Paper Review

Felix Erdmann

AUTHORS RESPONSE TO EDITOR COMMENTS

1 RC1: Anonymous Referee #3, 30 Jul 2019

The authors thank the referee #3 for the detailed and constructive comments. We included the general and specific comments
in the updated paper manuscript. Based on the general comments, the results section is revised in total. Two Figures (former 14
5 and 15) are removed and we introduced two tables summarizing the statistics of investigated flash characteristics. The specific
comments are addressed in the following.
[We indicate the line number in the document showing track changes for each comment.]

1.1 Specific Comments

1. Sentence on lines 47-48: networks with "widely-spaced" VLF/LF sensors (like Meteorage and the U.S. NLDN)
10 **report far more cloud pulses than return strokes, because they are sensitive to vertically-oriented current-carrying**
channels in "larger" IC flashes (see Cummins and Murphy, 2009 or Nag et al., 2015). I would say that they have "...
somewhat limited total lightning detection efficiency (DE)." Shorter-baseline VLF/LF systems like LINET and
portions of the ENTLN (cited in this work) have very high total lightning flash DE.

The statement is shortened and implemented after the introduction and citation of Cummins and Murphy (2009) and Nag et al.
15 (2015).
Comment: The comment is true nowadays looking at the overall numbers of detected strokes/pulses (although it is different for
the DE), however, our Meteorage data include no IC pulses until 2008. The first year exhibiting more IC pulses than CG strokes
in Meteorage data was 2015 (within the study region of this work). Meteorage has still a higher (spatial) accuracy and DE for
CG strokes than for IC pulses. The signal strengths (currents) are also significantly stronger for CG strokes than IC pulses. The
20 predominant vertical channels within the cloud (Nag et al. 2015) do not exist for all IC discharges, thus their detection can be
difficult for (V)LF LLSs. Return strokes associated with CG strokes do indeed very often produce sufficient (V)LF radiation
to be measured by those LLSs (attenuation with distance to the discharge).

We want to point out that the detection efficieny (DE) of CG strokes is in general high for VLF/LF networks. The IC DE
depends on the baseline distance. In particular, Meteorage is not a "widely-spaced" network as stated in the comment. The
25 baseline distance is similar to that of e.g. LINET in Germany. The overall DE of Meteorage, with high CG DE and somewhat
lower IC DE, is sufficient for a comparison to optical satellite instruments like ISS-LIS (it likely exceeds the ISS-LIS DE in

most cases).

[please see changes in lines 54-60]

2. Line 57: This is the first use of the term "relative DE", and (unfortunately) one of the referenced papers (Bitzer et al, 2016) use this term to mean the ratio of the conditional probabilities (see their equation (2)), but others do not use this definition. SO – you need to say that you use the common definition of relative DE used in many studies, which is the percentage of matched flashes divided by the number of flash in the other (reference) LLS (if that is what you are doing).

The relative DE is briefly introduced prior to the literature review.

[please see changes in lines 48-49]

3. Line 138: the sentence about "...within 330 ms and 5.5 km..." is not really correct. No group associated with these values will be included in the flash - these are normalizing parameters for a Euclidian distance measure (see Mach et al., 2007 - figure 2)

The weighted Euclidean distance concept is introduced in the paper to correct the sentence. **[please see changes in lines 154-156]**

4. Line 162: The Meteorage network only requires two sensors to report a lightning discharge (see Cummins and Murphy, 2009).

We confirm that two sensors are sufficient. They use the time and angle of arrival to create four variables for three target values (time, latitude, longitude).

[please see changes in line 182]

5. The description of LMA in lines 169-177 is not quite correct and should be reviewed by co-authors on your associated SAETTA paper. It reports leader development (associated with breakdown processes or fast leader propagation in established channels), serving to produce a spatial map of possible paths for later, high-current processes in the flash. The phrase "Fast CG discharges traveling between the cloud and ground" is not correct – you are probably referring to dart leaders that can occur in pre-established leader channels for earlier CG strokes, as well as in established channels within IC flashes.

We removed the phrase "Fast CG discharges travelling between the cloud and ground in time frames shorter than 80 s might be missed." It should indicate instead that flash components with continuous VHF radiation would be hardly mapped with a LMA due to time-of-arrival technique and the sampling method.

[please see changes in line 194]

6. Lines 211-213: regarding the authors' discussion of the rationale for using events rather than groups in the flash algorithm: This reviewer (and many others in the references cited in this work) firmly believes that reports from mid- to long-range ground-based VLF/LF LLS networks match well with LIS groups. Long vertical channels during periods of high currents provide localized light sources, and most of this light is produced during periods of less than 2 ms. The timing and centroid location of a LIS group are a good match for such sources. Figures 8 and 9 in Bitzer et al. (2016) show the very tight time- and space-correlations for these discharges. However, this does NOT mean that ALL LIS groups are space- and time-correlated with VLF/LF strokes/pulses. The authors' accompanying rationale related to lightning mapping and comparisons with SAETTA are quite reasonable. I ask that the authors refine/revise the rationale on these lines under the light of these comments.

The formulation was too stringent. The authors agree that LIS groups often represent similar physical phenomena (discharge processes) as (V)LF pulses/strokes. The significantly higher number of LIS groups than (V)LF sources shows that LIS groups can indicate additional discharge mechanisms that are not seen in the (V)LF range. The rationale is adapted.
[please see changes in lines 243-244]

7. Lines 284-292: regarding flash matching: This text seems to indicate that flashes are matched if any elements in the two LLS meet the time constraint, and if any elements (the same or different that the ones meeting the time constraint) meet the spatial constraint. This seems problematic, when multiple flashes are closely spaced in time and space. I ask that the authors clarify this point, to be sure that the algorithm is described properly.

The matching algorithm uses a combined distance-time criterion. The same elements must meet both the spatial and the time constraint. Flashes are only matched if at least two elements (one per flash) meet both the spatial and the temporal criteria for a match.
Section 2.5 is revised in total.
[please see changes in lines 317-323]

8. Lines 397-403: The group:pulse distances should not necessarily be greater than the event:pulse distance, because the group locations are interpolated to sub-pixel spatial resolution by the radiance-weighting of the spatial centroid. Also the position differences reported by Bitzer et al. (2016) are much larger than the actual position differences because of +/- 5 km latitude location offsets associated with LIS yaw maneuvers (see Zhang et al., 2019). This issue is fixed in ISS-LIS. You may want to refine your analysis in the light of these issues.

It is true that group centroids can be located in any location while event locations are fixed to the pixel centers. Group centroids can in fact be closer or further away from VLF/LF pulses/strokes (which are not bound to pixel centers obviously). One should expect statistically about the same distance between VLF/LF pulses/strokes and LIS events and groups. We removed the statement and added a short discription of the difference between event and group locations.
Zhang et al. (2019) provide very interesting findings and it is a valuable reference to be cited here. The correction during

TRMM yaw maneuvers results in slightly lower distances to ground-based LLS for TRMM-LIS groups (1-2 km in Zhang et al.) than for ISS-LIS events in this study (2-4 km).

90 **[please see changes in lines 447-460]**

9. Lines 469 and later: the units for the LIS radiance product is not correct, in terms of the spectral density. The units should be nm, not μm. See Zhang et al., 2019, which also shows that the TRMM-LIS minimum radiance is about 3 μJSr-1m-2nm-1, indicating that it has a lower threshold than ISS-LIS!

The unit issue was further examined in collaboration with D. Buechler (University of Alabama Huntsville, Huntsville, AL, USA). We found that the available version of ISS-LIS P0.2 had not yet included the calibrated radiance. The radiance variable in the data has the same values as (and is therefore identical to) the uncalibrated (raw) optical amplitude count. Hence, all analyses of optical signal strength ("radiance") are actually amplitude counts. The paper terminology is changed and the corresponding Figure labels are updated.
[please see changes in lines 162-166 and e.g. in lines 532-533]

100 ## 1.2 Editorial Corrections and Suggestions

10. Introduction, first paragraph: it might be helpful for some readers to also know that total lightning flash rate within a storm is associated with storm intensity features such as ice flux and updraft volume and rate (see Deierling and Petersen (2008) and Deierling et al., 2008, among others)

Suggested references were cited with a short summary of their work.
105 **[please see changes in lines 27-32]**

11. Sentence starting on line 41: "SAETTA" and "mapping" have not been introduced yet, and this sentence is probably not required at this point in the paper.

The second part of the sentence was indeed not required here and is removed.
[please see changes in lines 46-47]

110 **12. Line 50: the term "LIS groups" is used, but there has been no description of the LIS products (provided later in the manuscript). It might be helpful to do this in 2-3 sentences, or point the reader to section 2.1.**

Groups are briefly explained.
[please see changes in line 64-65]

13. Line 90: suggest replacing "for GLM" to "as GLM"

115 ok.
[please see changes in line 105]

14. Line 94: "orbits" should be "orbit", since there is only one orbit

ok.

[please see changes in lines 109]

120 **15. Sentence starting on line 105 ("Among the...), does not seen to be a complete sentence.**

The sentence was misleading and should be clear now.

[please see changes in lines 120-122]

16. Line 112: suggest changing "... chacteristics in lightning detection..." to "...lightning detection characteristics..."

ok.

125 **[please see changes in lines 127]**

17. Line 119: suggest changing "our paper" to "this work"

ok.

[please see changes in lines 134]

18. Lines 151-152: suggest changing "... from the lightning discharge on earth..." to be "... from the optical source at
130 **cloud-top..."**

ok.

[please see changes in lines 171-172]

19. Line 196: The sentence starting on this line does not make sense to me. Two events in a flash can be much farther
apart than 14.3 km. If you are referring to adjacent events then I do not see how they can be 11.9 km apart at nadir.
135 **Please clarify.**

We tried to give a value for the distance constraint of the NASA ISS-LIS clustering algorithm to be compared to our in-house
algorithm. We cannot really identify such a value between the events (additional communication to D. Mach [NASA MSFC,
Huntsville, AL, USA]). The NASA ISS-LIS clustering algorithm uses a weighted Euclidian distance (space and time) between
group centroids. The paper is corrected.

140 **[please see changes in lines 225-229]**

20. Line 204: suggest changing "group" to "collection", given the LIS definition of group

ok.

[please see changes in line 235]

21. Line 222: Use of the words "initial element" suggests that all the sources must be within the ds limit of the first source. This can cause spatially propagating flashes to be broken up. This should be clarified.

The term "the initial element" is replaced by "any element of the flash". The algorithm can indeed treat spatially propagating flashes.

[please see changes in lines 255-259]

22. Line 299: change "onl yup" to "only up"

ok.

[please see changes in lines 342]

23. Section 3.2: Just a comment: It seems that the distance and timing offset distributions are produced by comparing all LIS events and all Meteorage pulse/strokes, so (for LIS) it mergers any underlying space:time correlations for individual matched pairs with the timeevolution of light that is observed by LIS, due to things like leader propagation and continuing current in long channels.

We do not completely understand this comment. We would appreciate if you explain your comment further.

24. Lines 444-445: a contributing factor could also be related to increasing length of the optical sources due to finite leader and return-stroke velocities.

Do you mean we need to consider the time of propagation of the leader or return stroke to explain the time difference between the optical signal and the Meteorage CG stroke?

25. Lines 516-518: does this paragraph belong before the previous paragraph?

Yes, the brief summary fits better directly after the comparison and before discussing the exceptionally long flashes.

[please see changes in lines 581-590]

26. Line 528: The vertical displacement of LMA sources in the xlma display at large distances from the center of the network have always troubled me, even after speaking to the developer of xlma. It seems that refraction is not being handled, so that distant source heights are not really useable. Additional insights would be nice but are not required.

Currently the SAETTA data processing is similar to the other LMA processing chains and does not include any refraction correction. As a matter of fact electromagnetic waves propagating in the clear sky are downward deflected because of the refractive index gradient that is most often downward directed. If refraction is not taken into account in the calculation of the VHF source position, the calculated altitudes may overestimate the true altitudes for distant events from the network. This is something that should be investigated in the future.

VI

27. Line 543: suggest changing "It can be constituted that..." to "Overall, ..."

Ok.

[please see changes in line 618]

175 **28. The term "average mean radiance" (and similar terms that reflect statistics of statis- tics) require the reader to think hard to interpret the variables. Possible better wording could be "the mean radiance averaged over all heights" or something like that**

Okay, the wording is clarified.

[please see changes in e.g. lines 633-634]

180 **29. Line 550: is it "maximum radiance per flash" or "maximum event radiance per flash"?**

It is the highest amplitude count (former stated as radiance) observed for an event of a flash. The term maximum event amplitude count per flash is more precise here. **[please see changes in e.g. lines 633-634]**

30. Line 558: suggest changing "dark" to "darker", since they are not really dark (they can be seen).

Ok.

185 **[please see changes in line 636]**

31. Line 583: A reference for the 10 kA value would be helpful

The value was inititally chosen to discuss the observed differences in the results. For example, Cummins and Murphy (2009) also use values of 10 kA and 20 kA to distinguish "small events" and "larger events", respectively, and to identify CG strokes and IC pulses.

190 **[please see changes in lines 675-681]**

32. 591-592: This seems to be an over-statement. The relationship between current, polarity, and height will vary with storm polarity and type, and falls apart for hybrid (IC+CG) flashes.

The possibility of a different relationship in inverted polarity storms (and in other regions) is added to the discussion in the paper. The relationship of the polarity of the maximum current and the flash maximum altitude might also be distorted for

195 hybrid flashes. The fact is added to the discussion in paper. We do not have data to prove this idea.

[please see changes in lines 692-695]

VII

33. Conclusions section: Present tense should probably not be used, since the findings may not be universally applicable.

The use of present tense in the conclusions is revised and changed for study specific findings.

please see changes in section 4, page 23

34. Line 614: it would improve clarity if you replaced "to an overall equal proportion" with "when considering the complete dataset"

The author does not totally agree with the suggested wording. The study presents two distributions (for ISS-LIS and for Meteorage). The sentence provides the information that both distributions feature a similar ratio of negative values (indicating the match was detected prior to the source). The term "to an overall equal proportion" is replaced with "with similar probability".

[please see changes in line 716]

35. Line 621: Zhang et al. (2019) might be a good reference to add here

Ok.

[please see changes in line 723]

36. Line 626: suggest replacing "dark" with "darker"

Ok.

[please see changes in line 728]

37. Lines 626:640: It might be useful to add that all of these height-related behaviors are likely driven by the range of heights associated with CG flashes vs. IC flashes.

A short sentence is added mentioning the flash types.

[please see changes in line 729-730]

2   RC2: Anonymous Referee #1, 31 Jul 2019

The authors thank the referee #1 for his constructive comments. We included the general and specific comments in the updated paper manuscript. Based on the general comments, the results section is revised in total. Two Figures (former 14 and 15) are removed and we introduced two tables summarizing the statistics of investigated flash characteristics. The specific comments are addressed in the following.

[We indicate the line number in the document showing track changes for each comment.]

2.1 Specific Comments

1. Check the order of introduction of instruments in Section 1 (Introduction) to make sure each instrument is described before being referred to by its acronym. There are also cross-references between instruments in odd places, such a mention of LF/VLF systems in the paragraph (line 80) discussing VHF systems. Overall, I thought the introduction could be shortened somewhat to focus more on the aims of the study, with less context about the lightning detection problem as a whole.

The introduction was revised and refined. It should be clearer in terms of structure and the references. For example, papers of Cummins and Murphy (2009) and Nag et al. (2015) were added that provide summaries of the different LLSs. The introduction focuses more on the former studies and less on sensor characteristics than before.

[please see changes in section 1, pages 2-5]

2. Line 84: WWLLN relies on ionospheric reflections but operates at VLF. As noted on line 164, the LF measurements in this study do not include ionospheric reflection.

This technical paragraph was shortened in order to focus on the literature review.

[please see changes in line 93-99]

3. Fig. 1c may be two lightning flashes in the VHF data, and if it was automatically identified by an algorithm illustrates the challenges in flash classification. There is a large gap in the channels to the SE, and if this were one flash I would expect to see that gap filled given the otherwise very well-resolved channels. Was there evidence of a new channel developing and exhibiting bidirectional development in the flash to the SE at ∼0.1 elapsed seconds?

As a general comment the automatic algorithms to identify flashes from signals like VHF sources, LF pulses/strokes or optical events work in a statistical sense, however, neither the algorithms nor an observer can always distinguish individual flashes perfectly.

The flash in Figure 1c was identified as one flash by three different algorithms. Branches to both sides of the mentioned gap to the SE propagate towards each other. We looked at the flash development through a series of animations using SAETTA observations. We also looked in the raw LMA data to evaluate the idea given in the comment. In fact, there was a significant amount of VHF records and we could not clearly identify the single branches. Conclusively, we think that Figure 1c shows one flash that originated as two flashes whose channels connected. The connection is not directly visible in the data, as it is likely masked by the lack of reconstruction.

[please see changes in Figure 1, page 28]

4. Line 180: Recent studies by Chmielewski and Bruning (2016, 10.1002 / 2016jd025159) and Koshak, Mach and Bitzer (2018, 10.1175 / JTECH-D-17-0041.1) show that changes to the network geometry can have a significant influence on detection efficiency and location precision. This effect may be important for the SAETTA network which has a long N-S baseline, in contrast to the somewhat more circular and compact network in Thomas et al. (2004).

255 We agree with Referee #1 about the influence of the network geometry on the location precision. Coquillat et al. (2019) show that the displacement of 2 stations located in the south-west and south-east of Corsica towards 2 positions located in the extreme south and the extreme north of the island in 2016 led to a modification of the errors (Figure 3 in Coquillat et al., 2019). The radial error was markedly reduced except in the north of the studied domain meanwhile the altitude error increased almost everywhere at larger distances from the network. Therefore when different sets of at least 6 stations are involved in the

260 calculation of the VHF source position one would expect a different geometry of the network, which influences the location precision. Nevertheless given the maximum SAETTA location errors (and errors within 50 km of the LMA) that are given in section 2.1, the location accuracy is sufficient for an intercomparison to both Meteorage and ISS-LIS as the maximum SAETTA location errors are not greater than the location errors of the LSSs in the comparison.

As far as detection efficiency is considered, Chmielewski and Bruning (2016) found in general high DE (>90 %) for all

265 networks with different receiver thresholds up to a range of 150 km from the LMA center. Within a distance of 200 km, the modeled DE is still greater than 85 % in their study. The actual DE of SAETTA is difficult to address as the region and conditions are different from the regions used by Chmielewski and Bruning. One study is currently underway to assess SAETTA location accuracy through a comparison between SAETTA records and aircraft tracks in ice clouds.

The Koshak et al. (2018) paper is an interesting supplement and it is now given as a reference in section 2.1.

270 **[please see changes in line 200-208]**

5. Lines 182-183: Does the study domain refer to the 350 km max range of detection or the somewhat smaller lat/lon box at the beginning of section 3?

The study domain is limited to the region shown in Figure 1(a). It is introduced in the opening paragraph of section 2 to be within 40.5 °N to 44.0°N and 7.0°E to 11.0°E. The maximum distance to SAETTA's center approximates 271 km (the SW

275 corner of the domain) and 238 km (the NE corner of the domain).

The reference to Figure 1(a) was added to clarify the meaning of study domain.

[please see changes in line 211]

6. Line 211: I don't have a concern here, so much as I wish to highlight that the authors raise an essential question about the measurements: "it is questionable whether LIS groups really correspond to (V)LF pulses/strokes." I agree with prior studies that the group is the fundamental physical measurable from the instrument - it is an ~instantaneous light emission tied to heating by a "large" current flow along a channel, and the events register the extent of the light scattered by that process. However, the authors are also right to point out that not all (V)LF pulses/strokes have a corresponding group, which suggests that either the instruments are sensitive to different physics, or the (V)LF and optical measurements lack the necessary sensitivity to see what is actually the same physics. In the end, the authors' approach of clustering using the events instead of the group centroids is a good choice, since considering events will help them align with LIS data and better identify coincident ground strokes that might happen at some distance from the centroid of the light emission as observed at cloud top, but I would disagree that the events are the fundamental physical detection.

Yes, it is difficult to find a physical process representing the optical LIS events. Groups are likely the best representation of fundamental discharge processes in optical lightning data. Events should be seen as smallest, elementary measured units of LIS-like instruments. The rationale on these lines was refined (e.g., the word "fundamental" in line 234 was replaced by "elementary" to be consient with the further use of the term *element* for LIS events and LF pulses/strokes).

[please see changes in lines 237, 243-244]

7. In the paragraph beginning line 325, it is ambiguous whether the authors think the missed flashes nearest SAETTA were truly missed or if station downtime were to blame. This is especially interesting because the authors close the paragraph by stating that SAETTA is not an absolutely reliable DE reference. Pédeboy et al. (2018) is probably more explicit about the details, but it is not in a peer reviewed article; regardless it would be helpful to clarify here what the authors mean.

We looked in detail in our raw SAETTA data, with information from individual stations. It is found that VHF signals were observed for most of the missed (by SAETTA) flashes (114 of 120). The most important reasons for the classification as missed were signals at less than 6 stations and thus no reconstructed VHF source (77 of 120) and locating/timing differences just exceeding the algorithm criteria of 0.2 °/0.3 s between observed flashes and VHF sources (34 of 120) or reconstructed VHF sources had too high reduced chi² values and were disregarded by our algorithm. As we do not know the origin of the signal in the raw data, we cannot assign raw data signals at individual stations to single flashes. This needs the reconstruction of VHF sources using at least 6 stations.

The missed flashes nearest SAETTA were not detected by SAETTA due to a reduced number of active stations (5 of 15), reconstructed SAETTA sources showed a slightly too high chi² uncertainty (5 of 15) or SAETTA reconstructed sources are found slightly shifted in space/time to a Meteorage flash (3 of 15). Only 2 Meteorage flashes in the proximity of SAETTA were really not detected.

[please see changes in lines 368-383]

8. Given the predominance of flashes that occur outside SAETTA, altitude errors will be large and the total number of sources detected per flash will be small. How does this affect the results that depend on altitude retrieval from SAETTA in section 3.3?

SAETTA detection range allows in general for many (order 100) sources per flash up to distances of about 350 km. The issue
315 regarding the detection of low altitude sources in large distance to the LMA was added in the section 2.3. SAETTA is even capable of detecting sources lower than the theoretical values added in section 2.3, because most stations are well above sea level and the highest station is located at an altitude of 1950 mASL. In theory, this station has a direct line of sight to sources at about 1.6 kmASL altitude in a distance of 300 km to the station (which would be outside the study domain).

The altitude errors are given in section 2.3 and are expected to be less than 500 m within the entire domain (maximum
320 distance to the LMA center of 270 km in the SW corner, significantly lower altitude errors closer to the network). The results in section 3.3 are discussed taking the theoretical altitude errors into account, however, actual errors also depend highly on the set of stations used for the reconstruction of the VHF source and the source altitude. This is especially true for each individual VHF source. The overall error of a sample of VHF sources should be more consistent and closer to the theoretical model than individual source errors. Flash altitudes in section 3.3 use statistical values of concurrent SAETTA sources (i.e. the mean, 10th
325 and 90th percentile) and should represent the statistical flash altitudes within the given range of uncertainty.

9. Line 595: This statement is accurate for the authors' data, but I would predict storms with inverted polarity would have the opposite expectation for detectability as function of polarity. This context would be helpful if another region were to be studied with the same methodology.

It will need definitely further research to investigate how this relationship between polarity of the maximum (LF) current and
330 the flash altitude behaves in different regions and different storm types. The suggested context is added to discussion in the paper.
[please see changes in line 692-695]

2.2   Technical Corrections

After revisions are completed I recommend an additional read for flow and a few missing words. For instance, on line
335 **80: "uses very high frequency (VHF)" needs "radio signals" or some other noun at the end of the sentence.**

The entire paper is revised.

The colon at the end of line 46 seems like it was from an earlier revision where the instruments were introduced in a different way?

Yes, it was.
340 **[please see changes in line 53]**

Lines 609-13: "peak" here could be misinterpreted as "maximum." I suggest "mode" or "most frequent."

Ok.

[please see changes in line 712]

ADDITIONAL TRACK CHANGES

Instance	Description
Figure 1	Time series added (d)
Section 2.1	Raw amplitude count introduced as optical signal measure
Section 2.5	Revised and slightly shortened
Section 3.3	Revised
Table 2	New table added (section 3.3)
Table 3	New table added (section 3.3)
Figure 13	Labels updated (Radiance -> Amplitude count)
Figure 14	Removed (information gain after Figure 13 is low)
Figure 15	Removed (results in Table 3 and Figure 16)

345

Concurrent Satellite and ground-based Lightning Observations from the Optical Lightning Imaging Sensor (ISS-LIS), the LF network Meteorage and the SAETTA LMA in the northwestern Mediterranean region

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Abstract. The new space-based Lightning Imager (LI) on board the Meteosat Third Generation (MTG) geostationary satellite will improve the observation of lightning over Europe, the Mediterranean Sea, Africa and the Atlantic Ocean from 2021 onwards. In preparation of the use of the upcoming MTG-LI data, we compare observations by the Lightning Imaging Sensor (LIS) on the International Space Station (ISS), which applies an optical technique similar to MTG-LI, to concurrent records of the Low Frequency (LF) ground-based network Meteorage. Data were analyzed over the northwestern Mediterranean Sea from March 01, 2017 to March 20, 2018. Flashes are detected by ISS-LIS using illuminated pixels, also called events, within a given (2.0 ms) frame and during successive frames. Meteorage describes flashes as a suite of Intra-Cloud/cloud-to-cloud (IC) pulses and/or Cloud-to-Ground (CG) strokes. Both events and pulses/strokes are grouped to flashes using a novel in-house algorithm.

In our study, ISS-LIS detects about 57 % of the flashes detected by Meteorage. The flash detection efficiency (DE) of Meteorage relative to ISS-LIS exceeds 80 %. Coincident matched flashes detected by the two instruments show a good spatial and temporal agreement. Both peak and mean distance between matches are smaller than the ISS-LIS pixel resolution (about 5.0 km). The timing offset for matched ISS-LIS and Meteorage flashes is usually shorter than the ISS-LIS integration time frame (2.0 ms). The closest events and pulses/strokes of matched flashes achieve sub-millisecond offsets. Further analysis of flash characteristics reveals that longer lasting and more spatially extended flashes are more likely detected by both ISS-LIS and Meteorage than shorter duration and smaller extent flashes. ISS-LIS' relative DE is lower for daytime versus nighttime as well as for CG versus IC flashes.

A second ground-based network, the Very High Frequency (VHF) SAETTA Lightning Mapping Array (LMA), further enhances and validates the lightning pairing between ISS-LIS and Meteorage. It also provides altitude information of the lightning discharges and adds a detailed lightning mapping to the comparison for verification and better understanding of the processes. Both ISS-LIS and Meteorage flash detections feature a high degree of correlation with the SAETTA observations (without altitude information). In addition, Meteorage flashes with ISS-LIS match tend to be associated with discharges that

occur at significantly higher altitudes than unmatched flashes. Hence, ISS-LIS flash detection suffers degradation with low-level flashes resulting in lower DE.

1 Introduction

25 Lightning defines electrical discharges within the atmosphere. The discharges can happen within a cloud or between clouds (IC) or between a cloud and the ground (CG). The total lightning activity (IC+CG) is of interest for e.g. the numerical weather prediction (NWP) as lightning serves as tracer for deep convection. The total lightning flash rate is associated to storm intensity features. For example, Deierling and Petersen (2008) found strong correlation between the updraft volume above the $-5\,°\text{C}$ level in clouds and total lightning activity. Deierling et al. (2008) show a fairly stable relationship and strong correlation

30 between the precipitation ice mass flux, the nonprecipitation ice mass flux and their product on the one hand and the total lightning flash rate on the other. Especially graupel/small hail ice mass correlates well with the mean total lightning rate in their study. Among others Mattos et al. (2017) investigated the life cycle of thunderstorms and processes leading to the different discharge types. They found in their analysis of 46 isolated thunderstorms that in 98% of their cases, the first CG flash is preceded by IC lightning by approximately six minutes on average.

35 At this time, lightning observations in Europe are mainly made with ground-based sensors. To maximize the impact of lightning data on the assimilation in NWP systems, total lightning should be observed continuously over large areas.

At this time, lightning observations in Europe use mainly ground-based sensors. In a few years, the new Lightning Imager (LI) on board the Meteosat Third Generation (MTG) satellite (Stuhlmann et al., 2005) will provide continuous lightning observation over Europe, the Mediterranean Sea, Africa, the Atlantic Ocean and parts of Brazil. The satellite sensor will be

40 able to detect the total lightning including CG and IC flashes when it is launched in the 2021 time frame. The Lightning Imaging Sensor (LIS) on the International Space Station (ISS) (Blakeslee and Koshak, 2016) creates a unique opportunity to provide proxy data to help prepare research and operational applications for the MTG-LI data. It overpasses, among others, wide parts of Europe, including the entire Mediterranean region. ISS-LIS is in principle similar to the planned MTG-LI, so that ISS-LIS data can to some extent mimic the upcoming MTG-LI data. In addition, a comparison between European

45 ground-based lightning observation networks and ISS-LIS should improve the understanding of ground- and space-based lightning observationobservations. All instruments and networks are hereafter simply referred to as lightning locating systems (LLSs)bearing in mind ISS-LIS and SAETTA do actually map lightning.

This study uses the term relative DE. It is defined as the ratio of the number of matched flashes to the number of flashes in the other (reference) LLS, expressed as a percentage.

50 A LIS instrument was previously operational on the Tropical Rainfall Measurement Mission (TRMM) satellite (e.g., Christian et al., 1999; Cecil et al., 2005). Several LLS LLSs comparisons exist for regions covered by TRMM-LIS. The focus of the following (not exhaustive) literature review is on observational analyses rather than laboratory experiments, e.g. Boccippio et al. (2002). There are a variety of ground-based LLSs that have been utilized:

~~Very~~ Ground-based LLS observe different frequency ranges of the lightning radio signal. They are classified as e.g. very

55 low frequency (VLF) and low frequency (LF) LLSs ~~can detect lightning over hundreds or thousands of kilometers with the drawback of mainly detecting CG return strokes and hence a limitedoverall detection efficiency (DE)~~as well as very high frequency (VHF) LLSs (e.g. Figure 2, Cummins and Murphy, 2009). A summary of detection characteristics, (dis)advantages and the range of the various ground-based LLSs are provided in Nag et al. (2015). VLF/LF systems detect lightning on mid to long range. Their DE is somewhat limited. It varies for different networks and flash types (CG flash DE is ususally higher than

60 IC flash DE for VLF/LF LLSs) but increases in general with lower baseline distance.

~~Aiming~~ Thompson et al. (2014) aimed at exploring suitable proxy data for the Geostationary Lightning Mapper (GLM) (Goodman et al., 2013)~~, Thompson et al. (2014)~~. They report a pulse/stroke DE maximum for two long range LLSs, the World Wide Lightning Location Network (WWLLN) and the Earth Networks Total Lightning Location Network (ENTLN), of 18.9 % and 63.3 %, respectively, relative to 18-months records of TRMM-LIS groups (a combination of adjacent illuminated pixels in

65 the optical image that occur in the same 2 ms time frame). The maxima were found over the Pacific Ocean for WWLLN and near North America for ENTLN (within the analyzed region with the highest sensor density) in 2010 and 2011. They did not study how many WWLLN and ENTLN pulses/strokes had coincident TRMM-LIS groups.

Rudlosky et al. (2017) analyzed the performance of the Global Lightning Dataset 360 (GLD360) relative to TRMM-LIS from 2012 to 2014 in different regions. GLD360 was able to detect 63.6 % of the TRMM-LIS flashes in North America in

70 2014, the maximum DE reported in their study. The performance steadily increased from 2012 to 2014. The relative DE of GLD360 increased with the TRMM-LIS flash duration, flash extent and group number. The mean (median) location offset of the nearest GLD360 stroke to the matched TRMM-LIS flashes was 8.7 km (7.0 km). Rudlosky et al. (2017) applied the assumption that TRMM-LIS would detect all flashes within its field of view, but did not study the reverse problem, i.e. the relative DE of TRMM-LIS to the GLD360 flashes or strokes.

75 Defer et al. (2005) used both the UK Met Office long-range VLF sferics ATD system and TRMM-LIS to study the lightning activity in the eastern Mediterranean Sea for 20 days during winter 2008-2009. For their investigation of the flash scale, they developed and employed their own algorithm for TRMM-LIS flashes. The flash density analysis exhibits a general agreement between ATD and TRMM-LIS. The relatively small data set, and the fact that ATD detected mostly CG lightning, limited the ability for gaining overall statistics.

80 Bitzer et al. (2016) tested a Bayesian approach on the DE of TRMM-LIS and ENTLN by implementing the conditional DEs of the two LLSs relative to each other. They found a relative conditional group-to-pulse DE of 52 % (27 %) for TRMM-LIS to ENTLN (ENTLN to TRMM-LIS near North America in 2013). They also addressed peak timing differences and distances for the collocated discharges (again LIS groups, ENTLN pulses; see section 3.2 for details). Bitzer et al. tested further the effect of assimilating one dataset into the other on the detected number of discharges, i.e. 23.6 % of discharges could be added to

85 TRMM-LIS records.

While the previous papers focused on the DE, Höller and Betz (2010) analyzed TRMM-LIS and a VLF/LF lightning location network (LINET) in order to generate random proxy optical data from a given set of LINET data using model distribution functions. The outcomes are of specific interest for proxy data for the MTG-LI. Besides the relative DEs (approximately 50 %

for both LLSs), they investigated distribution functions and correlations between TRMM-LIS group and LINET pulse/stroke number per flash, flash extent and duration and between LINET pulse/stroke amplitude and TRMM-LIS group radiance. Although the Pearson correlation coefficients remained low, the approach can be further refined for high fidelity MTG-LI proxy data.

~~A second type of ground-based LLSs uses very high frequency (VHF). Recorded signals and related physical processes can differ from those recorded by VLF/LF LLSs.~~ VHF <u>LLSs are sensitive to lightning channel formation and leader processes, which occur multiple times during a single flash. Hence, VHF</u> LLSs typically feature high DE performances and three-dimensional (3D) ~~lightning channel mapping (Thomas et al., 2004). Their drawback is the limited range. These LLSs use direct line of sight signal~~ *mapping* of lightning channel propagation and spatial extent (Thomas et al., 2004). <u>VHF LLSs depend on direct line-of-sight</u> detection, and thus, the range suffers from the Earth's curvature and terrain shading effects. ~~LF LLSs can benefit from a reflection of the LF signal at the ionosphere.~~

Thomas et al. (2000) presented a case study of a storm in Oklahoma, USA, at local nighttime. The storm was observed by both the local Lightning Mapping Array (LMA) and TRMM-LIS. 108 of the 128 LMA lightning discharges were detected by TRMM-LIS and the LMA detected all TRMM-LIS flashes. The lightning missed by TRMM-LIS was mainly confined to low altitude discharges, i.e. below 7.0 km. Optical signals of lightning discharges that propagated via scattering to the upper part of the cloud were easily detected by TRMM-LIS.

Blakeslee et al. (2002) studied the São Paulo LMA (SP-LMA) dataset and its capability to serve ~~for~~ <u>as</u> GLM proxy data. TRMM-LIS events were in good agreement with the concurrent SP-LMA, ENTLN and LINET observations regarding latitude, longitude and timing. The records showed as expected more VHF (SP-LMA) sources than VLF pulses/strokes (ENTLN and LINET) per flash.

Due to TRMM satellite ~~orbits~~<u>orbit</u>, the comparisons of TRMM-LIS and ground-based LLSs records are restricted to tropical and subtropical regions between about 38 °N and 38 °S. As a result of its higher inclination orbit, ISS-LIS now allows the observation of extra-tropical thunderstorms to extend to 55 °N and S. The higher latitude storms might show different behaviors to their tropical and subtropical counterparts due to modified cloud vertical extent and forcing like the general wind field, average temperature and temperature gradients. Our study concentrates on the characteristics of lightning flashes over the northwestern (NW) Mediterranean Sea and should contribute to a better understanding of both European storms and European LLSs. This allows for the first time an intercomparison of LIS and European LLSs. Three LLSs operating in different spectral regions (Near-IR, VLF/LF, VHF) are compared: The satellite-based ISS-LIS being operational since March 2017, the French Meteorage VLF/LF LLS and the VHF SAETTA LMA on Corsica. The relative DE of ISS-LIS to Meteorage (and reverse) is analyzed, while SAETTA is used to verify and understand the results. Indeed, the spatially and temporally high resolution of SAETTA's measurements capture the structure and the life cycle of each lightning flash and gather additional information, i.e. discharge altitude, to assess more thoroughly ISS-LIS and Meteorage strengths and weaknesses. ~~Among~~ <u>Besides</u> the commonly investigated relative DEs, distances ~~, timing offsets~~and <u>and timing offsets, this work examines also</u> specific characteristics of matched ISS-LIS and Meteorage flashes~~are further examined. This work~~<u>. It</u> aims at providing the basis for mimicking optical, satellite-based lightning data from a VLF/LF LLS.

In section 2 ISS-LIS, Meteorage and SAETTA are introduced as well as the data processing, developed algorithms and the

125   investigation methodology. Results are presented in section 3. A brief summary and some discussion are given in section 4.

2   Instrumentation and Methodology

This paper aims at identifying the individual ~~characteristics in lightning detection~~ lightning detetection characteristics by the
satellite-based ISS-LIS, the VLF/LF Meteorage and the VHF SAETTA LLSs. ISS-LIS, installed on the International Space
Station in 2017, has been acquiring data since March 01, 2017. Our intercomparison of the LLSs covers the period from March

130   01, 2017 until March 20, 2018. The region was limited to 40.5°N to 44.0°N and 7.0°E to 11.0°E around the island of Corsica
in the NW Mediterranean Sea. Figure 1 shows the domain with accumulated data of one overpass (a), an infrared (IR) satellite
picture (b) and the example of one flash recorded by ISS-LIS, Meteorage and SAETTA (c). The three instruments are introduced
within this section. In total, ISS-LIS field of view (FOV) intersected the region of interest 851 times during the study period,
with 26 of the overpasses exhibiting lightning activity. In ~~our paper~~this work, all times are given in Coordinated Universal Time

135   (UTC). Altitudes are defined above sea level (ASL). Distances are calculated using Vincenty's formulae (Vincenty, 1975) based
on the WGS 84 reference ellipsoid which are more accurate on Earth than for example great circle distances (assumes the Earth
as oblate sphere rather than a sphere). The term *detection efficiency* (DE) means in the following the DE for flashes, not event
or pulse/stroke DE.

2.1   ISS-LIS

140   The ISS operates in Low Earth Orbit (LEO) and overpasses one region on the surface up to three times a day (up to two times in
the tropics). Lightning observation of a specific point lasts up to 90 seconds per overpass due to the ISS orbit characteristics and
the LIS FOV of approximately 655 x 655 km^2. The optical lightning detection is performed at a wavelength of 777.4 nm at the
atomic oxygen line. ISS-LIS observes both IC and CG discharges but cannot distinguish the lightning type. ISS-LIS captures
an image of the Earth every ~~two milliseconds~~2 ms referred to as a frame. The LIS focal plane consists of a 128 x 128 pixel

145   Charge Couple Device (CCD) that is read out every 2 ms. The pixel FOV ranges between 4.5 km (nadir) and 6.2 km at the edges
(Dennis Buechler, personal communication 2019). Blakeslee and Koshak (2016) apply a four-step filtering approach, involving
spatial, spectral, temporal and background subtraction filter, to identify pixels with lightning activity. This is required to detect
the lightning during daytime when the sunlight reflected off the cloud tops otherwise overwhelms and masks the lightning
signal (i.e., it is daytime lightning detection that drives the design of space-based lightning detectors such as LIS and the new

150   MTG-LI). An illuminated pixel that breaks a predefined threshold in a given 2 ms frame is identified as an event. Events
define the smallest units of the optical signals in the ISS-LIS data set. Their latitude and longitude correspond to the pixel
center. A group is the next unit of ISS-LIS data. An ISS-LIS group contains one or more events occurring within the same
time frame and in adjacent pixels of the ISS-LIS image (Christian et al., 2000). Next, groups are organized into flashes, so
that a flash can consist of one or multiple groups. ~~Groups occurring within~~ A Weighted Euclidean Distance (WED) employs

155   spatial and temporal clustering with 330 ms and 5.5 km~~belong to the same flash~~, respectively, to merge groups in flashes

(Mach et al., 2007). The locations of groups and flashes are defined by the radiance weighted average positions of their events and groups, respectively. Finally, an area contains all flashes with distances of less than 16.5 km to each other. The National Aeronautics and Space Administration (NASA) provides the ISS-LIS in different postprocessing levels. In the latest available version, P0.2, the quality control is already close to its (expected) final stage, but the data may contain some undetected minor errors (Blakeslee et al., 2017). The main difference will concern the detection efficiency. The fully validated flash density should not alter more than 5.0 % to 10.0 % from version P0.2 (R. Blakeslee, personal communication 2018). LIS data comprises the 2 ms scientific data, e.g. time, latitude, longitude and ~~radiance~~ optical amplitude count of events and instrument, platform or external errors to verify the data quality, and housekeeping data. The available ISS-LIS P0.2 version data have not yet included the (background-)calibrated radiance. The strength of the optical signal is defined by the raw amplitude count. It depends somewhat on the background value, but, in general, the radiance increases with the raw count (personal communication Dennis E. Buechler, NASA MSFC, August 2019). The housekeeping data, received every second, contains among others LIS' viewtime with information about the FOV at a time. It is provided on a 0.5° x 0.5° grid. ISS-LIS viewtime is fundamental for the intercomparison to continious observations at the ground.

The original ISS-LIS data contains times in the International Atomic Time with reference to 01 Jan. 1993 (TAI93) format. For the intercomparison of the LLSs, times are converted to UTC while taking the missing leap seconds into account. The ISS-LIS times include a time-of-flight (TOF) correction accounting for the time photons need to travel from the ~~lightning discharge on Earth~~ optical source at cloud-top to the satellite.

2.2 Meteorage

The Meteorage LF LLS uses Vaisala LS7002 sensors (Vaisala, 2013) at a frequency between 1 kHz and 350 kHz. It includes 21 ground sensors across France and contributes to the European Cooperation for Lightning Detection (EUCLID). EUCLID comprises lightning sensors all over Europe and helps to improve the performance of national LLSs (Schulz et al., 2016). The LS7002 sensors measure the signals related to CG strokes and IC pulses, thus the total lightning. Vaisala claims a CG DE of 95 % and a DE for IC of 50 %. Pédeboy et al. (2018a) stated that indeed 97 % of the CG flashes and 56 % of the IC flashes were detected by Meteorage (68.3 % overall DE relative to LMA flashes). The theoretical median location accuracy approximates 250 m and improves inside the network to about 150 m. Pédeboy et al. (2018a) found reduced median location accuracy for IC flashes of 1.64 km. Time synchronization applies a GPS receiver with an accuracy of 50 ns to UTC. The lightning location needs at least ~~four sensors~~ two sensors (each provides a time and an angle of arrival) by applying combined magnetic direction finding and time-of-arrival techniques. Lightning can be detected in a distance up to 1500 km from a sensor. In practice, the use of ionospheric reflection is avoided, hence, limiting the sensor range to about 625 km. It ensures the ground plane wave front of the signal is measured rather than a reflected wave of the lightning related signal. Our study makes use of the Meteorage lightning pulse/stroke data. For each pulse/stroke, the occurrence time, latitude, longitude, the amplitude with polarity and the type (IC/CG) are provided. Meteorage ~~data are then~~ observes lightning continuously within its range. In the intercomparison with ISS-LIS, data are disregarded if observation space or time do not fit the corresponding ISS-LIS viewtimes.

2.3 SAETTA (Suivi de l'Activité Electrique Tridimensionnelle Totale de l'Atmosphère)

190 The LMA technology was developed by New Mexico Tech (Rison et al., 1999). The SAETTA LMA operates in the 60-66 MHz VHF band, with an 80 μs analysis window (Coquillat et al., 2014), and consists of 12 LMA stations distributed over the island of Corsica. The distance between the network's northernmost and southernmost (westernmost and easternmost) stations approximates 180 km (70 km). The station altitude ranges from 3.8 mASL to 1950.2 mASL. SAETTA maps the total lightning activity. ~~Fast CG discharges travelling between the cloud and ground in time frames shorter than 80 might be missed.~~

195 A minimum of six stations is needed to capture a lightning source in 3D. Redundant information from more stations improves the location accuracy and consequently decreases the chance of mislocation and possible noise (e.g. single VHF sources in Figure 1(c)). As a drawback, less VHF sources and flashes are detected simultaneously by more than six stations. Aiming at a high flash DE, coincident signals at six stations are sufficient for the LMA data in this study.

SAETTA data include the time, latitude, longitude, altitude, amplitude of each lightning source. Lightning location reaches 200 up to a radius of 350 km from the center of the network. The VHF LLS depends on the direct line-of-sight to a lightning discharge. The altitude of the lowest detectable VHF source increases with the distance to the LMA due to Earth curvature. For example, sources in 100 (200) km distance to a station at sea level must be at least 0.8 (3.1) km in altitude to be visible to that station. Formula (6) of Koshak et al. (2018) is applied here. Their study also investigates effects of the LMA network geometry mainly on the altitude errors. The SAETTA location uncertainty increases with the distance to the network center.

205 According to the theoretical model of Thomas et al. (2004), the radial azimuthal and altitude errors are, at best for VHF sources at 10 km altitude, 15 m, 8 m and 40 m, respectively, within 50 km from the center of the network (Coquillat et al., 2019). These theoretical errors reach about 300 m, 20 m and 400 m, respectively, at the borders of the present study domain (Figure 1(a)). Hence, SAETTA location errors are in the same order of magnitude as those of Meteorage CG location while the LMA should capture lightning in more detail than the LF LLS.

210 SAETTA data are employed for locations and times of coincident ISS-LIS or Meteorage observations. Therefore, they are analyzed in space and time regarding the detected ISS-LIS and Meteorage lightning activity. A combined space-time filter identifies SAETTA sources within 0.2° (both latitude and longitude) and (simultaneously) 0.3 s of corresponding ISS-LIS events and Meteorage pulses/strokes. The filtering per (ISS-LIS or Meteorage) flash allows for analyzing the concurrent VHF measurements with e.g. altitude information. Furthermore, SAETTA data are not used to exclude any ISS-LIS or Meteorage 215 observations and neither ISS-LIS nor Meteorage data are confined to any SAETTA data condition. They are, however, used to verify the applied data processing approaches, i.e. grouping elements (events, pulses/strokes) to flashes and the analysis of possible false alarms within the lightning detection of ISS-LIS and Meteorage. The maximum altitude of SAETTA sources is bounded at 15.0 kmASL and the maximum $reduced\ \chi^2$, which defines a measure for the overall uncertainty of the time-of-arrival based system (Thomas et al., 2004), is set to 0.5.

2.4 Flash - Grouping algorithm

The NASA LIS flash clustering algorithm distinguishes events, groups and flashes (section 2.1). ~~The distance between two events of one flash (Christian et al., 2000) can be up to 14.3~~It makes use of a WED with maximum difference of 5.5 km ~~at the edge of the FOV (up to 11.9~~in space and 330 ~~nadir). The time constraint, dt_{merge}, confines the maximum time between two consecutive groups~~ms in time. The WED analyzes group centroids and not the events in the group to find if two groups are considered part of the same flash~~as 330~~. One flash cannot last longer than 2.0 s (Christian et al., 2000). An analysis of the (P0.2~~NASA flash sorting~~) NASA LIS flash clustering algorithm revealed that it tends to separate flashes when compared to concurrent SAETTA observations. Similar results were observed by Defer et al. (2005). Consequently, a new algorithm is developed to merge the ISS-LIS events to flashes. It has the additional advantage of treating both ISS-LIS events and Meteorage pulses/strokes. The ~~fundamental~~ lightning elements sensed by each LLS, that are the smallest available lightning signals (events and pulses/strokes), are merged into flashes. More explicitly, an event of ISS-LIS (pulse/stroke of Meteorage) should belong to exactly one flash and a flash is defined as a ~~group~~ collection of events (pulses/strokes). Flash characteristics are derived from the underlying element characteristics, e.g. the positions of its elements are used instead of the mean flash location. This study makes use of the ~~fundamental~~ elementary ISS-LIS event data as provided by the NASA prior to any data merging. It is accepted that ISS-LIS events do not have a direct representation in the Meteorage-like data. Former studies claimed that LIS groups roughly correspond to the physical processes detected by VLF/LF LLSs (e.g., Bitzer et al., 2016; Höller and Betz, 2010). Nevertheless, those studies found significantly more groups than pulses/strokes within the same region and time period. Bitzer et al. (2016) found for the number of TRMM-LIS groups to ENTLN pulses/strokes a factor of about 28.4 globally and even 3.7 in North America in 2013. Höller and Betz (2010) analyzed 6.7 groups per pulse/stroke on average. Due to those results, ~~it is questionable whether LIS groups really correspond to (V)LF pulses~~LIS optical groups emerge from both discharge processes measured by VLF/LF sensors but also processes lacking significant VLF/~~strokes~~LF radiation. In addition, the detected lightning sources of the applied VHF LLS comply more with the LIS events than the groups. Using events rather than group centroids improve in particular the finding of the coincident LMA data. The analysis of flash extents profits from the use of events in that the extent of an ISS-LIS flash corresponds to the full illuminated area rather than the ISS-LIS group centroid locations. The representation of the flash extent (density) will influence the future assimilation of lightning data in NWP models. A statistical analysis of (ISS-LIS) events and LF strokes/pulses will also be of interest for creating a proxy optical data set, e.g. for MTG-LI, derived from LF data.

Our *grouping algorithm* analyzes the ~~element~~ elements (events or pulses/strokes) and groups ~~the elements~~ them based on their relative location and time of occurrence to each other. First, the spatial and temporal constraints, ds_{merge} and dt_{merge}, for elements within one flash must be determined. Then, a combined space-time test merges the elements into flashes. It starts with the first available element (in the data of one LLS) and identifies all elements (of the same LLS' data) within the range of the constraints. Thereby, an element can only belong to the same flash if both the distance to ~~the initial element~~ any element of the flash is less than ds_{merge} and the time difference (to the same element) is shorter than dt_{merge}. All elements identified for a flash (~~and~~ including the initial element) are classified as *used*. For each used element within a flash, the test is repeated until no

unused element can be added to the flash. This step allows for considering propagating flashes and the potentially increasing

255 extent and duration of a flash ~~when~~ while adding new elements. The algorithm continues until all elements are classified as used. ~~In general, our~~ Our algorithm does not limit the duration of a flash. The number of elements per flash remains also free to the algorithm.

The algorithm verification includes a sensitivity study for ds_{merge} and dt_{merge} (Figure 2) and a comparison to NASA's algorithm and concurrent SAETTA observations (Figure 3).

260 Figure 2 gives the number of flashes analyzed from all observations of the approximately one-year-period by using different ds_{merge} (panels a and c) and dt_{merge} (b, d) for ISS-LIS (a, b) and Meteorage (c, d). In general, as expected, smaller ds_{merge} and dt_{merge} increase the flash numbers because less individual elements are part of a given flash and thus more flashes exist for the same elements. LIS flash numbers range from 236 to 4567 for the ds_{merge} (dt_{merge}) between 50.0 km and 1.0 km (1.0 s and 0.1 s), respectively. For the same constraints, Meteorage flash numbers vary from 340 to 1720 flashes.

265 The ISS-LIS flash number decreases rapidly for ds_{merge} between 0 km and 10 km (Figure 2(a)). The rapid change depends on the pixel size within the ISS-LIS image. Hence, it is expected that events of one flash are partitioned within the same frame if the ds_{merge} becomes smaller than the ISS-LIS image pixel size. ISS-LIS flash numbers remain constant for ds_{merge} greater than 15 km for all tested dt_{merge}. 0.3 s balances the need of consistency and the wish for a strict dt_{merge} (Figure 2(b)). The resulting flashes are verified against concurrent 3D SAETTA sources which supported the choice of our constraints and the

270 identification of resulting flashes. ~~Both chosen constraints for ISS-LIS flashes,~~ The spatial constraint (ds_{merge} of 15 km~~for~~ ~~ds_{merge} and 0.3 for~~) refers directly to event locations and not to group centroids (as for NASA's algorithm). The chosen time constraint for ISS-LIS flashes (dt_{merge} ~~, are~~ of 300 ms) is similar to the P0.2 NASA flash ~~sorting algorithm~~ clustering algorithm (330 ms).

The same algorithm is applied to group the Meteorage pulses/strokes into flashes. It needs, however, modified constraints

275 ds_{merge} and dt_{merge} since physical processes producing Meteorage pulses/strokes do not always ~~represent the same physical processes as~~ correspond to ISS-LIS events and occur with significantly lower counts. Meteorage pulses/strokes do not ~~always cover the~~ cover the full structure and duration of a lightning flash. Figure 2(b, d) is analyzed for Meteorage flash numbers as demonstrated for ISS-LIS flash numbers in (a, c). To find constant constraints suitable for various situations (e.g. vertical cloud structures, severity of a storm, flash rate), resulting flashes for different constraints are verified manually using SAETTA

280 observations. ~~Conclusively~~ Ultimately, Meteorage pulses/strokes belong to the same flash if they are detected within 20 km and 0.4 s. Due to the limited number of pulses/strokes, Meteorage ds_{merge} and dt_{merge} are coarser than the ISS-LIS merging constraints. Our constraints (20 km, 0.4 s) are consistent with the Meteorage-own flash grouping algorithm using a separation distance of less than 10 km for subsequent CG strokes and 20 km if IC pulses are involved. The delay between supsequent discharges of the same flash must be smaller than 0.5 s in Meteorage's algorithm. Höller and Betz (2010) provided a clustering

285 of LINET VLF/LF pulses/strokes to flash scale with 10 km and 1.0 s in space and time, respectively. Hence, their merging constraints for a flash are finer in space but coarser in time.

~~One must mention, though, that the~~ The determination of ds_{merge} and dt_{merge} does not ensure a perfect arrangement of the elements in flashes though. The objective is to find constraints leading to statistical representations of flashes in the ISS-LIS

and Meteorage data. Therefore, all identified flashes are double-checked against concurrent 3D SAETTA observations. Even if it is sometimes challenging to separate the flashes in the SAETTA data, the detailed VHF mapping helps to understand the processes leading to the identification of the ISS-LIS and Meteorage flashes. The SAETTA data ~~are also~~ should also be used to find possible false alarms in the ISS-LIS and Meteorage data.

Figure 3 demonstrates the behavior of NASA's flash merging algorithm and our developed algorithm for one example. It shows a short time period (6.5 s) during one ISS overpass on September 10, 2017. In Figure 3(a), there is the map of flash locations from the P0.2 NASA flash merging algorithm and our developed algorithm as well as concurrent VHF SAETTA sources. The ISS-LIS events (not plotted) coincide in general well with the SAETTA observations in both location and time. The mapped observations are presented in latitude, longitude and altitude time series in Figure 3(b). 20 flashes from NASA's algorithm are confronted with 11 flashes from our developed algorithm for the same ISS-LIS events. NASA's merging algorithm somehow splits some flashes, e.g. the flash between 5.3 s and 6.2 s where NASA's algorithm identifies 5 flashes. Our algorithm finds a single flash for that period and the concurrent SAETTA observations support this result.

Our developed algorithm was in addition tested versus GLM flash-scale data ~~using~~. The GLM algorithm uses the distance between events and not (as ISS-LIS) between group centroids in order to merge events/groups to flashes. ~~Time~~ WED time and spacing for events/groups of one GLM flash are 330 ms and 16.5 km, respectively (Goodman et al., 2013). The GLM flash-scale data agree very well with the flashes identified by our algorithm ~~from~~ utilizing the underlying GLM events.

2.5 Flash - Matching algorithm

ISS-LIS and Meteorage detect ~~flashes~~ lightning in a different way. It was described how the different signals can be merged into a common entity, namely a flash. ~~Our~~ The intercomparison of LLSs uses flash-scale to find concurrent observations. ~~Our algorithm scans the individual~~ Individual flashes of both LLSs ~~and sorts them~~ are sorted into one of the four following categories: *LIS detected by both* (i.e. an ISS-LIS flash has a coincident Meteorage flash), *LIS-only* (i.e. no coincident Meteorage observations), *Meteorage detected by both* (i.e. a Meteorage flash with concurrent ISS-LIS events) or *Meteorage-only* (i.e. ISS-LIS does not detect the flash). Matching criteria in space (ds_{match}) and time (dt_{match}) are specified. The criteria ds_{match} and dt_{match} do not address the flash mean position and time, respectively, but the single events or pulses/strokes within a flash. ~~Flashes detected by one LLS are characterized~~ Two flashes observed by different LLSs are defined as *matched* ~~to a given flash of the second LLS if they occur within the range of~~ if at least two elements (one per flash) meet both ds_{match} and dt_{match} ~~to the elements of the given flash. As a given flash detected by a given~~. A given flash of the reference LLS does not necessarily correspond to exactly one matched flash, ~~the two categories *LIS detected by both* and *Meteorage detected by both* are expected to have different counts.~~ It is also possible that a flash ~~fulfills~~ meets the matching criteria of more than one given flash ~~and thus~~ (and is collocated to more than one flash.~~)~~. Hence, the two categories *LIS detected by both* and *Meteorage detected by both* are expected to have different counts.

The detailed analysis of distances and timing offsets between matched flashes refines the ~~results~~ matching algorithm further: The ~~closest flash detected by the second LLS is identified for a given flash of the given LLS . Our algorithm analyses the underlying elements of each flash. The algorithm starts with~~ number of cases where one flash is matched to multiple flashes

10

of the second LLS should be reduced. Therefore, the refined algorithm initiates with finer matching criteria, i.e. one percent
of both ds_{match} and dt_{match}~~seeking for any~~. It seeks for one element detected by the second LLS ~~around~~ that meets both
the finer criteria to any element of the given flash. ~~The~~ Only if no match is found, the allowed distance and time difference
increase ~~iteratively~~ by one percent of ds_{match} and dt_{match}, respectively~~, for a given flash (and all its elements) until~~. The
process repeats iteratively until either a match is found ~~.If~~ or the allowed distance (timing offset) exceeds the original ds_{match}
(dt_{match}). In the latter case, the algorithm stops and the flash is ~~called~~ labelled *unmatched* (Note: The refined analysis is
performed for matched flashes only, however, the algorithm can treat the unmatched flashes, too.). One or more matches for
the given flash are still possible because of the discrete increments from one iteration to the following. There might also be
flashes within ~~similar distance and similar~~ equal distance and equal time offset to the given flash.

The criteria ds_{match} and dt_{match} are determined through a sensitivity study of the relative DEs of ISS-LIS and Meteorage
(Figure 4). A spatial criterion lower than 10.0 km reduces the relative DE of both ISS-LIS and Meteorage rapidly (Figure 4~~(afor~~
~~; a:~~ ISS-LIS, c~~for Meteorage))~~: Meteorage). In general ISS-LIS' relative DE is more sensitive to both ds_{match} and dt_{match}
than Meteorage relative DE. This result is triggered by the low number of Meteorage pulses/strokes (compared to the number
of ISS-LIS events) hampering effectively the finding of suitable elements, i.e. pulses/strokes, for a collocation. ISS-LIS relative
DE decreases within the entire range of investigated times dt_{match}. The most sensitive behavior occurs for dt_{match} up to 1.5 s
(Figure 4(b)). Meteorage appears to be sensitive to dt_{match} ~~onl yup~~ only up to 0.5 s (Figure 4(d)). Despite the differences in
sensitivity to the criteria between ISS-LIS and Meteorage, it is aimed at using the same ds_{match} and dt_{match} for both LLSs.
Finally, ds_{match} of 20 km and dt_{match} of 1.0 s are chosen to balance the individual sensitivities of the LLSs to the criteria.
They allow to identify matches if, for example, ISS-LIS detects primary IC discharges of a flash and Meteorage only detects
a CG stroke occuring during the final stage of the same flash. Our criteria are relatively coarse compared to some former
studies (section 1). Höller and Betz (2010) applied the same dt_{match} but an even coarser ds_{match} (i.e. 30 km) to match LINET
VLF/LF flashes and TRMM-LIS flashes. Further investigation of the matched flashes, e.g. the distributions of the distances
and timing offsets, will demonstrate to which extent matches rely on the fairly coarse criteria.

3 Results

The different LLSs detect flashes in different ways and with distinct characteristics. In this section, flash observations are
compared and analyzed. As an example, the ISS overpass with the corresponding observations of ISS-LIS, Meteorage and
SAETTA in Figure 1 comprises (almost) the entire study region. It lasted 169 seconds, from FOV entering to leaving the
region. The effective viewtime per 0.5° x 0.5° grid box is indicated in grayscale in Figure 1(a). Wide parts of the domain have
been seen for at least 60 seconds. Figure 1(b) shows additionally an IR satellite image indicating the cloud tops. The example
of a single flash observed by all three LLSs during this overpass is given in Figure 1(c). SAETTA captures the most detail of
the flash structure and there are significantly more ISS-LIS events than Meteorage pulses/strokes. All but the first Meteorage
signals indicate an IC pulse. Since the first stroke is of type CG, the entire flash is characterized as CG-flash.

355 First, relative DEs of ISS-LIS and Meteorage are elucidated. The comparisons of matched flash locating and timing differences are discussed and finally characteristics of flashes, with special interest in differences between matched and unmatched flashes, are analyzed.

3.1 Detection comparison

This DE analysis is realized on the flash scale. Flashes were preliminarily identified by our in-house algorithm, which
360 merges ISS-LIS events and Meteorage pulses/strokes according to their locations and times of occurrence. Further investigations break the flash scale down into events and pulses/strokes, e.g. for the flash characteristics.

The period of observations spans from March 01, 2017 to March 20, 2018. In total, 330 ISS-LIS flashes and 569 Meteorage flashes are identified by our algorithm.

Besides the DE, the Probability of False Alarm (POFA) characterizes the quality of detection. Quantifying the POFA needs
365 knowledge about the truth, that is the real number of flashes. SAETTA could provide a reference value to quantify the POFA, however, not all stations have operated continuously for the entire study period. Signals from at least six stations are needed to reconstruct and locate a discharge signal. 31 of 330 (89 of 569) ISS-LIS (Meteorage) flashes were not detected by SAETTA. SAETTA's detection efficiency and accuracy decreases also with the distance to the network's center. A vast majority of more than 90 % of all flashes occurred outside a distance of 100 km from SAETTA's center. There was evidence of VHF activity
370 for the vast majority (95.0 %) of the missed flashes. In 59.2 % of the cases, at least 1 and less than 6 stations recorded signals. SAETTA lightning was observed just outside our matching criteria (0.2 °, 0.3 s) for 28.3 % of the missed flashes. The result indicates that the POFA is low for both ISS-LIS and Meteorage (although it cannot be quantified).

Only 3 (12) of the ISS-LIS (Meteorage) flashes missed by SAETTA are located within 100 km of SAETTA's center. Due to the low total flash number within this close domain to SAETTA, a statistical analysis is ambiguous. 13 of the 15 missed flashes
375 near SAETTA were missed due to station downtimes or filter criteria (spatial, temporal, reduced chi² [section 2.3]). Pédeboy et al. (2018b) reported Meteorage flashes missed by SAETTA with (absolute) peak currents exceeding 100 kA. Two of the twelve missed Meteorage flashes close to SAETTA exhibit an (absolute) current above 100 kA. SAETTA data can in fact not serve for the desired true flash numbers, mainly due to station downtimes, and the POFA of Meteorage and ISS-LIS cannot be calculated.

380 60.7 % (54) of the Meteorage flashes without concurrent SAETTA sources contain one pulse/stroke only. Those flashes with only one pulse/stroke (or one event in case of ISS-LIS) are referred to as *single element flashes*. Missing SAETTA observations for a single element flash might be indicative of a locating and timing problem of the ISS-LIS event or Meteorage pulse/stroke (possible false alarm). The DE analysis distinguishes results for the complete dataset and excluded single element flashes. 316 ISS-LIS and 367 Meteorage flashes remain after excluding
385 the single element flashes. Thus, the ISS-LIS (Meteorage) flash number is reduced by 14 (202) flashes compared to the overall count. ISS-LIS single element flashes are rare while there is a significant amount of Meteorage single element flashes. The result is related to the differences in optical and LF lightning detection. The entire data set contains 16,881 ISS-LIS events and 2,144 Meteorage pulses/strokes (487 CG, 1657 IC). 15,578 events (92 %) are distributed over the ISS-LIS flashes with match.

Table 1. Relative detection efficiencies (DEs) of Meteorage and ISS-LIS. The values in parentheses give the relative DEs for flashes with at least two elements. The flash numbers (100 %) to calculate the DEs are indicated. Note: ISS-LIS (Meteorage) DE uses Meteorage (ISS-LIS) flash numbers.

	Overall	Daytime	Nighttime	IC-flash	CG-flash
ISS-LIS DE [%]	57.3 (62.4)	53.9 (60.8)	58.7 (63.0)	59.3 (68.8)	53.5 (55.4)
Meteorage flash number	569 (367)	167 (102)	402 (265)	369 (192)	200 (175)
Meteorage DE [%]	83.3 (83.9)	80.0 (80.2)	84.8 (85.5)	-	-
ISS-LIS flash number	330 (316)	100 (96)	230 (220)	-	-

For Meteorage, 1,439 pulses/strokes (271 CG, 1168 IC) constitute the flashes with matches (67 %). Hence, 55.6 % (70.5 %) of the CG strokes (IC pulses) belong to flashes with coincident ground-space-detection. Despite coarser ds_{merge} and dt_{merge} for a Meteorage flash than an ISS-LIS flash, Meteorage observed 239 flashes more than ISS-LIS within similar regions and time frames.

Figure 5(a) presents a histogram of the total flash detection counts within the four categories introduced in section 2.5. The number of single element flashes is marked. Additionally, it includes a map of the locations of the flashes within each category. Figure 5(b) maps the flashes as 2D-histogram on a 0.1° x 0.1° grid. Flashes are detected all over the study domain for both ISS-LIS and Meteorage without any apparent pattern even if the number of flashes is not sufficient to be statistically representative.

Table 1 summarizes all relative DEs. Daytime covers the time from 05:00 UTC to 17:00 UTC. Nighttime flashes are defined between 17:00 UTC and 05:00 UTC.

ISS-LIS was able to detect 326 of the 569 recorded Meteorage flashes from March 01, 2017 to March 20, 2018, a relative DE of 57.3 %. If the notable number of Meteorage single element flashes is neglected, ISS-LIS detected 229 of the remaining 367 Meteorage flashes (62.4 %). ISS-LIS shows a low relative DE of less than 54 % for daytime flashes. 58.7 % of the Meteorage nighttime flashes are detected by ISS-LIS. In particular the nighttime relative DE cannot reach literature expectations of over 90 % for LIS (Boccippio et al., 2002). ISS-LIS' relative DE does significantly depend on the Meteorage flash type. A flash with at least one CG stroke, referred to as CG-flash, is detected in only 53.5 % of the cases while a pure IC-flash is detected with 59.3 % relative DE. ISS-LIS could detect 68.8 % of the occurring Meteorage IC flashes with at least two pulses. If flashes with at least two pulses/strokes are considered, the relative DE of IC flashes surpasses that of CG-flashes by almost 14 % and increases compared to the total IC flash relative DE by 9.5 %. Hence, especially CG-flashes and single pulse IC-flashes decrease the total DE of ISS-LIS. All relative DEs use ds_{match} of 20 km and dt_{match} of 1.0 s. Finer criteria would further decrease the relative DE of ISS-LIS (higher sensitivity to the criteria than Meteorage).

Out of the total 330 ISS-LIS flashes, Meteorage detected 275 (83.3 %). The DE of Meteorage relative to ISS-LIS flashes with at least two events equals 83.9 % (265 of 316 flashes). The relative DE of the VLF/LF Meteorage LLS appears to be significantly higher than in former studies (section 1) using LF LLSs and TRMM-LIS. It is assumed that the ISS-LIS detection efficiency is similar to that of TRMM-LIS in general (personal communication, R. Blakeslee) and thus Meteorage provides a

415　high quality LF LLS. Moreover, especially Meteorage detection efficiency appears to be quite resistant to changes of ds_{match} and dt_{match}. For example, halving both criteria (10 km in space, 0.5 s in time) results in a relative detection efficiency of about 78 %. More details about the sensitivity to the matching criteria can be found in Section 2.5.

　　Meteorage detected 80.0 % of the 100 ISS-LIS daytime flashes. Its relative DE reaches 84.8 % for 230 ISS-LIS nighttime flashes. The relative DE depends on both the performance of the LLS itself but also the performance and locating accuracy

420　relative to the reference LLS. As ISS-LIS detects flashes optically, the influence of different lighting on ISS-LIS daytime and nighttime accuracy is investigated as part of the following section 3.2.

3.2　Distances and timing offsets between collocated flashes

In this section, the matched ISS-LIS and Meteorage flashes are studied regarding their relative location and time of occurrence. For each element of a flash detected by one LLS the closest (in time or in space, not a combined filter here) element of the

425　matched flash(es) accounts for the statistic. One element can be closest to multiple elements of the second LLS. The entirety of elements of flashes with matches is analyzed statistically. Figure 6 presents the results for distances (a) and timing offsets (b) between events and pulses/strokes.

　　Figure 6(a) shows histograms of the distance between a given ISS-LIS event (Meteorage pulse/stroke) and the closest pulse/stroke (event) of a matched flash. The distribution given an ISS-LIS event peaks primarily between 2.50 km and 3.00 km

430　and secondarily at about 4.50 km with a median (mean) of 4.74 km (5.68 km). The distribution given a Meteorage pulse/stroke has a broad maximum from 0.75 km to 2.75 km with a median (mean) of 2.31 km (3.60 km). The Meteorage pulse/stroke distance distribution features a more pronounced (if wider) peak for less distance than the distribution given an ISS-LIS event. This is due to the calculation method and the numbers of available events and pulses/strokes. The higher number of (and smaller distance between) ISS-LIS events allows in general for finding a closer event to a given Meteorage pulse/stroke than vice versa.

435　The cumulative distribution functions (CDFs) within the plotted interval (Figure 6(a)(ii)) show that the distance distribution given an ISS-LIS events has a larger tail than the distribution given a Meteorage pulse/stroke. The 60th percentile is found at approximately 5.5 km and 2.6 km for a given ISS-LIS event and Meteorage pulse/stroke, respectively. Both Meteorage IC pulses and CG strokes exhibit similar distributions to the overall Meteorage pulses/strokes (also in the CDFs, Figure 6(a)(ii)) with the peak between 0.75 km and 2.75 km. The median (mean) distance for IC pulses and CG strokes and their match equals

440　2.36 km (3.63 km) and 2.22 km (3.51 km), respectively. Hence, CG strokes feature a slightly lower distance to matched events than IC pulses.

　　Distances are calculated between the closest ISS-LIS events (not groups or flashes as in former studies) and LF pulses/strokes. ~~The group-pulse/stroke distances should be equal or greater than event-pulse/stroke distances as events~~ Events provide a finer resolution of the lightning discharge than groups while group centroids use radiance-weighted event locations and

445　are interpolated to sub-pixel locations. Bitzer et al. (2016), who used TRMM-LIS groups and ENTLN pulses, found for both conditional distributions median (mean) differences in location between 7.0 km and 7.2 km (7.6 km and 7.9 km) in North America. Those values are in accordance with the distances observed by Rudlosky et al. (2017) between TRMM-LIS flashes and GLD360 strokes. ~~All their results show a similar order of magnitude to our findings~~ Zhang et al. (2019) thoroughly

investigated TRMM-LIS performance and found additional latitudinal location offset resulting from TRMM yaw maneuvers.

450   Their analysis suggests a correction for TRMM-LIS group locations that improves the TRMM-LIS group distance to NLDN pulses/strokes for summers of 2012 and 2013. The distance distribution peaks for 1-2 km after the correction (5-6 km before the correction). The distribution peak distances before the correction were similar to those reported by Bitzer et al. (2016) and Rudlosky et al. (2017). Hence, it is assumed that distances between TRMM-LIS groups and VLF/LF LLS pulses/strokes are similar or slightly smaller than distances comparing ISS-LIS events and Meteorage LF pulses/strokes ~~.~~ in this work. It should

455   be mentioned that TRMM-LIS pixel size is slightly smaller than that of ISS-LIS, i.e. 4.3 (3.7) km nadir after (before) TRMM boost versus 4.5 km nadir.

The optical ISS-LIS sensor might be affected by different lighting. Therefore, the accuracy of ISS-LIS flashes relative to ground-based LLSs is explicitly investigated during day and night (not shown as Figure). Daytime flash distances are concentrated mainly between 2.0 km and 5.0 km and the distribution peaks at about 3.5 km. The ISS-LIS nighttime flash

460   distribution peaks at about 5.5 km distance to matched Meteorage flashes. Given an ISS-LIS flash, the CDF distribution also rises faster for daytime than for nighttime flash distances. Hence, distances between coincident flashes are in fact smaller during daytime than during nighttime. The comparison of ISS-LIS flashes to SAETTA reveals a small difference of up to $0.05°$ latitude and longitude during both day- and nighttime. ISS-LIS flashes tend to occur slightly south and west of the corresponding SAETTA observations. The small locating difference, considering ds_{match} of 20 km and ISS-LIS spatial resolution of 4.5 km

465   (nadir), does not significantly influence our results. In particular, ISS-LIS maintains its locating accuracy during daytime and during nighttime.

The timing offset subtracts the time of the matched element from the time of the given element. It yields positive and negative values according to which element occurred first, with a positive value indicating that the given element occurred later than its match. Again, the two conditions *given an ISS-LIS event* and *given a Meteorage pulse/stroke* are applied. The

470   resulting distribution (Figure 6(b)) peaks between ~~-0.5~~ −0.5 ms and 0.5 ms for a given ISS-LIS event and between ~~-1.0~~ −1.0 ms and 1.0 ms for a given Meteorage pulse/stroke. The ~~tails of the distribution~~ distribution tails, with absolute timing offset longer than 10 ms and up to 1.0 s, are not plotted. They are larger for a given ISS-LIS event than for a given Meteorage pulse/stroke. It is observed that Meteorage pulses/strokes often do not cover the entire duration of a flash. ISS-LIS events reflect the actual flash duration (reference to concurrent SAETTA sources) better than the Meteorage pulses/strokes. Hence,

475   given an ISS-LIS event and looking for a matched Meteorage pulse/stroke, the number of available pulses/strokes is often limited. Several events can have the same closest pulse/stroke even if the events occurred in different time frames. It increases the probability of larger timing offsets especially for a given ISS-LIS event compared to a given Meteorage pulse/stroke. The CDFs (Figure 6(b)(ii)) reveal that about 20 % (5 %) of the ISS-LIS events (Meteorage pulses/strokes) shown here exhibit timing offsets of more than positive 2.5 ms. About 20 % (10 %) of ISS-LIS events (Meteorage pulses/strokes) have values lower than

480   -2.5 ms). In the overall distribution (not shown), time ~~offset~~ offsets exceed positive 10.0 ms for 43 % (22 %) of ISS-LIS events (Meteorage pulses/strokes)~~and are more negative than -10.0.~~ Negative time offsets exceed $−10.0$ ms for 25 % (22 %) of ISS-LIS events (Meteorage pulses/strokes). The distribution given an ISS-LIS event is slightly skewed towards positve time offsets (given ISS-LIS event occurred later than its best match stroke/pulse). The overall median (mean) values yield 2.36 ms

(54.60 ms) and ~~-0.00~~ −0.00 ms (2.70 ms) given an ISS-LIS event and Meteorage pulse/stroke, respectively. The ~~large mean value~~ mean for a given ISS-LIS event is an artifact of ~~skewed distribution~~ the skewed distribution (also in the tails). Considering the ISS-LIS integration frame time of 2.0 ms, the remaining average statistics are close to the temporal accuracy of ISS-LIS. Both conditional distributions given ISS-LIS and given Meteorage show an overall similar shape (~~also~~ Figure 6(b)~~(ii)~~). The matched element, considering both the ISS-LIS and Meteorage distributions, occurs with similar probability earlier or later (or simultaneously) than the element itself and the distribution peak is centered at zero time offset. This is an interesting finding since e.g. Höller and Betz (2010) and Bitzer et al. (2016) found that TRMM-LIS detected lightning on average one to two milliseconds later than the ground-based LLSs. This is not ~~generally~~ the case for ISS-LIS in our study (and again one must consider the ISS-LIS integration time frame of 2.0 ms). Although the order of magnitude of the time offsets agrees well with our results. Timing differences can in fact be directly compared to those studies as the closest event provides the same time as the closest group (groups merge several events within the same time frame and in adjacent pixels of ISS-LIS).

The distribution given an IC pulse is also symmetric around zero and shows a maximum between ~~-1.0~~ −1.0 ms and 1.0 ms (Figure 6(b)). Its median (mean) is 0.00 ms (4.29 ms). For CG strokes, however, the distribution ~~maximum reaches from -1.0~~ peaks between −1.0 ms ~~to~~ and 0.0 ms. The negative distribution peak and median (mean) of ~~-0.07~~ −0.07 ms (~~-4.32~~ −4.32 ms) ~~indicates~~ indicate that ISS-LIS detected CG lightning slightly later than Meteorage. It might account for the time the light of the CG lightning needs to propagate towards the higher parts of the cloud and to become visible from space.

3.3 Characteristics of detected flashes

The previous sections dealt with relative DEs, location and times of coincident ISS-LIS and Meteorage records. In this section the unmatched flashes (42.7 % Meteorage, 16.7 % ISS-LIS) are also considered to investigate the following flash characteristics: the number of elements (events, pulses/strokes) per flash, flash extent, flash duration, flash mean absolute (pulse/stroke) amplitudes and individual pulse/stroke amplitudes, flash mean (event) ~~radiance~~ amplitude count, flash maximum (event) ~~radiance~~ amplitude count and the flash mean, minimum and maximum altitudes based on SAETTA observations. They are separated per matched and unmatched flash, per daytime (05:00 UTC to 17:00 UTC) and nighttime (17:00 UTC to 05:00 UTC) as well as per ~~IC and CG~~ flash type (IC, CG). The ISS-LIS flash IC or CG attribute depends on the type of the matched Meteorage flash. There is no flash type associated with ISS-LIS-only flashes. As explained in section 2.4, ISS-LIS events are analyzed. The statistical results obtained would be similar using groups instead of events, except for the flash extent and maximum ~~radiance~~ amplitude count per flash. It should be ~~mentioned~~ noted that especially the number of daytime CG flashes is very limited (24 ISS-LIS, 42 Meteorage meaning <10 %). Flash extents add the north-south and the east-west distance of a flash. The north-south distance uses the maximum and minimum latitude of the flash elements. The east-west distance of a flash is defined as the distance between the longitudinal maximum and minimum of the elements at the mean latitude (as that distance depends also on the latitude). Flash durations, or the times from the first to the last element of a flash, are ~~in general not limited~~ ~~in this study~~ only limited by viewing time of ISS-LIS. Theoretically, one flash could last for up to 90 s. Meteorage flash durations are not limited.

16

Table 2. Overall average and the average, minimum and maximum values of ISS-LIS flash characteristics for matched and unmatched flashes. The minimum (maximum) altitude represents the 10th (90th) percentile of concurrent SAETTA sources. Amplitude count is recorded per event.

	Average	Average matched	Average unmatched	Minimum matched	Minimum unmatched	Maximum matched	Maximum unmatched
Number of Events	51.2	56.6	23.7	1	1	518	116
Extent [km]	27.9	29.5	19.8	0.0	0.0	92.3	57.8
Duration [s]	0.32	0.35	0.20	0.0	0.0	1.69	0.89
Mean altitude [km]	8.2	8.2	8.2	2.9	3.6	10.9	11.6
Minimum altitude [km]	6.3	6.2	6.7	1.9	2.2	9.6	10.5
Maximum altitude [km]	9.9	10.0	9.6	2.9	4.6	12.8	12.9
Mean amplitude count	17.9	18.2	16.0	9.0	9.0	40.2	32.5
Maximum amplitude count	50.6	53.3	36.9	9.0	9.0	127.0	88.0

Statistics for ISS-LIS flash characteristics are summarized in Table 2. It contains the overall average and averages, minima and maxima observed for matched and unmatched ISS-LIS flashes. ISS-LIS flashes have on average 51.2 events~~(minimum 1, maximum 518). The flashes with a coincident Meteorage flash contain 56.7 events (minimum 1, maximum 518), while the .~~ Matched ISS-LIS flashes ~~without a match show only 23.7 events on average (minimum 1, maximum 116). About half (25) of the unmatched (matched) ISS-LIS flashes exhibit less than 10.0 events and about half (30)of those flashes were recorded during one single ISS-LIS frame. The~~ had more than twice as many events as the unmatched flashes (56.6 versus 23.7). The detailed event number distributions for matched and unmatched ISS-LIS flashes are shown in Figure 7(a) and (b), respectively. ~~It includes the histogram~~They include the histogram (i) and the CDF (ii). Daytime and nighttime flashes are distinguished for the flash types (only for matched flashes). The histogram's bars add the numbers of the different categories for the corresponding bin. All following Figures make use of the same layout. ISS-LIS nighttime flashes have about two times more events than daytime flashes. The background subtraction threshold for the optical signal is usually greater during daytime than during nighttime and the sensor is then more sensitive (difference of ~~about~~ 1.0 ~~$\mu J \cdot sr^{-1} m^{-2} \mu m^{-1}$, ,~~ overall minimum event ~~radiance~~ ampltiude count at 9.0 ~~$\mu J \cdot sr^{-1} m^{-2} \mu m^{-1}$)~~ and can influence the number of events per flash with a relative reduction of event numbers on bright backgrounds (daytime) compared to dark backgrounds (nighttime). Additionally, ISS-LIS CG flashes comprise on average approximately 11 % more events than IC flashes.

~~The average number of events in matched ISS-LIS flashes is more than two times higher than in unmatched flashes.~~ ~~Accordingly~~In accordance with the difference in event numbers, the matched ISS-LIS flashes feature ~~also~~ a larger extent and longer duration than the unmatched ISS-LIS flashes ~~. Figures 8 and 9 present the results for the~~ (Table 2). ISS-LIS flash

535 extents ~~and durations. ISS-LIS flash extents~~ range from 0 km (single events) to ~~about~~ 92 km~~. The average flash extents equal 19.8 and 29.5 for unmatched and matched ISS-LIS flashes, respectively~~, as shown in distributions in Figure 8. Peterson et al. (2018), who studied the evolution and structure of extreme flashes observed by TRMM-LIS, found a LIS flash with maximum event separation of 162 km. The size results likely from an elongation due to scattering of optically bright discharges.

~~Here, it is observed that almost all flashes (except for~~ This work found coincident Meteorage flashes for all but 4 ~~) with~~

540 ~~an extent~~ ISS-LIS flashes with extents exceeding 40 km~~have coincident Meteorage flashes~~. ISS-LIS nighttime flashes are on average about 5 km larger than daytime flashes (and comprise more events). The result ~~might indicate a better detection of dim events on very dark backgrounds during night compared to bright sunlit clouds at daytime~~is likely caused by the higher background subtraction threshold during the day than during the night. It could also result from an optical elongation of nighttime flashes. Large flashes with the maximum event separations in Peterson et al. (2018) also occurred at nighttime, but

545 the groups of these flashes were not separated by a significant fraction of the event separation. Fundamentally different cloud structures or types during day and night might also influence the results. It would need additional information, e.g. measuring infrared brightness temperatures for the cloud tops, to verify this hypothesis. Referring to the flash types, the mean extent of ISS-LIS CG flashes is about 5 km longer than for ISS-LIS IC flashes, however, the longest ISS-LIS flash is a nighttime IC flash.

550 One observed ISS-LIS flash lasted about 1.7 s (a CG nighttime flash), the longest duration found in this study. Peterson et al. (2018) found spurious flash durations up to 28 s in convective clouds, which result from high flash rates and slow storm motion. One large propagating flash lasted 5.04 s in their study. Figure 9 presents the duration distributions for the matched (a) and unmatched (b) ISS-LIS flashes. Matched ISS-LIS flashes last on average almost twice as long as ISS-LIS-only flashes, i.e. 0.35 s versus 0.20 s (~~Figure 9).~~ also Table 2). About 10 % (25 %) of the matched (unmatched) flashes were recorded during

555 a single LIS frame. Long lasting flashes (duration longer than 0.5 s) ~~have a~~ were detected by both LLSs with high probability (92.6 %)~~of being detected by both LLSs~~. ISS-LIS nighttime flashes last statistically 0.1 s longer than the daytime flashes. The result accords with the higher relative DE of ISS-LIS and more detected events during the night than during the day. ~~The ISS-LIS flash duration distribution does not significantly depend on the flash type.~~

~~Meteorage flashes contain between 1 and 54 pulses/strokes (overall~~ The averages, minima and maxima of Meteorage flash

560 characteristics are summarized in Table 3. Meteorage flashes contain on average 3.8 ~~)~~(but up to 54) strokes/pulses. The distributions of pulse/stroke numbers per matched and unmatched flash are presented in Figure 10(a) and (b), including the stacked histogram (i) and the CDFs (ii) as explained for Figure 7. Meteorage flashes seen by ISS-LIS are composed of 4.4 pulses/strokes on average. Meteorage-only flashes contain 2.9 pulses/strokes on average. ~~97 of the 326~~ 29.8 % of Meteorage flashes with coincident ISS-LIS flash have only one pulse/stroke (10 CG, 87 IC). ~~Among the 243~~ Single element flashes constitute about

565 43.2 % of the unmatched Meteorage flashes ~~, 105 single pulse/stroke flashes are found~~ (15 CG, 90 IC). As for ISS-LIS flashes, Meteorage flashes with match not only contain more pulses/strokes, but extend and last also longer than the unmatched flashes

Table 3. Overall average and the average, minimum and maximum values of Meteorage flash characteristics for matched and unmatched flashes. The minimum (maximum) altitude represents the 10th (90th) percentile of concurrent SAETTA sources. The maximum pulses/stroke amplitude is the maximum current observed for the flash (negative or positive).

	Average	Average matched	Average unmatched	Minimum matched	Minimum unmatched	Maximum matched	Maximum unmatched
Number of pulses/strokes	3.8	4.4	2.9	1	1	54	26
Extent [km]	9.9	12.1	6.8	0.0	0.0	147.5	109.2
Duration [s]	0.17	0.22	0.11	0.0	0.0	2.32	0.98
Mean altitude [km]	7.6	8.1	6.7	2.7	2.1	11.2	11.0
Minimum altitude [km]	5.7	6.1	5.1	1.7	1.6	10.3	9.4
Maximum altitude [km]	9.2	9.8	8.2	3.2	2.1	12.8	12.9
Mean absolute amplitde [kA]	9.5	8.0	11.6	1.3	1.1	98.1	102.9
Maximum pulse/stroke amplitude [kA]	-4.8	-2.2	-8.2	-149.0	-144.0	150.0	128.5

(Table 3). The flash extent distributions in Figure 11(a) and (b) show a mean (maximum) of 12.1 km (147.5 km) and 6.9 km (109.2 km) for matched and unmatched flashes, respectively. ISS-LIS detected all IC Meteorage flashes with extents above 32 km. The longest flashes are categorized as CG nighttime. In general, Meteorage CG flashes extend further than IC flashes. The mean extent equals 18.2 km (11.6 km) and 9.2 km (3.9 km) for matched (unmatched) CG and IC flashes, respectively. It is particularly small for unmatched IC flashes (as ISS-LIS can detect the longer IC flashes).

Meteorage flash durations support the findings, with matched flashes lasting on average (maximal) 0.22 s (2.3 s) and unmatched flashes lasting on average (maximal) 0.11 s (1.0 s). Figure 12(a) and (b) provide the duration distributions for Meteorage flashes. Distributions of both matched (a) and unmatched (b) flashes are sharply peaked for flashes shorter than 0.05 s (first bin; including single element flashes, maximum of 13 pulses/strokes per flash). The CDF (Figure 12(a) and (b) (ii)) illustrates that Meteorage CG flashes (mean 0.28 s) last statistically longer than Meteorage IC flashes (mean 0.11 s).

To conclude, matched flashes contain more elements, are more extended and last longer than unmatched flashes for both ISS-LIS and Meteorage records. Meteorage flashes appear to be on average both smaller in extent and shorter in duration than ISS-LIS flashes. The finding is in accordance with the different expectations on optical (LIS) and LF (Meteorage) signals.

580 Seven (2 ISS-LIS, 5 Meteorage) exceptionally long flashes (extent > 90 km or duration > 1.5 s) are analyzed using concurrent SAETTA observations. The VHF sources highlight that there can be concurrent flashes that either merge and form one flash or propagate at different height levels. ISS-LIS and Meteorage detect both types as continuous flashes as the LLSs capture the flashes two-dimensionally. They can in particular not distinguish the different altitudes for the latter.

585 ~~Conclusively, matched flashes contain more elements, are more extended and last longer than unmatched flashes for both ISS-LIS and Meteorage records. Meteorage flashes appear to be on average both smaller in extent and shorter in duration than ISS-LIS flashes. The finding is in accordance with the different expectations on optical (LIS) and LF (Meteorage) signals.~~

Then, the VHF SAETTA LLS is used to determine the altitude range of each ISS-LIS and Meteorage flash. Three additional flash characteristics are defined: flash mean altitude, flash minimum altitude and flash maximum altitude. The minimum altitude is defined as the 10th percentile of the altitudes of concurrent SAETTA sources rather than the true minimum. It is

590 aimed at reducing the influence of noise in the data. In the same manner the flash maximum altitude equals the 90th percentile of the concurrent SAETTA source altitudes instead of the true distribution maximum. Since not every ISS-LIS and Meteorage flash ~~happens within the SAETTA detection range~~ was detected by SAETTA, flash numbers are reduced compared to the ones discussed in section 3.1 to 256 ISS-LIS flashes with match, 43 ISS-LIS-only flashes, 292 Meteorage flashes with match and 188 Meteorage-only flashes. ISS-LIS mean event ~~radiance~~ amplitude count and Meteorage mean pulse/stroke amplitude

595 distributions are examined for the analyzed altitude levels of flashes. ~~It should be mentioned that the~~ The lowest altitude of detectable VHF sources increases with distance to the LMA network mainly due to Earth's curvature and also due to shading by the relief, especially in the South of the domain (Coquillat et al., 2019). Hence, flash ~~mean and especially minimum~~ minimum (and mean) altitudes might suffer from undetected VHF sources in low altitudes.

Figure 13 presents the mean altitude of matched (a) and unmatched (b) ISS-LIS flashes as histograms (i) and CDFs (ii).

600 The distribution of the flash mean ~~radiances~~ amplitude counts in each altitude bin ~~is~~ are included as blue boxplot diagram (with mean marked as diamond, outliers not plotted). ~~LIS flashes with coincident Meteorage flash have an average mean altitude of 8.2 (Figure 13(a)(i)). The same average mean flash altitude is found for ISS-LIS-only flashes (Figure 13(b)(i)). The distribution of~~ The distribution of mean altitudes for unmatched ISS-LIS flashes fits that of the matched ISS-LIS flashes (although the number of unmatched flashes is low~~.~~). The mean flash altitudes average about 8.2 km (Table 2).

605 The overall ISS-LIS flash mean ~~(maximum)~~ altitude distribution, that is dominated by 83.3 % flashes with match, peaks at about 9.5 km~~(11.0 to 11.5)~~, as shown in the ~~histograms in in Figures~~ histogram in Figure 13(a)(i)~~and ??(a)(i)~~. The daytime distribution has a second mode near 5.0 km ~~(7.0)~~ of altitude. ~~About half of the flashes reach altitudes of 11.0~~ISS-LIS flashes reach on average altitudes of 9.9 km and ~~above (Figure ??(ii)),~~ were observed up to almost 13 km of altitude (Table 2) (a noteworthy high value considering the tropopause in ten to twelve kilometers of altitude).

610 Differences between matched and unmatched ISS-LIS minimum flash altitudes approximate 0.5 km, with matched flashes showing lower minima (distributions not shown). The difference is significant as it exceeds the predicted SAETTA altitude error (about 0.2 km over wide parts of the domain). 89.7 % of the 126 ISS-LIS flashes with minima less (or equal) than 6.0 km of altitude have a coincident Meteorage flash. ISS-LIS flashes with minima above 6.0 km (173) are detected by Meteorage in

82.7 % of the cases. ~~It can be constituted that~~ Overall, Meteorage better detected low altitude ISS-LIS flashes than ISS-LIS flashes restricted to mid and high levels.

The ~~radiance~~ amplitude count of ISS-LIS flashes increases in general with the mean altitude (Figure 13(a)(i) and (b)(i)). The highest observed flash mean altitudes occur mainly for pure IC flashes and show statistically high ~~radiances. They are likely~~ amplitude counts. They likely origin within high reaching convective clouds like cumulus congestus and cumulonimbus. ~~The average mean radiance for the matched and unmatched ISS-LIS flashes yields 18.2 $\mu J \cdot sr^{-1} m^{-2} \mu m^{-1}$ and 16.0 $\mu J \cdot sr^{-1} m^{-2} \mu m^{-1}$, respectively.~~ Similar results regarding the ~~mean radiance distributions can be~~ amplitude count distributions were identified for the ISS-LIS flash maximum altitude distributions and maximum event amplitude count per flash at different altitudes (not shown).

~~Figure ??(i) illustrates the results of maximum radiance per flash within the maximum flash altitude histogram. The highest flashes do not necessarily contain the brightest events, however, the average trend is that maximum radiances slightly increase with the altitude of the flash. The maximum radiance distributions for the altitude bins show a wide spread. Medians of bins between 8 and 13 maximum flash altitude remain within the corresponding Inter-Quartile-Ranges (IQRs)of neighboring bins. The~~ Matched and unmatched ISS-LIS flashes feature almost similar mean amplitude counts (Table 2). The overall brightest event (127.0 ~~$\mu J \cdot sr^{-1} m^{-2} \mu m^{-1}$)~~) occurred during nighttime for a matched flash. The strongest optical signal during the day (105.0 ~~$\mu J \cdot sr^{-1} m^{-2} \mu m^{-1}$)~~) is recorded within a matched flash, too. Accordingly, ~~an average maximum radiance per flash~~ maximum event amplitude counts averaged over all flashes of about 53.3 $\mu J \cdot sr^{-1} m^{-2} \mu m^{-1}$ and 36.9 $\mu J \cdot sr^{-1} m^{-2} \mu m^{-1}$ ~~characterizes~~ characterize matched and unmatched flashes, respectively (Table 2). Flashes containing the optically brightest events have a higher chance of producing significant LF signals and being detected by Meteorage than the optically ~~dark~~ darker flashes.

~~As Figure 13 for the maximum altitude of ISS-LIS flashes with (a) and without (b) match. Here, the blue boxplots (median as line, mean as diamond, IQR as box, 1.5 IQR as whiskers - outliers not plotted) represent the distributions of ISS-LIS maximum (event) radiance per flash for each altitude bin.~~

The comparison of altitudes of Meteorage flashes with and without ISS-LIS matches aims at studying how ISS-LIS can detect low altitude flashes. Flash mean (absolute) amplitude and maximum pulse/stroke amplitude ~~and maximum amplitude~~ per flash are additionally analyzed for each altitude bin in the histograms. The maximum amplitude per flash can either show a positive or negative current. Results are presented in Figure ~~?? for the altitude mean and in Figure~~ 14 for the altitude maximum with the mean absolute ~~pulse/stroke~~ amplitude per flash. ~~An average mean altitude of~~ Mean flash altitudes average 8.1 km ~~is found~~ for Meteorage flashes with match (~~Figure ??(a)). It is about~~ Table 3). They are on average 1.4 km lower for the Meteorage-only flashes~~(Figure ??(b))~~. The mean altitude of matched flashes is similar to that of ISS-LIS matched flashes (Figure 13(a)). The unmatched flashes, however, differ by about 1.5 km in altitude ~~. Meteorage-only flashes occur at significantly lower altitude than the matched flashes(matched and unmatched ISS-LIS flashes are found at similar altitudes~~(difference well above the SAETTA altitude uncertainty). Meteorage flash maximum altitudes confirm this result (Figure 14): Flashes with coincident ISS-LIS flash reach on average 9.8 km of altitude. The Meteorage-only flashes feature ~~an average~~ a maximum altitude of 8.2 km on average. The maximum altitude distribution peaks, as for the ISS-LIS matched flashes, at about 11.0 km altitude

21

(Figure 14). For the Meteorage-only flashes, another mode exists between 6.5 km and 7.0 km of altitude. The ~~low altitude mode is also found in the mean flash altitude distribution (at about 5.0 in Figure ??)~~ daytime distribution of unmatched Meteorage flashes peaks at the lower altitudes. It is indicative of ISS-LIS' reduced DE ~~of~~ for low altitude flashes (during all times and even more pronounced at daytime than during the night). Meteorage flashes with maxima exceeding 10.0 km (248) are detected by ISS-LIS in 75.4 % of the cases. ISS-LIS' relative DE for Meteorage flashes with maxima lower (or equal) than 10.0 km (232) is only 45.3 %. This trend still influences the flash ~~minima. Figure 15 shows the distribution of minimum flash altitudes with the maximum (pulse/stroke)amplitude per flash in each altitude bin. Matched and unmatched flash average minimum altitudes approximate 6.1 and 5.1 , respectively.~~ mean and minimum altitudes (Table 3). Hence, it is confirmed that ISS-LIS flash detection declines from high to low altitude flashes. The result agrees with the case study of (Thomas et al., 2000), who found significantly less skill of TRMM-LIS for (CG) discharges near the cloud base than for lightning channels propagating to near the top of the clouds.

~~Flash mean altitude of Meteorage flashes (from concurrent SAETTA observations) with coincident ISS-LIS flash (a) and unmatched Meteorage flashes (b). Daytime, nighttime and flash type, IC and CG, are indicated by the colors. Histogram (i) and corresponding CDF (ii) use the same colors. The CDF shows in addition a black curve for all data. The histogram bin width equals 0.5 . The mean value is plotted as dashed line. The total number of flashes is indicate above the histogram. The blue boxplots (median as line, mean as diamond, IQR as box, 1.5 IQR as whiskers - outliers not plotted) represent the distributions of Meteorage mean absolute pulse/stroke amplitude per flash for each altitude bin.~~

Figure 15 shows the distribution of minimum flash altitudes with the maximum (pulse/stroke) amplitude per flash in each altitude bin. Low altitude flashes (minimum altitudes below 5.0 km) feature statistically higher flash mean (not plotted) and maximum amplitudes than flashes occurring above 5.0 km of altitude (Figure 15). Those flashes are mainly identified as CG flashes. The analysis of the flash maximum amplitude shows that those low altitude flashes are dominated by negative maximum currents. The flashes with minimum altitudes above 5.0 km exhibit statistically more positive than negative maximum currents. Further investigation reveals that about 94 % of the (absolute) currents above 22.5 kA ~~are observed for~~ belong to CG strokes. The strongest currents reach up to 150.0 kA (both negative and positive currents) and are related to CG strokes. CG strokes have almost exclusively negative currents in this study. IC pulse currents do not exceed 50 kA. About 90 % of pulses/strokes with amplitude below 10.0 kA are IC pulses. ~~CG strokes have almost exclusively negative currents in this study~~A similar result is provided by Cummins and Murphy (2009). They found that 90 % of positive LF currents with less than 10.0 kA belong to IC pulses. Negative currents are ~~also~~ observed for approximately 26 % of the IC pulses. ~~The strongest currents reach up to 150.0 (both negative and positive currents) and occur for CG strokes. IC pulse currents do not exceed 50 .~~

The Meteorage mean (maximum) flash absolute amplitude equals 8.0 kA (13.2 kA) and 11.6 kA (18.1 kA) for matched and unmatched flashes, respectively. The difference between matched and unmatched flashes is attributed to some low to mid-level flashes producing strong currents and being not detected by ISS-LIS (~~difference between~~ compare (a) and (b) in Figures ?? 14 and 15). However, the overall distributions of absolute flash amplitudes appear to be similar for matched and unmatched Meteorage flashes. ~~One might use~~

Flashes observed in this study show a statistical relationship between the polarity of the maximum ~~current to identify whether the flash occurs below or above 5.0 . This accounts~~ (LF) current and the altitude. The relationship was detailed for the flash minimum ~~altitude. The relationship~~ altitudes and appears also for the flash maximum altitudes. ~~Here~~In this study, flashes with maximum altitudes below 10.0 km exhibit mainly negative maximum currents. ~~It was investigated~~ As it was found that ISS-LIS' DE is 30 % higher for flashes with maximum altitudes above 10.0 km than for flashes restricted to lower levels~~. Hence~~, the polarity of the flash maximum current ~~can~~ might provide a first information whether a flash is detected by ISS-LIS. This finding is probably specific for storm types and flashes used in this study (and region). The observed relationship between the polarity of the maximum current of a flash and its altitude might change for inverted polarity storms or hybrid (IC+CG) flashes.

4 Conclusions

This study compares the results of the LF ground-based Meteorage LLS, the satellite sensor ISS-LIS and the VHF ground-based LMA SAETTA. The study domain is bounded to a region near Corsica in the Mediterranean Sea where SAETTA data are available. As ISS-LIS has been operating since March 2017, the period is confined to about one year from March 01, 2017 to March 20, 2018.

A new algorithm is developed to group both ISS-LIS events and Meteorage pulses/strokes to flashes. The algorithm is validated using concurrent SAETTA observations and the results of the existing NASA LIS algorithm.

ISS-LIS ~~detects~~ detected in total 16,881 events distributed over 330 flashes during its overpasses over the study domain. Meteorage data are filtered for the times of ISS overpasses. It contains 2,144 pulses/strokes (487 CG, 1657 IC) in 569 flashes. ISS-LIS ~~detects~~ detected about 57.3 % of the Meteorage flashes. Especially ~~CG flashes~~ cloud-to-ground (CG) flashes and single pulse ~~IC flashes decrease the total~~ intra-cloud/cloud-to-cloud (IC) flashes decrease the overall relative detection efficiency (DE) of ISS-LIS. A relative DE of 53.9 % ~~is~~ was observed for flashes detected by Meteorage at daytime. LIS detected Meteorage ~~intra-cloud/cloud-to-cloud (IC)~~ IC flashes with about six percent higher relative DE than ~~cloud-to-ground (CG)~~ CG flashes. The LF Meteorage LLS ~~is~~ was able to detect more than 80 % of all occurring ISS-LIS flashes.

Distances and timing offsets between matched ISS-LIS and Meteorage flashes are analyzed. A ~~peak~~ mode (median) distance (given a Meteorage flash) of about 1.8 km (2.3 km) states a fairly accurate collocation of the flashes. Given an ISS-LIS flash, the ~~peak~~ mode (median) distance equals about 3.0 km (4.7 km). ~~It is generally smaller than~~ The majority of distances between matched flashes are within the ISS-LIS pixel resolution (4.5 km nadir, 6.2 km at the edge of the field of view). The absolute timing offset distribution between a given Meteorage flash and the matched ISS-LIS flash is sharply peaked for less than 1.0 ms. Considering the ISS-LIS frame integration time of 2.0 ms, this is a very satisfying result. An analysis of the closest elements (events and pulses/strokes) reveals that, ~~to an overall equal proportion~~with similar probability, ISS-LIS or Meteorage ~~detect~~ detected a lightning first while the ~~peak~~ mode timing offsets remain within the LIS frame integration time. For CG strokes, however, ISS-LIS ~~tends~~ tended to detect the lightning activity later than Meteorage. All offsets increase relatively from the distribution given a Meteorage flash to the distribution given an ISS-LIS flash. This finding is likely caused by the significantly

lower number of pulses/strokes than the number of events. Thus, it is more likely to find an event close to a pulse/stroke than vice versa.

For an enhanced understanding of the flash detection by ISS-LIS and Meteorage, characteristics of the flashes are investigated. In accordance with e.g. Rudlosky et al. (2017) and Zhang et al. (2019), the probability of a match increases with larger flash extent and flash duration. A matched flash ~~can extend~~ extended on average almost twice as wide and ~~last~~ lasted twice as long as a flash not seen by both ISS-LIS and Meteorage. In a similar manner, the matched flashes ~~contain~~ contained on average twice the number of elements than a flash observed by only one of the LLSs. ISS-LIS is sensitive to optical signals while Meteorage detects LF signals of electrical discharges. Nevertheless, ISS-LIS flashes with at least one very bright event ~~are~~ were more likely to be detected by Meteorage than optically ~~dark~~ darker flashes. Using the 3D lightning location of concurrent SAETTA observations, ISS-LIS and Meteorage flash altitudes are compared. Altitude-related behaviors are likely driven by the range of IC and CG flash altitudes. Detailed results for the flash types are given in section 3.3. Matched flashes of both ISS-LIS and Meteorage feature similar ~~average~~ mean altitudes near 8.2 km on average. Unmatched Meteorage flashes occurred on average 1.4 km lower than Meteorage flashes seen by ISS-LIS. Especially the maximum altitude of a flash ~~influences~~ influenced significantly the detectability by ISS-LIS (compare e.g. Thomas et al. (2000)). Meteorage flashes with maxima exceeding 10.0 km of altitude ~~are~~ were detected by ISS-LIS in 75.4 % of the cases. ISS-LIS' relative DE for Meteorage flashes with maxima lower than 10.0 km of altitude is only 45.3 %. The Meteorage flash detection ~~depends~~ depended slightly on the flash minimum altitude. 89.7 % of the ISS-LIS flashes with minima less than 6.0 km of altitude ~~have~~ had a coincident Meteorage flash. ISS-LIS flashes with minima above 6.0 km of altitude ~~have~~ had a coincident Meteorage flash in 82.7 % of the cases.

Further investigation revealed that the ~~radiance~~ optical brightness of ISS-LIS flashes is somewhat correlated to the flash altitude with increasing (both mean and maximum event) ~~radiances~~ amplitude counts for increasing flash altitudes. Meteorage amplitudes ~~increase~~ increased statistically with decreasing flash altitudes. Especially the polarity and the current of the strongest pulse/stroke within a Meteorage flash ~~show~~ showed potential to gain qualitative flash altitude information. Flashes with maximum currents of negative 10 kA or lower ~~remain~~ remained mainly below 10.0 km of altitude. As stated earlier, ISS-LIS' relative DE ~~is~~ was 30 % higher for those flashes than for flashes with maximum altitudes above 10.0 km. This finding will need ~~some~~ additional proof, but it can be useful for mimicking satellite lightning products using LF LLSs.

This study analyzes satellite observed lightning over an extra-tropical region and compares the observations to ground-based LLSs. Our results including the statistics use about one year of data within the limited region around Corsica island. This results in a limited number of lightning cases. The limited region enables the direct unique comparison of not only ISS-LIS and LF Meteorage but also the VHF SAETTA LLS. Hence, ISS-LIS and Meteorage flash detection is investigated in more detail, e.g. considering the concurrent SAETTA lightning source altitudes. The coincidences between ISS-LIS and Meteorage flashes do not always have a one-to-one correspondence. It is, in addition, an artifact of the relatively coarse match constraints of 20.0 km in space and 1.0 s in time. The constraints are validated and their influence on the results is seen in the matched distance and timing offset distributions. It should be mentioned that the available ISS-LIS data is the provisional P0.2 version for this work. It is close to but not quite the fully validated data of ISS-LIS. Due to our limited number of cases, all ISS-LIS data are treated in the same way independent of the position within ISS-LIS field of view (FOV). It is known that the ISS-LIS

pixel (event) resolution and the DE decrease near the edge of the FOV. However, it is decided to not filter and reduce the observed cases further in order to allow a statistical analysis. Our method can be applied to geostationary satellite LLSs, i.e. GLM and the future MTG-LI, and the comparison of their observations to ground-based LLSs. It is planned to study GLM and NLDN lightning observations in America using our methodology. The geostationary satellite observes one region continuously and thus there will be many more cases for the statistics. The results might be compared to our results of the comparison of ISS-LIS and Meteorage.

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27
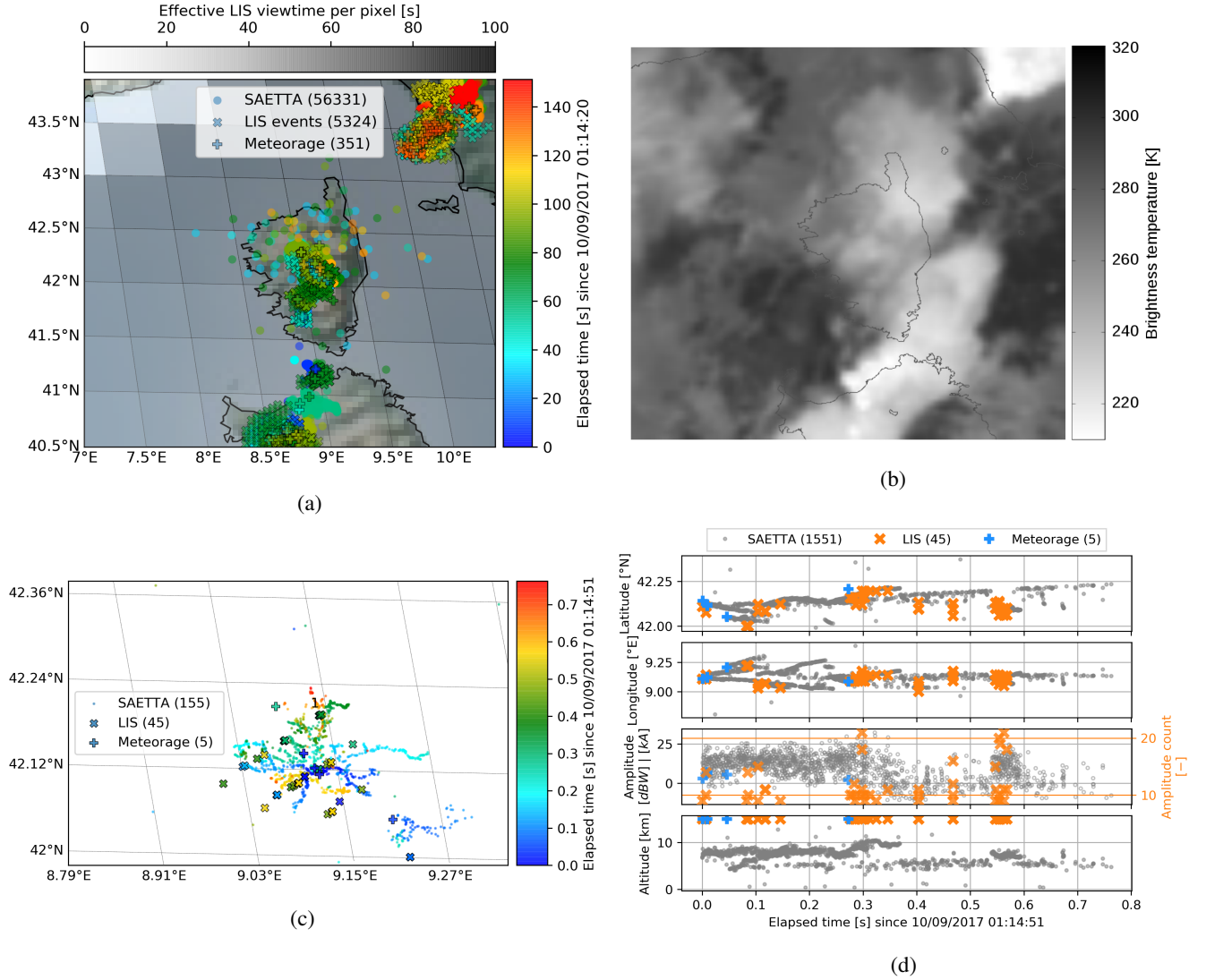
Figure 1. Observations of ISS-LIS events (as pixel centers), Meteorage pulses/strokes and SAETTA VHF sources (as indicated) during one ISS overpass over Corsica on Sept. 10, 2017 (a). The ISS-LIS viewtime is presented as grayscale of the background. Numbers in parentheses give the number of SAETTA VHF sources, ISS-LIS events and Meteorage pulses/strokes, respectively. (b) shows the infrared (IR 10.8 μm) satellite image of the same day at 01:15:00 UTC. One flash over Corsica detected by the three LLSs during the same ISS overpass is shown in (c) as map and (d) as time series of latitude, longitude, signal strength (amplitude [dBW] for SAETTA, [kA] for Meteorage and amplitude count for ISS-LIS) and altitude (LIS and Meteorage set to 15 km).

Observations of ISS-LIS events (as pixel centers), Meteorage pulses/strokes and SAETTA VHF sources (as indicated) during one ISS overpass over Corsica on Sept. 10, 2017 (a). The ISS-LIS viewtime is presented as grayscale of the background. Numbers in parentheses give the number of SAETTA VHF sources, ISS-LIS events and Meteorage pulses/strokes, respectively. (b) shows the infrared (IR 10.8) satellite image of the same day at 01:15:00 UTC. One flash over Corsica detected by the three LLSs during the same ISS overpass is shown in (c).
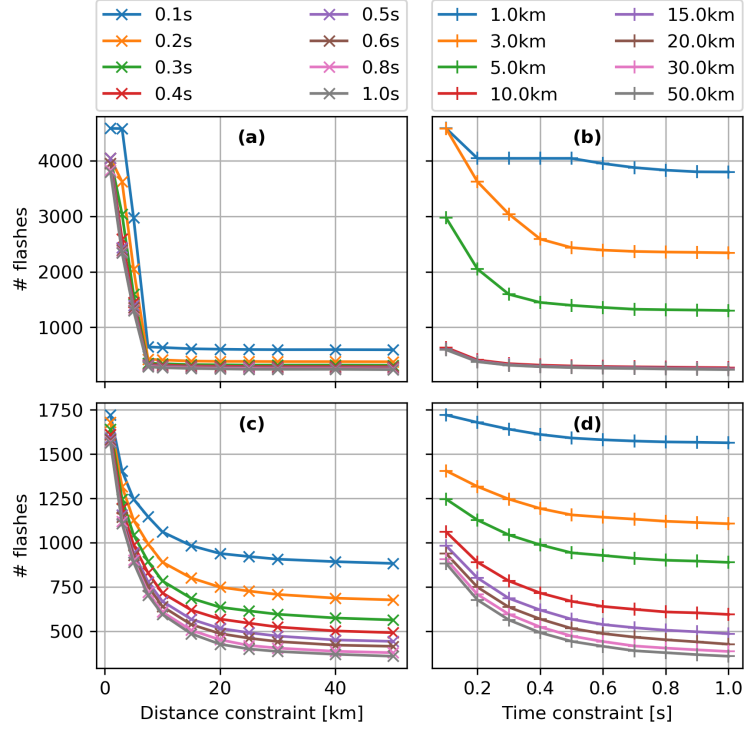
Figure 2. Total flash number based on constant, equal time constraint dt_{merge} (line color) with varying distance ds_{merge} for elements of ISS-LIS (a) and Meteorage (c) flashes. (b) and (d) with constant, equal distance ds_{merge} (line color) and varying dt_{merge} for ISS-LIS and Meteorage, respectively.
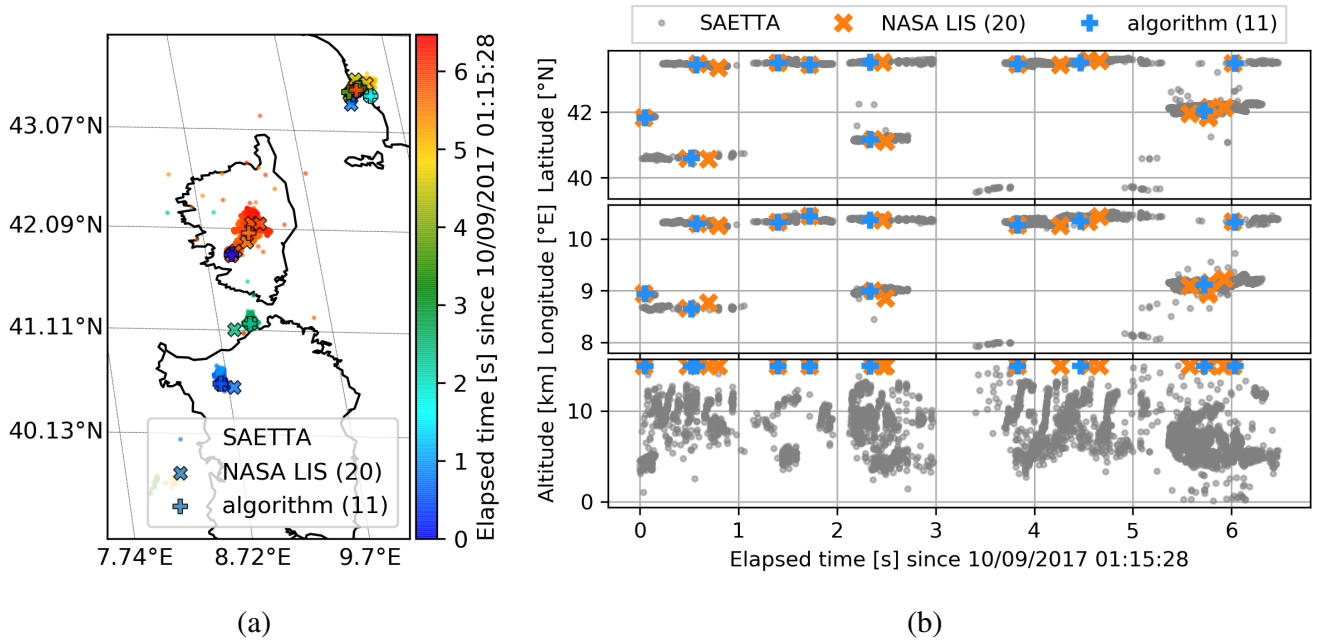
Figure 3. The map (a) and time series (b) of SAETTA observations and the mean flash positions based on the NASA LIS algorithm and our developed algorithm for one situation on Sept. 10, 2017. The numbers in parentheses in the legend indicate the number of identified flashes (both algorithms analyzed the same LIS events). Colors in (a) represent the elapsed time from the initial lightning activity. Note: Flash altitudes in (b) are not known for ISS-LIS flashes but arbitrarily plotted at 15 km.

One situation on Sept. 10, 2017. The map (a) and time series (b) with SAETTA observations and the mean flash positions based on the NASA LIS algorithm and our developed algorithm. The numbers in parentheses in the legend indicate the number of identified flashes for both algorithms analyzing the same events. Colors in (a) represent the elapsed time from the initial lightning activity.
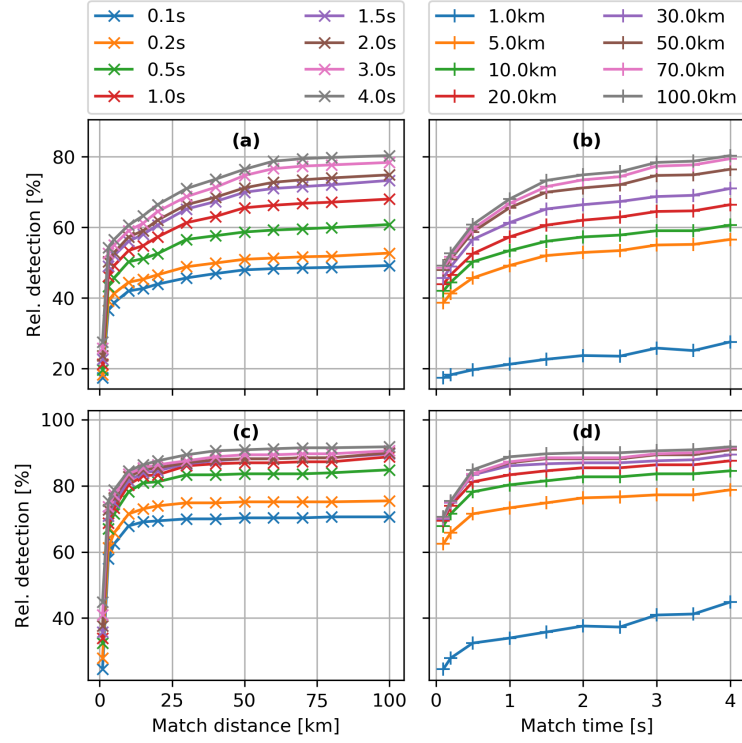
Figure 4. Relative detection efficiency based on constant, equal time criterion dt_{match} (line color) with varying distance ds_{match} for ISS-LIS (a) and Meteorage (c). (b) and (d) with constant, equal distance ds_{match} (line color) and varying dt_{match} for ISS-LIS and Meteorage, respectively.

(a)   (b)
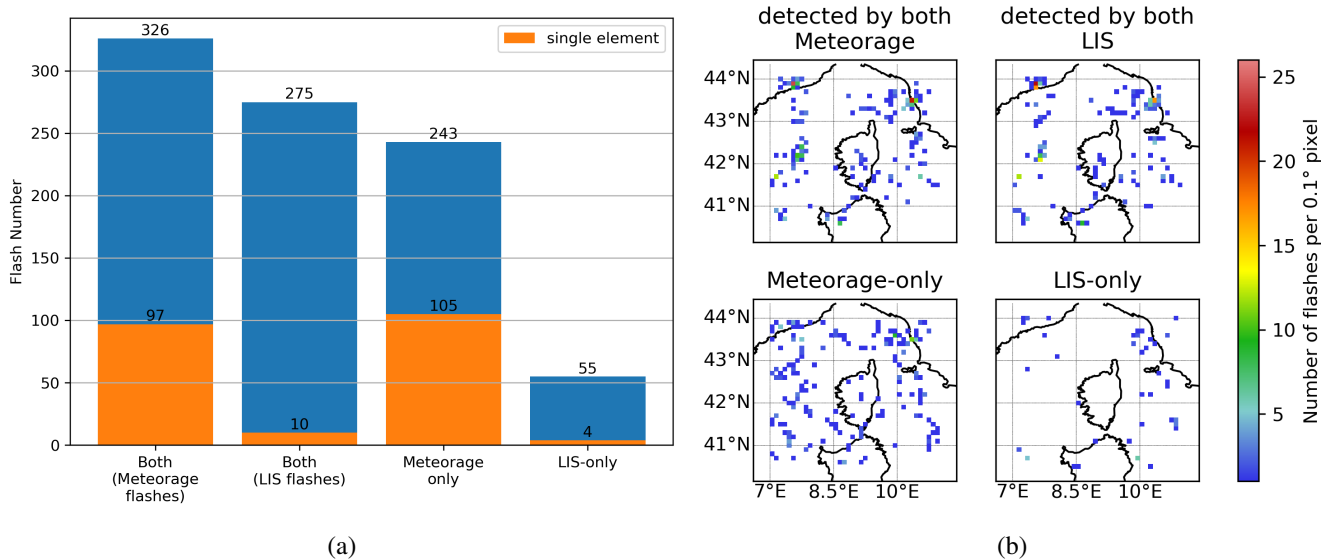
Figure 5. Flash category histogram (a) and spatial distribution (b) while matching ISS-LIS and Meteorage flashes for the available ISS overpasses from March 2017 to March 2018. Categories show the number of Meteorage flashes seen by ISS-LIS [Both (Meteorage flashes)], the number of ISS-LIS flashes detected also by Meteorage [Both (LIS flashes)] and flashes detected either by Meteorage [Meteorage-only] or ISS-LIS [LIS-only]. The numbers of single element flashes (event for ISS-LIS, pulse/stroke for Meteorage) are marked for each category as indicated.

Flash category histogram (a) and spatial distribution (b) while matching ISS-LIS and Meteorage flashes for the available ISS overpasses from March 2017 to March 2018. Categories show the number of Meteorage flashes seen by ISS-LIS Both (Meteorage flashes), the number of ISS-LIS flashes detected also by Meteorage Both (LIS flashes)and flashes detected either by Meteorage Meteorage-onlyor ISS-LIS LIS-only. The numbers of single element flashes (event for ISS-LIS, pulse/stroke for Meteorage) are marked for each category as indicated. The color scale in the maps (b) states the number of flashes per 0.1° pixel.
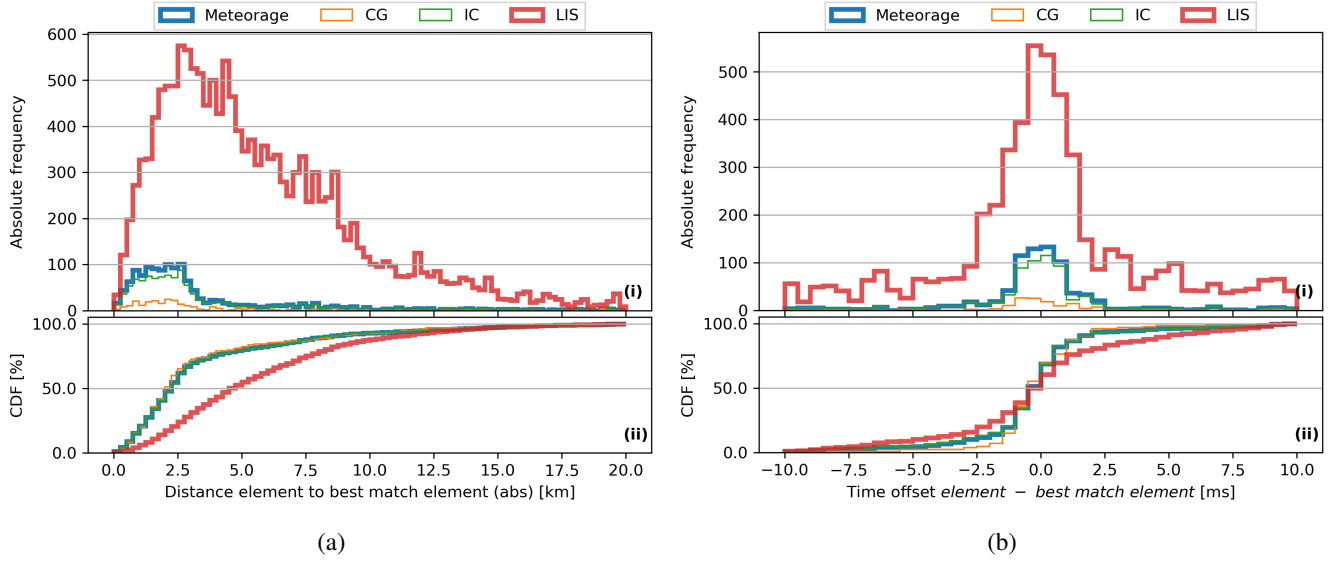
Figure 6. Best match distance (a) and time offset (b) between a given Meteorage pulse/stroke (LIS event) and the closest ISS-LIS event (Meteorage pulse/stroke). Histogram (i) and cumulative distribution function (ii) with bins of $0.25\,\mathrm{km}$ (a) and $0.5\,\mathrm{ms}$ (b). Analyzed pulses/strokes (events) belong to flashes with matches (spatial and temporal filters) while the pulses/strokes (events) of unmatched flashes are not considered. For Meteorage, the discharge types [CG, IC] are distinguished. Only elements with absolute timing offsets of less than $10\,\mathrm{ms}$ are included in the plotted time offset distribution. A positive time indicates the given element occurred later than the best match.
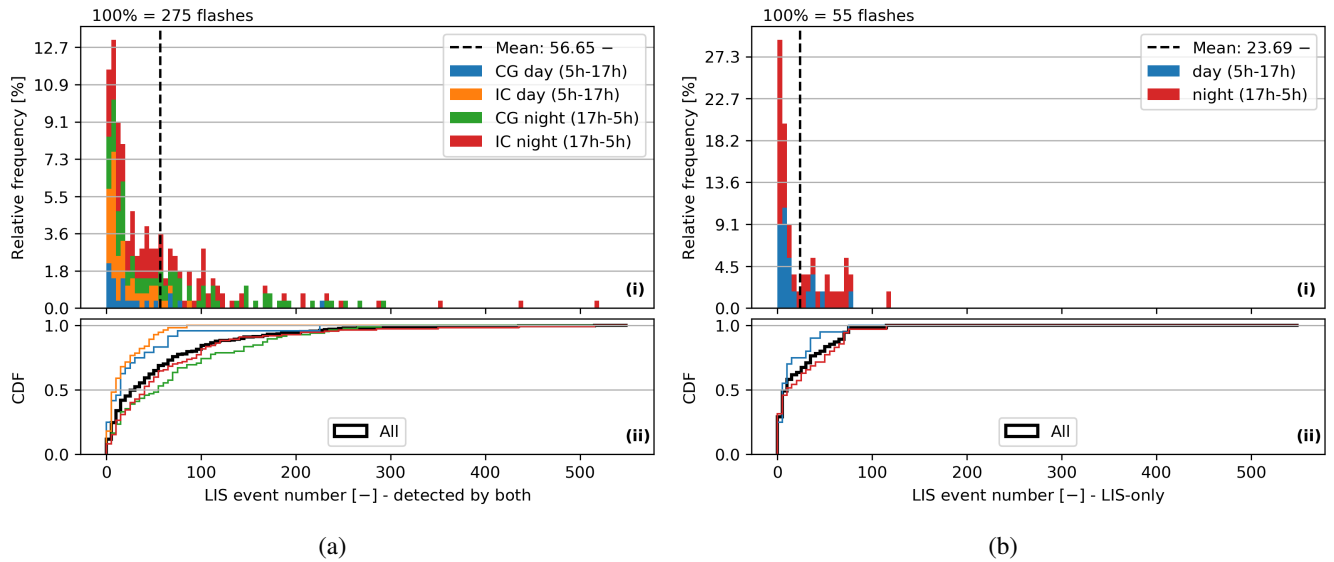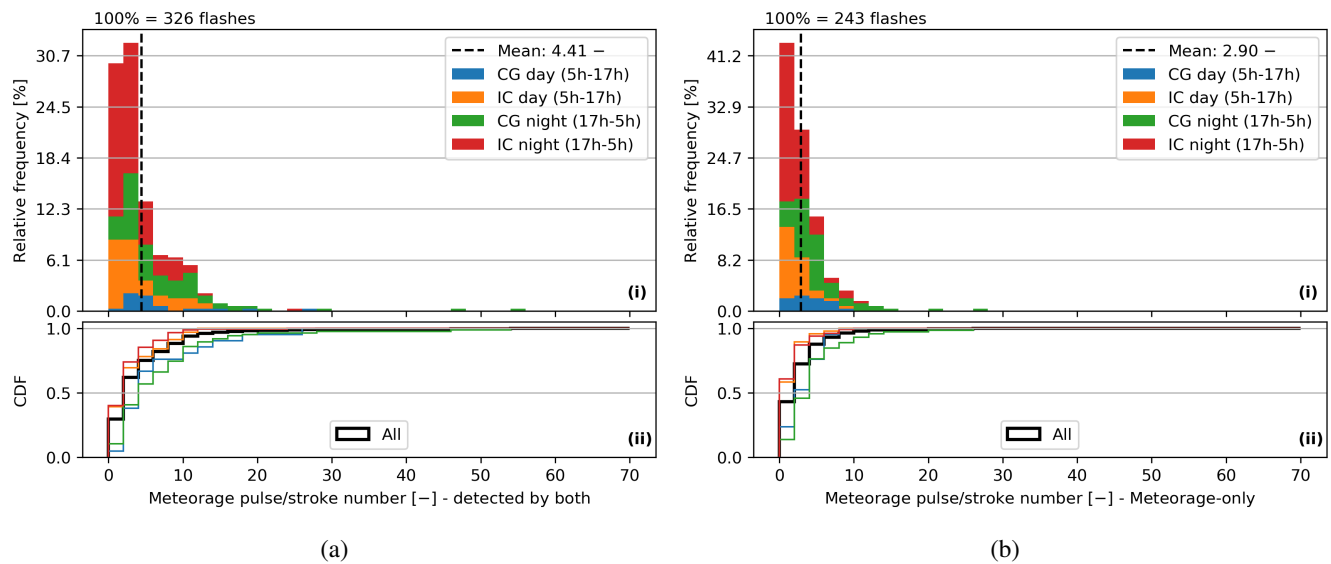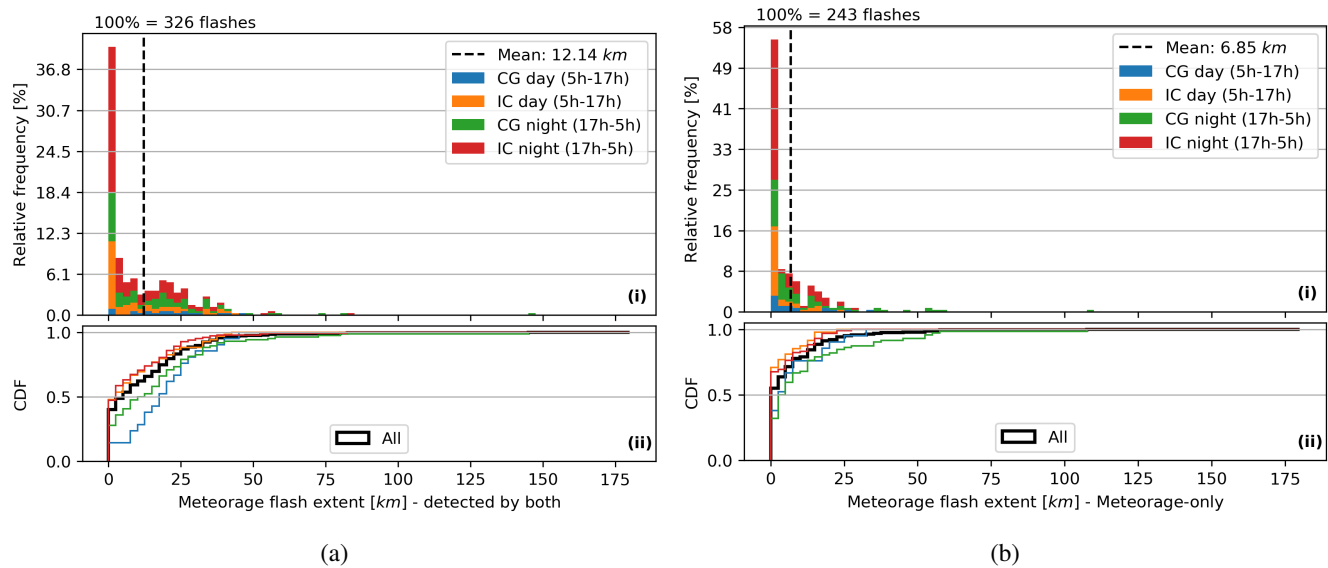
Figure 7. LIS event numbers of ISS-LIS flashes with coincident Meteorage flash (a) and unmatched ISS-LIS flashes (b). Daytime, nighttime and flash type, IC and CG, are indicated by the colors. Histogram (i) and corresponding CDF (ii) use the same colors. The CDF shows in addition a black curve for all data. The histogram bin width is constant at 5 events. Note: The CG/IC attribute for ISS-LIS flash needs the matched Meteorage flash and does not exist for ISS-LIS-only flashes. The mean value is plotted as dashed line. The total number of flashes is indicate above the histogram.
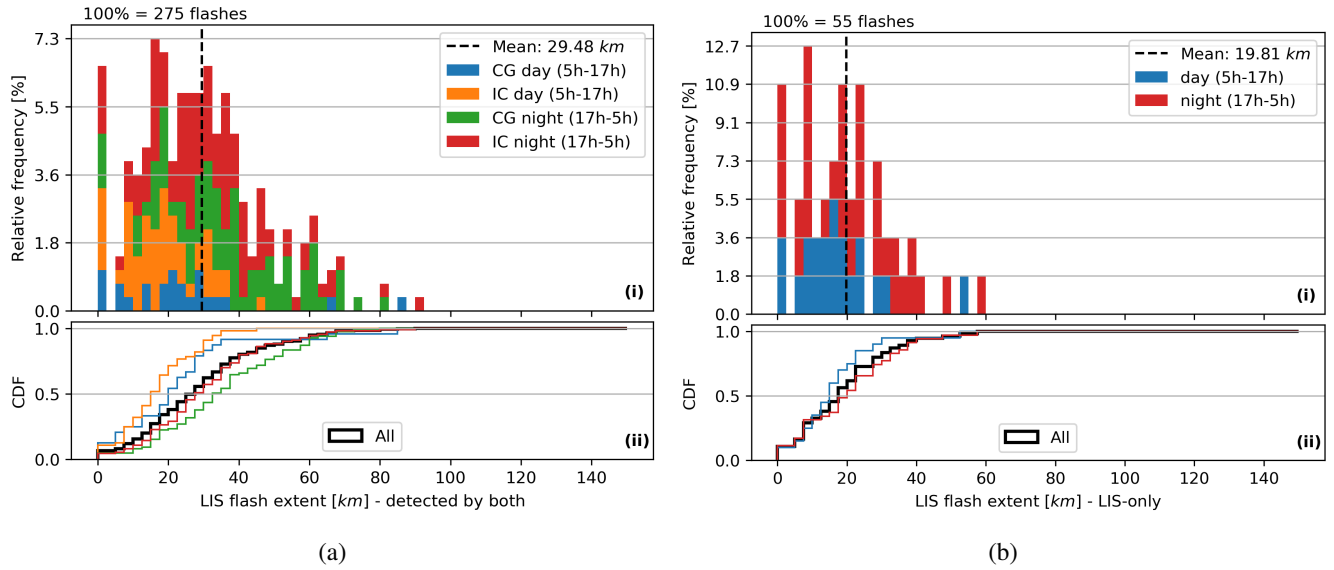
Figure 8. Flash extents of ISS-LIS flashes with coincident Meteorage flash (a) and unmatched ISS-LIS flashes (b). Daytime, nighttime and flash type, IC and CG, are indicated by the colors. Histogram (i) and corresponding CDF (ii) use the same colors. The CDF shows in addition a black curve for all data. The histogram bin width is constant at 2.5 km. Note: The CG/IC attribute for ISS-LIS flash needs the matched Meteorage flash and does not exist for ISS-LIS-only flashes. The mean value is plotted as dashed line. The total number of flashes is indicate above the histogram.
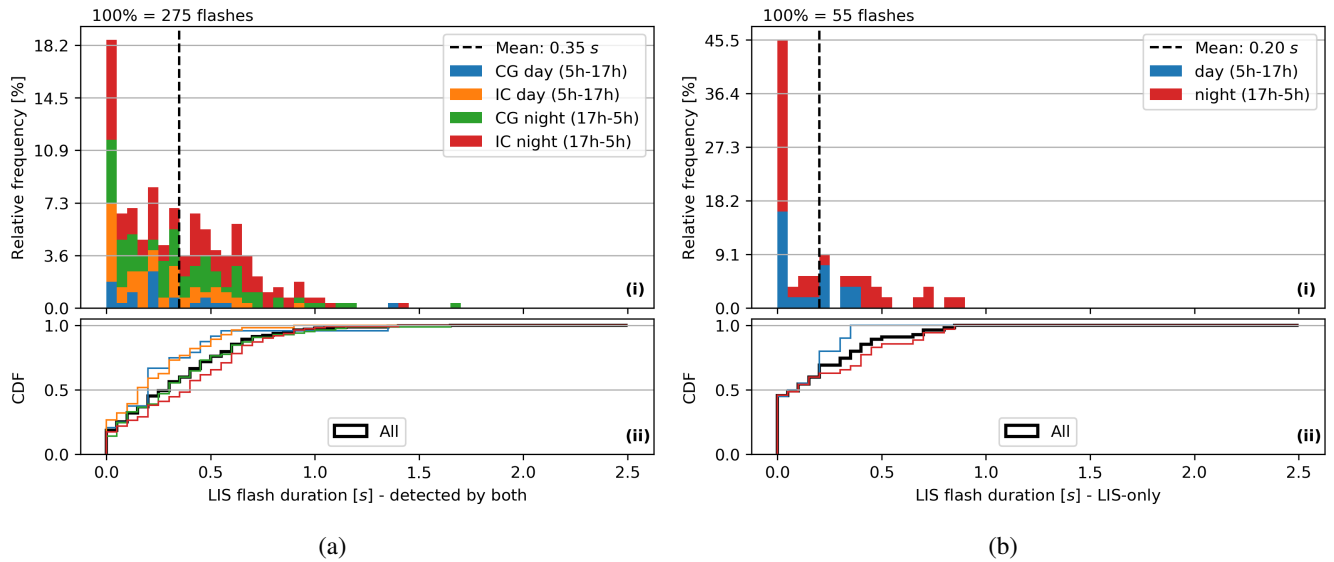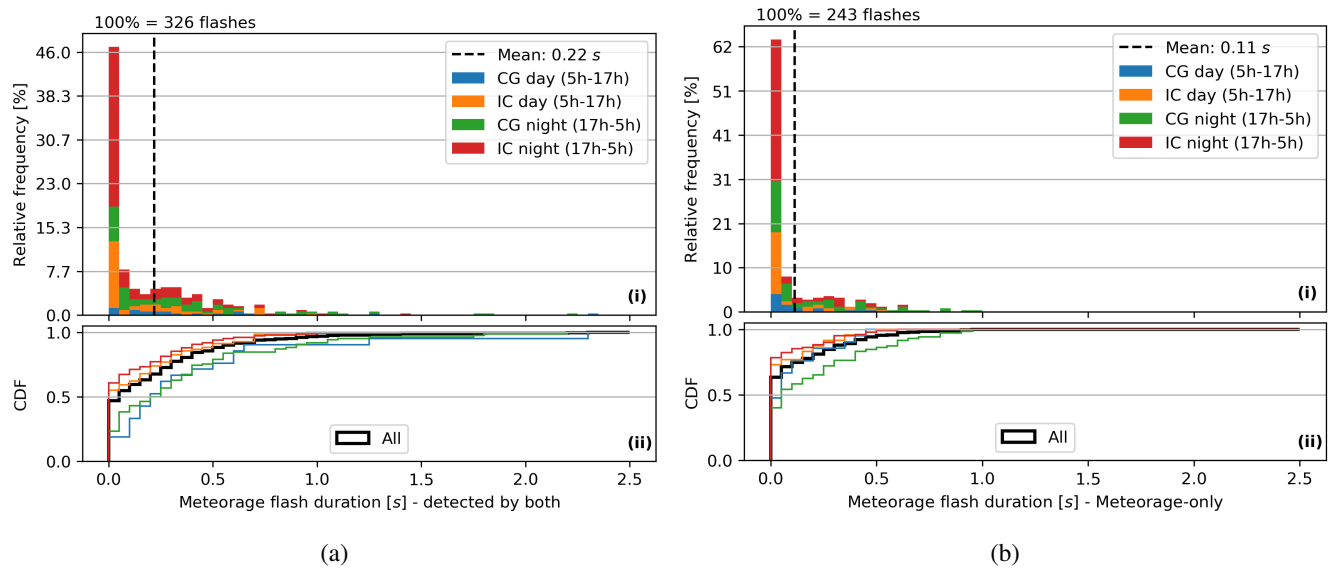
Figure 9. Flash duration of ISS-LIS flashes with coincident Meteorage flash (a) and unmatched ISS-LIS flashes (b). Daytime, nighttime and flash type, IC and CG, are indicated by the colors. Histogram (i) and corresponding CDF (ii) use the same colors. The CDF shows in addition a black curve for all data. The histogram bin width is constant at 0.05 s. Note: The CG/IC attribute for ISS-LIS flash needs the matched Meteorage flash and does not exist for ISS-LIS-only flashes. The mean value is plotted as dashed line. The total number of flashes is indicate above the histogram.
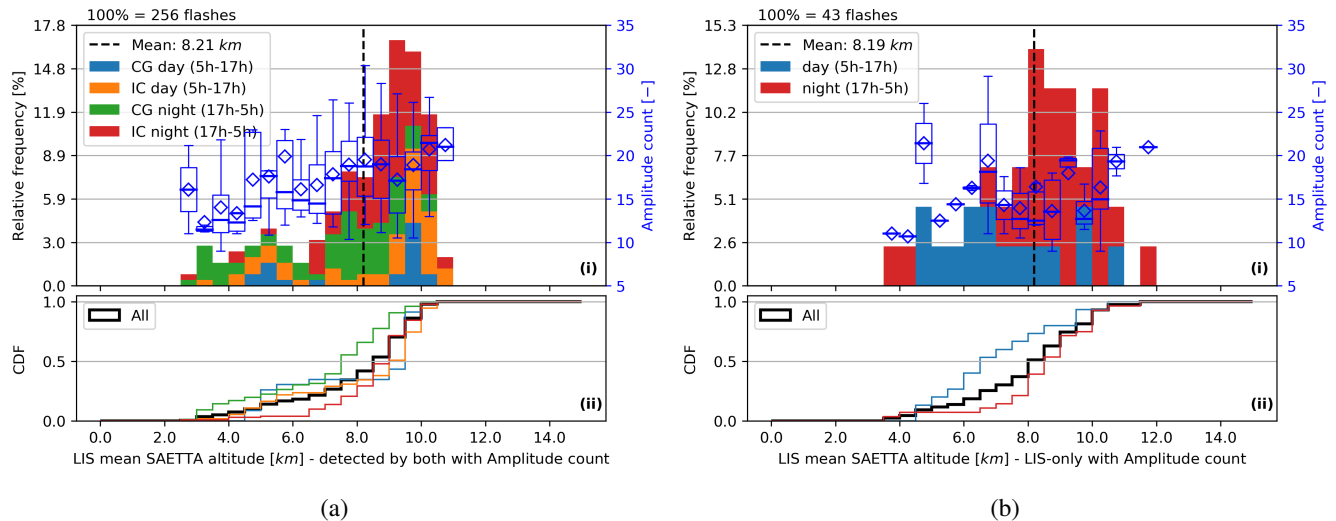
Figure 10. Pulse/Stroke number of Meteorage flashes with coincident ISS-LIS flash (a) and unmatched Meteorage flashes (b). Daytime, nighttime and flash type, IC and CG, are indicated by the colors. Histogram (i) and corresponding CDF (ii) use the same colors. The CDF shows in addition a black curve for all data. The histogram bin width is constant at 2 pulses/strokes. The mean value is plotted as dashed line. The total number of flashes is indicate above the histogram.

Figure 11. Flash extent of Meteorage flashes with coincident ISS-LIS flash (a) and unmatched Meteorage flashes (b). Daytime, nighttime and flash type, IC and CG, are indicated by the colors. Histogram (i) and corresponding CDF (ii) use the same colors. The CDF shows in addition a black curve for all data. The histogram bin width is constant at 2.5 km. The mean value is plotted as dashed line. The total number of flashes is indicate above the histogram.
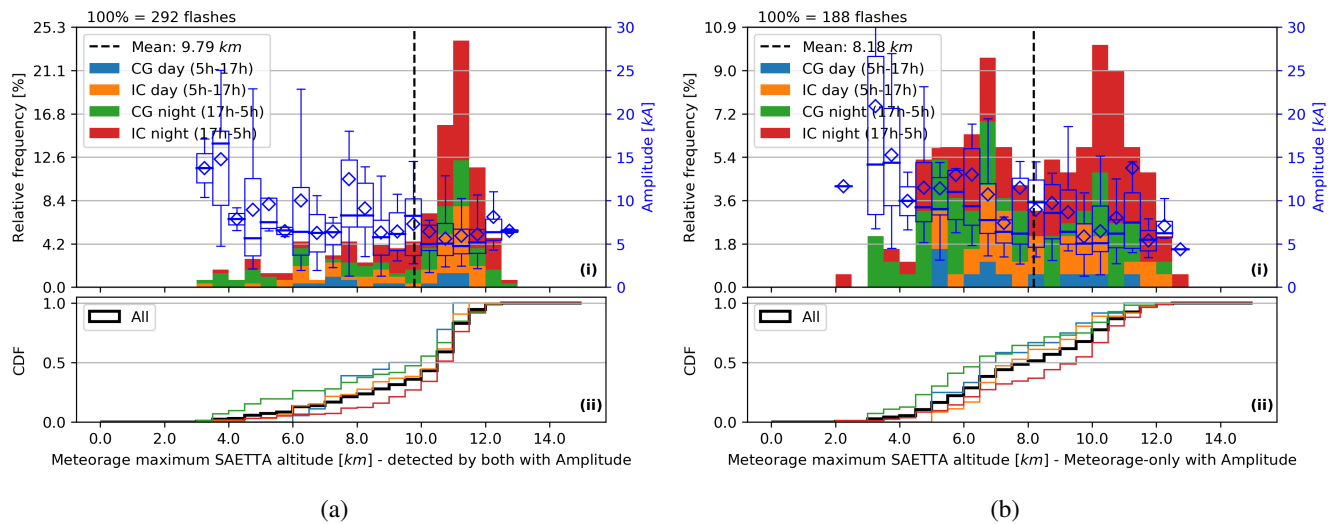
Figure 12. Meteorage flash duration with coincident ISS-LIS flash (a) and unmatched Meteorage flashes (b). Daytime, nighttime and flash type, IC and CG, are indicated by the colors. Histogram (i) and corresponding CDF (ii) use the same colors. The CDF shows in addition a black curve for all data. The histogram bin width is constant at 0.05 s. The mean value is plotted as dashed line. The total number of flashes is indicate above the histogram.

Figure 13. Flash mean altitude of ISS-LIS flashes (from concurrent SAETTA observations) with coincident Meteorage flash (a) and unmatched ISS-LIS flashes (b). Daytime, nighttime and flash type, IC and CG, are indicated by the colors. Histogram (i) and corresponding CDF (ii) use the same colors. The CDF shows in addition a black curve for all data. The histogram bin width equals 0.5 km. Note: The CG/IC attribute for ISS-LIS flash needs the matched Meteorage flash and does not exist for ISS-LIS-only flashes. The mean value is plotted as dashed line. The total number of flashes is indicate above the histogram. The blue boxplots (median as line, mean as diamond, Inter Quartile Range -IQR- as box, 1.5 IQR as whiskers - outliers not plotted) represent the distributions of ISS-LIS mean event ampltiude count per flash for each altitude bin.

Flash mean altitude of ISS-LIS flashes (from concurrent SAETTA observations) with coincident Meteorage flash (a) and unmatched ISS-LIS flashes (b). Daytime, nighttime and flash type, IC and CG, are indicated by the colors. Histogram (i) and corresponding CDF (ii) use the same colors. The CDF shows in addition a black curve for all data. The histogram bin width equals 0.5 . Note: The CG/IC attribute for ISS-LIS flash needs the matched Meteorage flash and does not exist for ISS-LIS-only flashes. The mean value is plotted as dashed line. The total number of flashes is indicate above the histogram. The blue boxplots (median as line, mean as diamond, Inter Quartile Range -IQR- as box, 1.5 IQR as whiskers - outliers not plotted) represent the distributions of ISS-LIS mean event radiance per flash for each altitude bin.

Figure 14. Flash maximum altitude of Meteorage flashes (from concurrent SAETTA observations) with (a) and without (b) coincident ISS-LIS flash. Daytime, nighttime and flash type, IC and CG, are indicated by the colors. Histogram (i) and corresponding CDF (ii) use the same colors. The CDF shows in addition a black curve for all data. The histogram bin width equals 0.5 km. The mean value is plotted as dashed line. The total number of flashes is indicated above the histogram. The blue boxplots (median as line, mean as diamond, IQR as box, 1.5 IQR as whiskers - outliers not plotted) represent the distributions of Meteorage mean absolute amplitude per flash for each altitude bin (scale fixed range from 0 to 30 kA).

~~As Figure 11 for the maximum altitude of Meteorage flashes with (a) and without match (b).~~
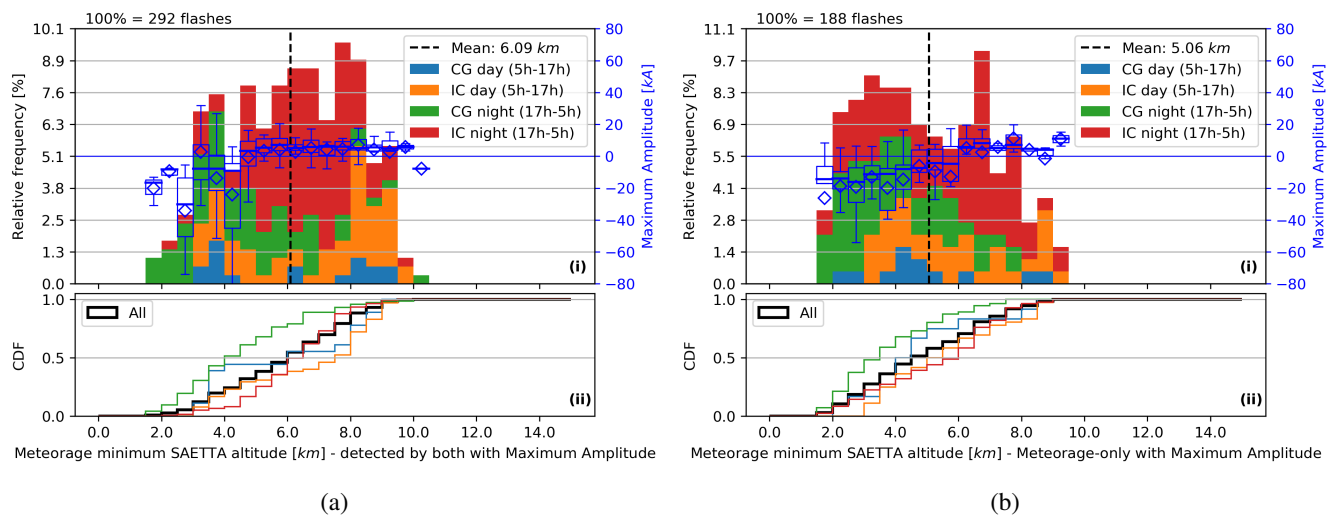
Figure 15. As Figure 14 for the minimum altitude of Meteorage flashes with (a) and without (b) match. Here, the blue boxplots represent the distributions of Meteorage maximum (pulse/stroke) amplitude per flash (positive or negative currents) for each altitude bin.

As Figure 11 for the minimum altitude of Meteorage flashes with (a) and without (b) match. Here, the blue boxplots (median as line, mean as diamond, IQR as box, 1.5 IQR as whiskers - outliers not plotted) represent the distributions of Meteorage maximum (pulse/stroke) amplitude per flash (positive or negative currents) for each altitude bin.