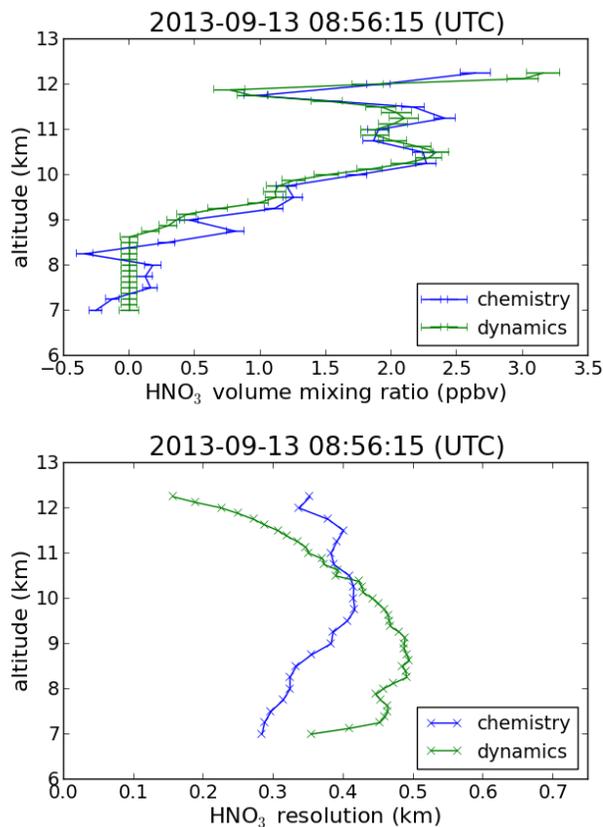


Fig. 8. Typical retrieval result for HNO₃ of 13 September 2012 between Cape Town and Antarctica. A data cube with chemistry mode spectral resolution has been processed with the chemistry mode retrieval code and for comparison also with dynamics mode spectral resolution and dynamics mode retrieval code.

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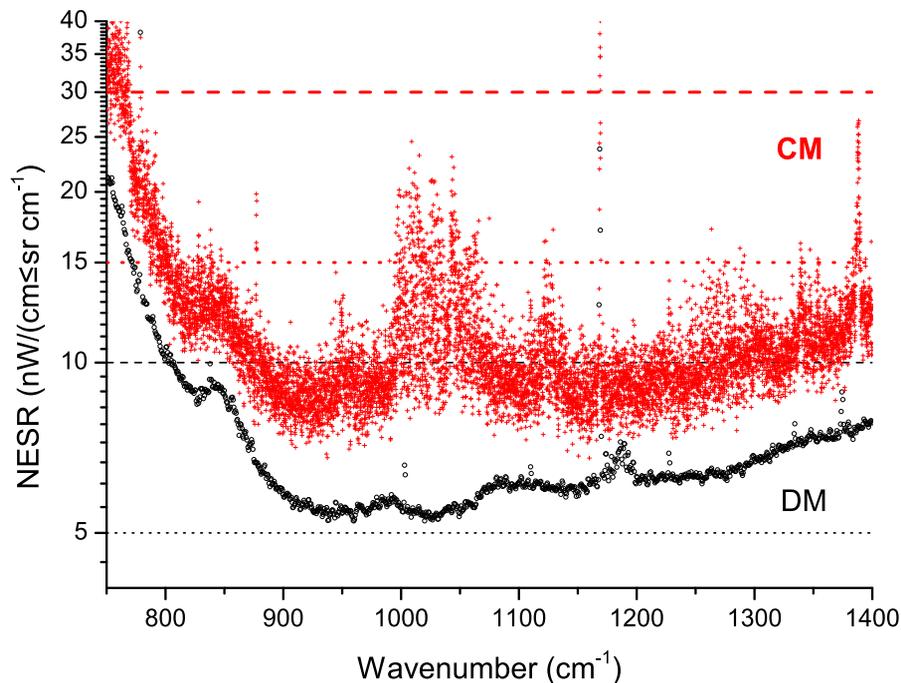


Fig. 10. Sensitivity analysis for a calibration sequence of the flight of 26 September 2012. The dotted/dashed lines in black and red depict the target/threshold requirements for DM and CM, respectively. The black circles represent the mean Noise Equivalent Spectral Radiance (NESR) value for the L1 product based on statistics of 43 consecutive real blackbody spectra. Spectral sampling is 0.625 cm^{-1} , total integration time is 1.2 s. The red crosses show the mean NESR for one “deep space” measurement based on the statistics over the 128 rows. Spectral sampling is 0.0625 cm^{-1} , total integration time is 12 s.

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