Prediction Model for Diffuser Induced Spectral Features in Imaging Spectrometers

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Abstract. Wide-field spectrometers for Earth Observation missions require inflight radiometric calibration, for which the sun can be used as a known reference. Therefor a diffuser is placed in front of the spectrometer in order to scatter the incoming light into the entrance slit and provide homogeneous illumination. The diffuser however, introduces interference patterns known as speckles into the system, yielding potentially significant intensity variations at the detector plane, called Spectral Features.

There have been several approaches implemented to characterize the Spectral Features of a spectrometer, e.g. end-to-end measurements with representative instruments. Additionally, in previous publications a measurement technique was proposed, which is based on the acquisition of monochromatic speckles in the entrance slit following a numerical propagation through the disperser to the detection plane. Based on this measurement technique we present a standalone prediction model for the magnitude of Spectral Features in imaging spectrometers, requiring only few input parameters and therefor mitigating the need for expensive measurement campaigns.

1   Introduction

Many current and future Earth Observation missions carry wide field spectrometer payloads such as the ENVISAT Medium Resolution Imaging Spectrometer (Olij et al. (1997)), the Sentinel-2 MultiSpectral Imager (Martimort et al. (2012)), the Sentinel-3a Ocean and Land Colour Imager (Nieke and Mavrocordatos (2017)), the Sentinel-4 UVN instrument (Clermont et al. (2019)), the Sentinel-5-UVNS instrument (Guehne et al. (2017)), or the GHGIS instrument of CO2M or former CarbonSat (Fletcher et al. (2015)). These space based instruments require inflight radiometric calibration, for which the Sun can be used as a known reference. In order to ensure homogeneous illumination of the instrument a diffuser is used to scatter the incoming sunlight into the entrance slit. However, the diffuser introduces a interference pattern known as speckles into the optical system. The speckles propagate through the disperser and are integrated at the detector plane, yielding intensity variations described as *Spectral Features* by van Brug et al. (2004). When the solar spectrum is used for the calibration of the top of atmosphere measured spectra, the Spectral Features create a radiometric error as they alter the calibration function. Speckle effects are commonly known from applications involving highly coherent laser light such as holographic imaging (Bianco et al. (2018)) or Laser Speckle Contrast Imaging (Heeman et al. (2019)). In general, diffuse sunlight does not yield a significant net speckle pattern when incident on a detector due to its broad spectrum. However, for spectrometers with a

fine spectral resolution, the spectral width per channel is narrow enough to give rise to an amplitude of the Spectral Features, that can be as large as other radiometric errors, e.g. due to straylight and polarisation. Spectral Features depend on various geometric conditions implying that their exact position at the detector plane cannot be reliably predicted. This also renders any mitigating post-processing steps ineffective. Widely known speckle suppression techniques, such as rotation or tilting of elements in the optical system (Goodman, 2007, section 5), are only viable for on-ground calibration with a static setup. For space applications additional moving parts are typically not implemented, because they pose a supplementary risk of failure. In early planning phases, each contribution to the radiometric error needs to be quantitatively estimated to allow a global optimization of the instrument. This work presents a novel method to predict the radiometric error due to Spectral Features. As in most applications, the radiometric error is driven by the amplitude of the Spectral Features, van Brug and Courrèges-Lacoste (2007) introduced the Spectral Features Amplitude (SFA) as a standardized measure.

The issue of diffuser induced Spectral Features in imaging spectrometers has first been documented by Richter et al. (2002) and Wenig et al. (2004) in the context of the Global Ozone Monitoring Experiment (GOME). Ahlers et al. (2004) observed spectral oscillations caused by the onboard diffuser in the Scanning Imaging Absorption Spectrometer for Atmospheric Chartography (SCIAMACHY) instrument, as well as van Brug et al. (2004). Several approaches for characterization and modeling have been proposed since. van Brug and Courrèges-Lacoste (2007) presented an end-to-end measurement setup featuring a complete instrument including source, telescope optics, disperser, and detector. However, since Spectral Features need to be suppressed "by design" of certain instrument parameters a quantitative analysis with a representative spectrometer is usually difficult to achieve. van Brug and Scalia (2012) introduced models for different speckle averaging effects. However, a comprehensive and reliable model has not been presented, yet. Isolating the Spectral Features by eliminating all other error sources in an representative end-to-end setup remains the main challenge to gain quantitative insights into the SFA dependence.

A different approach to quantify Spectral Features was taken by Burns et al. (2017) and improved by Richter et al. (2018). It is based on the subsequent acquisition of monochromatic speckle patterns in the slit plane over several spectral channels, which are then propagated numerically through the disperser to the detection plane. Some simplifying assumptions are made about the optical system which reduces the complexity and limits systematic error contributions. It is only limited by the SNR and the resolution achieved in the entrance slit and therefor capable of yielding comprehensive measurement data for most instrument designs. Especially, it allows for a step-by-step tracing of the speckle statistics from the slit to the detector plane.

Based on this SFA measurement technique we present a novel standalone SFA prediction model, which solely relies on mathematical descriptions of the speckle statistic and its SFA impact. It includes polarization effects of the diffuser, spatial and spectral averaging as well as pixel averaging at the detector.

First we review the definition of the Spectral Features Amplitude (SFA) in Sect. 2. In Sect. 3 the revised SFA measurement technique used for our measurements is shown. We then present the standalone SFA prediction model in Sect. 4, which can be understood as a mathematical formulation of the SFA measurement technique. Finally, we compare the results of the prediction model to our measurement chain in detail in Sect. 5 to show its validity. In the last section we discuss the applicability to a real instrument.

2 Spectral Features Amplitude

The term *Spectral Features Amplitude* (SFA) was first proposed by van Brug and Courrèges-Lacoste (2007) as standardized way to quantify diffuser induced "wiggles" in a spectrum measured by a space spectrometer instrument. They describe it as the magnitude of the "features" in a spectrum that are solely caused by the diffuser altering the solar reference spectrum. The SFA value is calculated as the standard deviation of the mean normalized detector signal over multiple spectral channels (see details in Sect. (3.1)). The SFA value holds information about the amplitude of features. However, the data produced in this work, which is used to calculate the SFA, also allows for the estimation of the spectral extend of features. One usually may not draw conclusions about the absolute spectral positions of features with this approach. We will show that the instrument parameters used in this work lead to a spectral speckle extend smaller than the instrument detector pixel. This essentially allows for the treatment of the SFA as white noise at the detector level.

3 SFA Measurement Chain

In this Sect. the used SFA measurement technique introduced by Burns et al. (2017) is presented in a revised state. The goal of this technique is the reduction of experimental complexity and therefore systematic error contributions during data acquisition. First, the measurement principle is explained. Second, the used materials and the measurement procedures for the near infrared (NIR) and the short wavelength infrared (SWIR) channel are presented.

3.1 Principle

Figure 1 depicts the optical setup of an imaging spectrometer during solar calibration. The incoming sunlight is scattered by the diffuser. The scattering origin lies in the aperture plane with spatial coordinates g and h which is perpendicular to the light's direction of propagation. The angular field distribution at the aperture plane is imaged to the slit plane with coordinates x and y. The light is collimated onto a dispersive element (e.g a diffraction grating), which is splitting it into its spectral components. The spatial information in the y direction of the slit is transformed into spectral information at the detector with coordinate b by imaging the the diffracted beams of different wavelengths onto a 2D array detector. Beams of the same wavelength (within the spectral resolution) are assigned the same spectral detector coordinate b. The spatial information in x direction is retained in the detector coordinate a. We relate the coordinates via the simplified linear spectrometer equations

$$a = M_x x, \tag{1}$$
$$b = M_y y + k\lambda, \tag{2}$$

where M_x and M_y are the respective magnification factors in x and y direction, $k = db/d\lambda$ denotes the dispersion, and λ the wavelength. For these simplified equations to hold the magnification factors and the dispersion are assumed to be independent of the wavelength and the spatial position (x, y). Also, the instrument point spread function (IPSF) is not accounted for. Sunlight is spatially coherent (Agarwal et al. (2004)) and we can assume collimated light illuminating the diffuser. Additionally, the
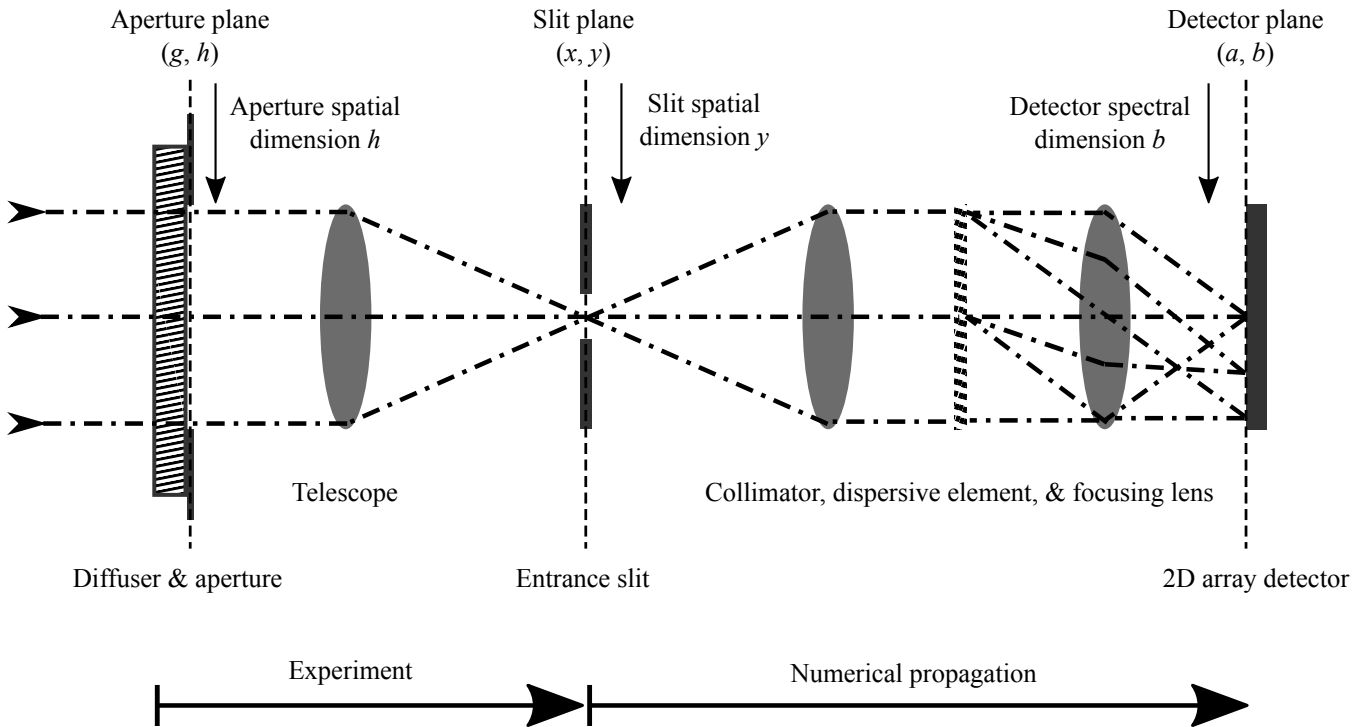
Figure 1. Optical setup of an imaging spectrometer during solar calibration. The sequence of optical components is subdivided into two parts. The first part is covered by the experimental setup in the lab starting at the illuminated diffuser and ending at the slit in the telescope focal plane. The second part numerically propagates the images recorded in the slit plane to the instrument focal plane.

sunlight's temporal coherence is on the order of femtoseconds (Hecht and Lippert (2018)) and much smaller than the typical
90   detector integration time of several hundred milliseconds of an optical instrument. As a consequence, cross coherence terms of interfering fields of different wavelengths average out and the net intensity distributions at the slit and the detector planes are well approximated by the superposition of monochromatic intensities. The Sun disk comprises of many incoherent point sources, which should be considered for angular averaging contributions and is not part of this work, as it will only account for a single point source. For the purpose of the SFA measurement the sequence of optical components is subdivided into
95   two parts. The first part ranging from the illuminated diffuser through the telescope to the entrance slit is represented by the optical setup in the lab. The second part comprises the rest of the optical system from the slit plane to the instrument detector plane. The data acquired in the first part is used as input for a numerical simulation of the optical setup after the slit plane. The setup layout is shown in Fig. 3. The Sun is mimicked as a single field point with a tunable laser source, which is spectrally stabilized by a wavemeter, and illuminates the diffuser through a linear polarizer at normal incident with respect to the diffuser
100   plane. The distance between the single mode fiber output and the diffuser is chosen such that the divergent beam illuminates the diffuser homogeneously over the size of the apertures. The second aperture blocks any unwanted angular contributions. A

powermeter placed next to the diffuser records a fixed fraction of the emitted laser power. The telescope images the scattered light onto a 2D array detector positioned in the focal plane. The focal plane of the telescope represents the slit plane in Fig. 1. The diffuser plane is tilted by 10° with respect to aperture and slit plane. This ensures that only scattered light contributes to the measurement. The telescope is aligned perpendicular to the aperture and slit plane.
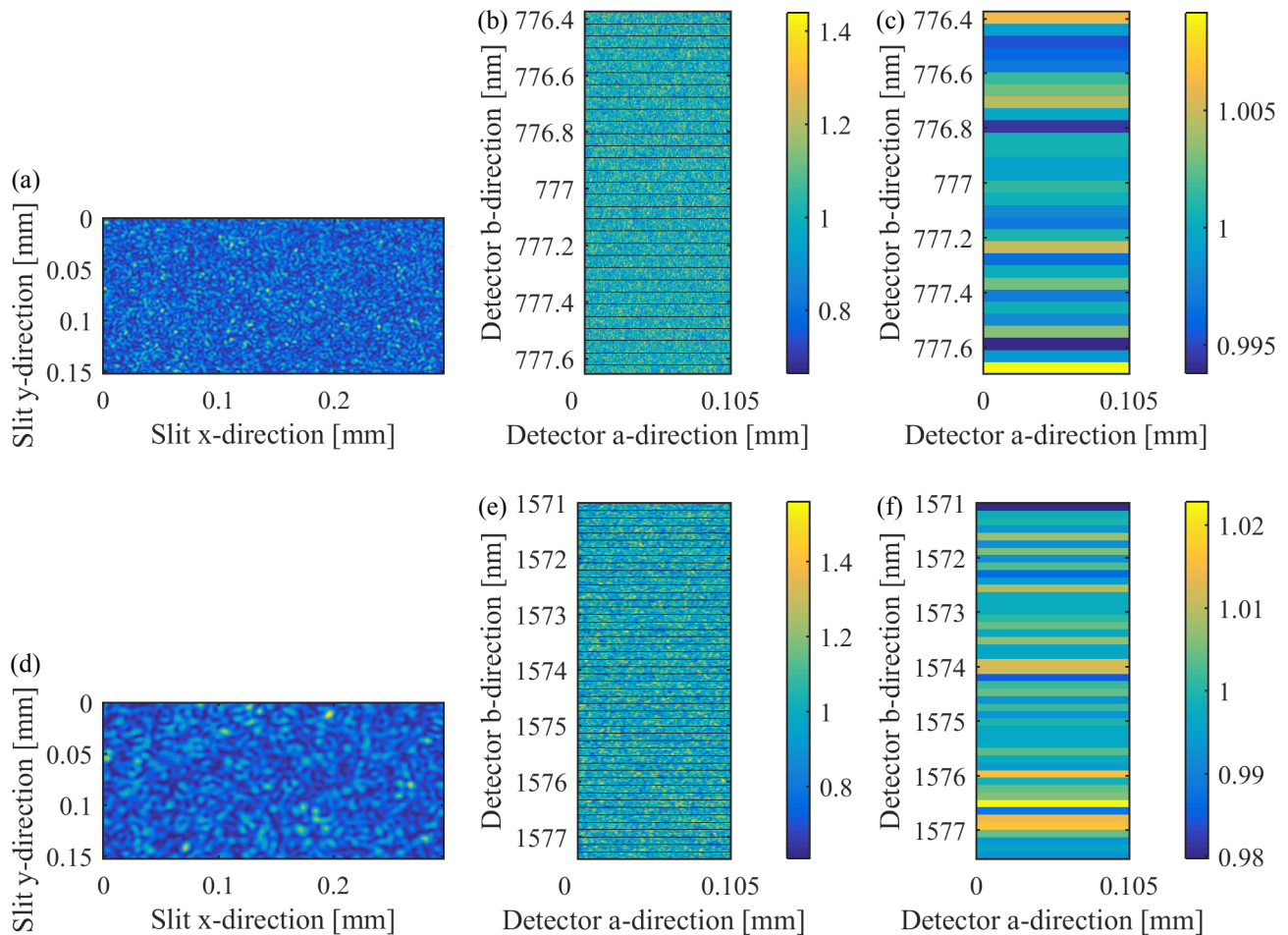
Figure 2. Speckle patterns in the NIR band (a-c) and SWIR band (d-f) at different stages in the measurement chain: (a/d) is an example of a monochromatic speckle pattern in the slit plane, (b/e) is the speckle pattern integrated at the detector plane using Eq. (3) and normalized to its mean, where the horizontal lines denote the instrument detector pixel grid $(\tilde{a}, \tilde{b})$, and (c/f) is the final normalized integrated detector signal. The standard deviation taken over the pixel rows is the SFA.

For a measurement, monochromatic speckle intensities are recorded subsequently over a wavelength range $\lambda_1...\lambda_N$ several times the spectral resolution λ_{res} of the real spectrometer that is being mimicked. An example for such an monochromatic speckle pattern is shown in Figure 2 (a) and (d) in the slit plane. The spectral tuning step size $\Delta\lambda$ in between images needs to be

sufficiently small, in order to properly sample the change of the speckle patterns. The intermediate result is a three dimensional data set $I_{slit}(x,y,\lambda)$ consisting of a spectrum of monochromatic speckle images, where x and y are the spatial coordinates in the slit plane and λ is the wavelength. Every speckle image is mapped to a certain position (a,b) at the focal plane, where all images are summed up in intensity. The summation on intensity basis is justified as cross coherence terms involving interference of different wavelengths of actual sunlight will vanish for sufficiently long integration times. The summation procedure is detailed in Burns et al. (2017) and can be summarized as

$$I_{det}(a,b) = \frac{\Delta\lambda}{\lambda_{res}} \sum_{\lambda=\lambda_1}^{\lambda_N} I_{slit}\left(\frac{a}{M_x}, \frac{b-k\lambda}{M_y}, \lambda\right) \Theta(b-k\lambda), \tag{3}$$

where slit coordinates are expressed in terms of the detector coordinates using Eq. (1) and (2) and the Heavyside function with $\Theta(y)=0, y<0$ and $\Theta(y)=1, y\geq0$. The result of the sum is a two dimensional intensity distribution in the focal plane of the instrument $I_{det}(a,b)$, which is depicted in Figure 2 (b) and (e). In a last step $I_{det}(a,b)$ is overlayed with the actual instruments detector pixel grid $(\tilde{a},\tilde{b})$ and intensities belonging to the same pixel are summed, which is shown in Figure 2 (c) and (f). The SFA is calculated as standard deviation of the mean normalized detector pixel intensity distribution $I_{det,binned}(\tilde{a},\tilde{b})$.

3.2 Materials and Procedure

Measurements are conducted in the NIR regime around 777 nm and in the SWIR regime around 1570 nm which represent wavelength bands with commonly monitored data products, such as water vapour, clouds, CO_2, aerosols, or the O_2 absorption which is commonly used to calculate the effective path length and the air mass factor (see Irizar et al. (2019), or Meijer et al. (2019), or Voors et al. (2017)). The experimental setup is shown in Sect. 3.1. As light sources serve tunable monochromatic external cavity diode lasers with single mode output and an integrated optical isolator. They are stabilized via a proportional-integral-derivative (PID) loop with feedback data from a wavemeter, which uses a Fizeau interferometer. As connecting fibers between laser, wavemeter, and output serve single mode (SM) fibers, since the spectral tuning range will be narrow. Also SM fiber will introduce no additional speckle in contrast to multi mode fibers. A linear polarizer ensures polarization stability. The round diffuser plate has a diameter of 70 mm and a thickness of 3 mm. It is made of highly scattering fused silica HOD®-500 material featuring inhomogeneities of 20 microns or less. The spatial coherence length of the Sun's light is around 60 microns in the NIR and 120 microns in the SWIR band, according to Divitt and Novotny (2015). Hence, the material, which was selected for the Sentinel-5/UVNS according to Irizar et al. (2019), is suited for the wavelengths in question and deemed a good choice for this study. The data collected with the powermeter is used to normalize the acquired images. The round apertures are used to control the size of individual speckle correlation areas. The laser beam's uniformity at the diffuser plane was measured to be around 3% in the NIR band and 6% in the SWIR band over the size of the aperture. There were no additional speckle contribution by the fibers detected. The telescope has a focal length of $f_{tel} = 1100mm$. For the NIR the laser source has a center wavelength of 780 nm and a nominal linewidth of 6×10^{-7}nm. The Thorlabs 780HP SM fiber type was used for all fiber connections in the NIR range. The CCD detector features a 12.5 mm x 10.0 mm active area with 2750 x 2200 pixels of size 4.54 μm x 4.54 μm. The expected average one dimensional width of the speckle correlation area S_d in the slit plane is

calculated using

$$S_d = \frac{2\lambda f_{tel}}{D\sqrt{\pi}},\tag{4}$$

which was derived by Goodman (2007). The apertures' diameter is set to $D = 10$mm, and thus the speckles are sampled by 19 pixel in one dimension, which is deemed sufficiently fine to ensure that the sampling with the detector does not alter the speckle pattern significantly. The laser wavelength was tuned over the range of 776.4 nm-777.7 nm with a step size of $\Delta\lambda = 1$pm. The step size is chosen, so that there is a non zero correlation between subsequent speckle images. For the SWIR measurement the laser source as well as the detector and the fiber splitter are replaced and the Thorlabs SMF-28 SM fiber was used for fiber connections. The SWIR laser source center wavelength is 1550 nm, with single mode output of nominal 2×10^{-6}nm linewidth. The detector is a 640 x 512 pixel InGaAs camera with a pixel size of 15.5 μm x 15.5 μm. The tuning range was 1571nm-1577.5nm with a step size of $\Delta\lambda = 3.1$pm. The apertures' diameter is set to $D = 13$mm which means a sampling of 10 pixel per speckle in one dimension. This is smaller than for the NIR configuration, but is a trade off between sampling and SNR.
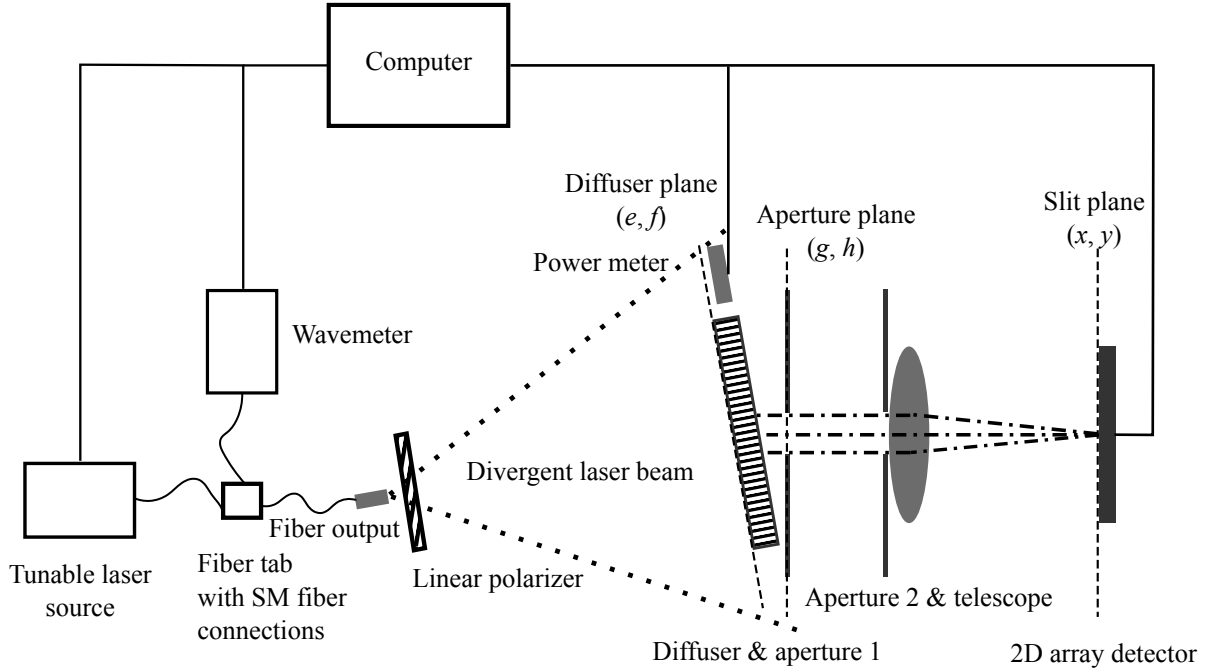
Figure 3. Layout of the experimental setup for measuring diffuser induced monochromatic speckle patterns in the slit plane. Single mode fibers are used to connect the laser source with the wavemeter and the output via a fiber tab and indicated by curved lines.

4 Spectral Features Amplitude Prediction Model

The prediction model presented in the following is a mathematical formulation of the measurement method described in Sect. 3. It relies on the determination of the speckle statistics at different steps of the measurement chain. The relevant physical information about speckle averaging effects in the measurement chain lies in the intensity distributions. A single pattern I is sampled by a finite but sufficient amount of pixel, so that the individual pixel size is small compared to the speckle size. The magnitude of the speckle effect in I is described as the *speckle contrast* (Goodman, 2007, p. 28)

$$C = \frac{\sigma_I}{\langle I \rangle}, \tag{5}$$

where σ_I is the standard deviation and $\langle I \rangle$ is the mean value of I over all pixel. Under the general assumption, that the individual statistics of the underlying fields are circular complex Gaussian, a fully developed speckle pattern generated with linear polarized monochromatic light has a contrast of $C = 1$ (Goodman, 2007, p. 29). We adopt this assumption for this model. The speckle contrast is reduced by several averaging effects introduced by the spectrometer instrument. A reduction of C is only achieved by the summation of intensity distributions showing a correlation smaller than unity. If the summation is on amplitude basis (when the distributions can interfere), C is not reduced (Goodman, 2007, section 3.1.1). From this follows, that only distributions which can not interfere will impact C and are therefore subject of further discussions. Each one of the N independent averaging effects attributes to a certain amount of *degrees of freedom M_n* or effectively uncorrelated intensity distributions, which can be combined to a total averaging factor M (Goodman, 2007, p. 186) by

$$M = \prod_n^N M_n. \tag{6}$$

The reduced speckle contrast will then calculate to

$$C_{reduced} = \frac{1}{\sqrt{M}} = \left(\sqrt{\prod_n^N M_n} \right)^{-1}. \tag{7}$$

In order to predict the contrast reduction we identified $N = 3$ contributors, which can be assigned to different steps of the SFA measurement chain:

1. Generation of monochromatic diffuse depolarized light in the aperture plane (g, h) yielding a factor $M_{polarization}$,

2. mapping intensities in the slit plane (x, y) to instrument detector positions (a, b) contributing a factor $M_{spectral}$,

3. and integration of the instrument detector pixels with a factor denoted by $M_{detector}$.

The predicted reduced speckle contrast at the instrument detector plane using Eq. (7) corresponds to the SFA and is

$$SFA = C_{reduced} = \frac{1}{\sqrt{M_{polarization} M_{spectral} M_{detector}}}. \tag{8}$$

4.1 Polarization Averaging

The laser source emits a single polarization state, which is ensured with the polarizer. The diffuse light leaving the volume diffuser can be treated as depolarized due to multi scattering (Lorenzo, 2012, p.85). This corresponds to two orthogonal polarization configurations or two effective intensity distributions which can not interfere. Therefore step $n = 1$ introduces two degrees of freedom $M_{polarization} = 2$ (Goodman, 2007, p. 49).

4.2 Spectral Averaging

Step $n = 2$ leads to spectral averaging at the detector. We recall the finding from Sect. 3.1, that the net intensities in the field planes (slit and detector) can be treated as superposition of monochromatic intensities for integration times much greater than the coherence time. Let us consider the acquired speckle intensities $I_n(x, y)$ and the underlying fields $A_n(x, y)$, which are related by $I_n = |A_n|^2$. They are recorded at various wavelengths λ_n. The magnitude of the statistical change of subsequent speckle intensities I_m and I_n with a wavelength difference $\Delta\lambda_{nm} = |\lambda_n - \lambda_m|$, can be described in terms of the first order field correlation coefficient μ_{mn}, with

$$\mu_{mn}(\lambda_n, \lambda_m) = \frac{\langle A_m A_n^* \rangle}{\sqrt{\langle I_m \rangle \langle I_n \rangle}}, \tag{9}$$

where $*$ denotes the complex conjugate. The field correlation is influenced by two effects, which in our case are both frequency dependent. The first effect is due to changing light paths through the diffuser medium. The second effect takes into account the spatial offset $\Delta b = k\Delta\lambda_{nm}$ at the detector plane between individual speckle patterns I_n induced by the dispersion (see Eq. (2)). We start with the former contribution to the field correlation and follow the approach of Zhu et al. (1991), who presented an analytic equation for the angular correlation function of a slab geometry of a scattering media of thickness d, which can also be used for wavelength correlations:

$$F(\lambda_n, \lambda_m) = \frac{(d + 2B)/(z_0 + B)\left[\sinh(z_0\sqrt{q^2 + \alpha^2}) + B\sqrt{q^2 + \alpha^2}\cosh(z_0\sqrt{q^2 + \alpha^2})\right]}{[1 + B^2(q^2 + \alpha^2)]\sinh(d\sqrt{q^2 + \alpha^2}) + 2B\sqrt{q^2 + \alpha^2}\cosh(d\sqrt{q^2 + \alpha^2})}, \tag{10}$$

where we set $q = \sqrt{i6\pi\left|\frac{1}{\lambda_n} - \frac{1}{\lambda_m}\right|\beta n_s/l_t}$, with n_s denoting the refractive index of the scattering material, l_t the transport mean free path used for anisotropic multi scattering systems, and β a constant factor taking into account the contribution of the tilted diffuser plane (e, f) with respect to the other planes and the specific geometry. The value z_0 describes the average penetration depth after which the light is scattered for the first time. It does not have a great influences in a transmission geometry and is approximated with the transport mean free path l_t. We set $\alpha = 0$, thereby ignoring absorption. The parameter for the boundary condition is given by $B = l_t\frac{2(1+R)}{3(1-R)}$, where R is the reflection coefficient which is calculated using the Fresnel equations. It accounts for internal reflection due to index of refraction mismatch at the boundaries.

The second contribution to the field correlation is due to changing spatial positions of speckle patterns which are distributed over the instrument detector in accordance with the spectral dispersion. This constitutes a spatial offset Δb between the speckle intensities I_n and I_m at the detector plane (a, b). Keeping in mind Eq. (1) and (2) for the transformation from the slit to the

detector plane, the correlation of speckle fields, which are separated spatially by Δb can be expressed as

$$
\left.\Psi(\Delta a, \Delta b)\right|_{\Delta a=0} = \left.\frac{\int_{-\infty}^{\infty} |P(g,h)|^2 \, e^{-i\frac{2\pi}{\tilde{\lambda} z}(g\Delta a + h\Delta b)} \, dg \, dh}{\int_{-\infty}^{\infty} |P(g,h)|^2 \, dg \, dh}\right|_{\Delta a=0}, \tag{11}
$$

where $P(g,h)$ is the aperture function of the imaging system and may be defined, e.g. for a circular aperture of diameter D, as $P(g,h) = circ\left(\frac{2\sqrt{g^2+h^2}}{D}\right)$ which is 1 inside the aperture and 0 otherwise, g and h are the a- and b-coordinate representations in the aperture plane, z is the distance between aperture and slit plane, and $\tilde{\lambda}$ the mean wavelength (Goodman, 2007, p. 169). We set $\Delta a = 0$ since we are only interested in the offset in the spectral direction. Combining the two effects we can model the correlation between the speckle fields as

$$
\mu_{mn}(\lambda_n, \lambda_m) = F(\lambda_n, \lambda_m) \, \Psi(\Delta b). \tag{12}
$$

The accumulation of individual speckle patterns I_n with field correlations μ_{mn} at the detector can be interpreted as the summation of partially correlated speckle intensities

$$
I_{det}(a,b) = \sum_{n=1}^{N=\lambda_{res}/\Delta\lambda} I_n\left(\frac{a}{M_x}, \frac{b-k\lambda}{M_y}, \lambda_n\right). \tag{13}
$$

The amount of individual speckle intensities I_n contributing to the sum at arbitrary detector coordinates (a,b) is equal to the ratio of the spectral resolution λ_{res} with the step size $\Delta\lambda$. This also applies to the mean intensities,

$$
\langle I_{det}(a,b) \rangle = \sum_{n=1}^{N=\lambda_{res}/\Delta\lambda} \langle I_n\left(\frac{a}{M_x}, \frac{b-k\lambda}{M_y}, \lambda_n\right) \rangle. \tag{14}
$$

Using an established method shown by Bevan (2009) and Goodman (2007) we define a coherency matrix with entries $J_{nm} = \langle \boldsymbol{A}_m \boldsymbol{A}_n^* \rangle$ and use Eq. (9) to get

$$
\mathbf{J} = \begin{bmatrix}
\langle I_1 \rangle & \sqrt{\langle I_1 \rangle \langle I_2 \rangle}\mu_{1,2} & \cdots & \sqrt{\langle I_1 \rangle \langle I_N \rangle}\mu_{1,N} \\
\sqrt{\langle I_1 \rangle \langle I_2 \rangle}\mu_{1,2}^* & \langle I_2 \rangle & \cdots & \sqrt{\langle I_2 \rangle \langle I_N \rangle}\mu_{2,N} \\
\vdots & \vdots & \ddots & \vdots \\
\sqrt{\langle I_1 \rangle \langle I_N \rangle}\mu_{1,N}^* & \sqrt{\langle I_2 \rangle \langle I_N \rangle}\mu_{2,N}^* & \cdots & \langle I_N \rangle
\end{bmatrix}. \tag{15}
$$

By diagonalization of $\mathbf{J}$ with a unitary linear transformation $\mathbf{L}_0$, the ensemble of correlated speckle fields is transformed to a basis with no correlation between them.

$$
\mathbf{J}' = \mathbf{L}_0 \mathbf{J} \mathbf{L}_0^\dagger = \begin{bmatrix}
\langle \tilde{I}_1 \rangle & 0 & \cdots & 0 \\
0 & \langle \tilde{I}_2 \rangle & \cdots & 0 \\
\vdots & \vdots & \ddots & \vdots \\
0 & 0 & \cdots & \langle \tilde{I}_N \rangle
\end{bmatrix}, \tag{16}
$$

where $\dagger$ denotes the Hermitian transpose operation. The total mean intensity $\langle I_{det} \rangle = \sum_n \langle I_n \rangle = \sum_n \langle \tilde{I}_n \rangle$ is conserved under this transformation but in general $\langle I_n \rangle \neq \langle \tilde{I}_n \rangle$. The complex coherence factor $\mu_{mn} = |\mu_{mn}| \exp(i\Phi_{nm})$ includes a phase Φ_{nm}.

However, due to the specific construction of $\mathbf{J}$, these phase terms can be omitted when calculating the eigenvalues (Dainty et al., 1975, section 4.7.2). Finally, for the spectral degrees of freedom we use the eigenvalues $\langle \tilde{I}_n \rangle$ of the coherency matrix to get

$$M_{spectral} = \left(\frac{\langle I_{det} \rangle}{\sigma_{det}} \right)^2 = \frac{\left(\sum_n \langle \tilde{I}_n \rangle \right)^2}{\sum_n \langle \tilde{I}_n \rangle^2} . \tag{17}$$

235 Note that changing $\Delta\lambda$ and thus changing N will not change the result of $M_{spectral}$ as long as $\Delta\lambda$ is sufficiently small to sample the covariance μ_{mn}. The enabling property of the coherency matrix $\mathbf{J}$ is called *Toeplitz*, which implies an asymptotically behavior of its eigenvalues found by Grenander and Szegö (1958). Gray (2006) gives a simplified prove in Corollary 2.1 and 2.2 that both, numerator and denominator in Eq. (17) converge for large N.

4.3 Detector Averaging

240 In step $n = 3$ an averaging due to the integration of the instrument detector pixel takes place. We already established, that the resultant intensity distribution at the detector $I_{det}(a, b)$ is given by the summation in Eq. (13). This effect impacts the speckle contrast if individual speckles are not sufficiently oversampled by the instrument detector pixel grid $(\tilde{a}, \tilde{b})$. An analytical expression for the degrees of freedom $M_{detector}$ introduced by stationary speckles in one detector pixel with relative coordinates $(\Delta a, \Delta b)$ is given by

245
$$M_{detector} = \left[\frac{1}{A_D^2} \iint\limits_{-\infty}^{\infty} K_D(\Delta a, \Delta b) \left| \mu_{det}(\Delta a, \Delta b) \right|^2 d\Delta a\, d\Delta b \right]^{-1} , \tag{18}$$

where A_D is the area of a detector pixel, $K_D(\Delta a, \Delta b)$ is the autocorrelation function of the detector pixel, and $\mu_{det}(\Delta a, \Delta b)$ is the field correlation at the detector plane (Goodman, 2007, p. 108). In order to accurately describe μ_{det} one needs to account for the evolution of the speckle size during the summation in Eq. (13). Let us consider a single speckle correlation area $I_1(S_d/2 \leq |x - x_1|, S_d/2 \leq |y - y_1|, \lambda_1)$ with a spatial extend denoted by S_d centered at (x_1, y_1) in the slit. Its correlation
250 relative to this position is described by

$$\Psi(\Delta x, \Delta y) = \frac{\int_{-\infty}^{\infty} |P(g, h)|^2 e^{-i\frac{2\pi}{\lambda z}(g\Delta x + h\Delta y)} dg\, dh}{\int_{-\infty}^{\infty} |P(g, h)|^2 dg\, dh} , \tag{19}$$

with $\Delta x = x - x_1$ and $\Delta y = y - y_1$ being relative coordinates. The function $F(\lambda_1, \lambda_n)$ introduced previously, characterizes how the correlation area develops after n wavelength steps at the same position, $I_n(S_d/2 \leq |\Delta x|, S_d/2 \leq |\Delta y|, \lambda_n)$ with $n > 1$. In other words, it denotes the amount of spectral steps after which a single speckle seizes to exist at a fixed position in the slit.
255 The initial position of the speckle at the detector is $(a_1, b_1) = (M_x x_1, M_y y_1 + k\lambda_1)$. The subsequent contributions relative to the initial position are shifted in the b−direction by $k|\lambda_1 - \lambda_n|$ and have a magnitude denoted by $F(\lambda_1, \lambda_n)$. Therefore, the resultant speckle correlation function at the detector $|\mu_{det}(\Delta a, \Delta b)|^2$ is a convolution of $|\Psi(\Delta a, \Delta b)|^2$ with $|F(\lambda_n, \lambda_m)|^2$:

$$|\mu_{det}(\Delta a, \Delta b)|^2 = |\Psi(\Delta a, \Delta b)|^2 \circledast |F(\lambda_n, \lambda_m)|^2 . \tag{20}$$

The symbol $\circledast$ denotes the convolution.

5 Results and Discussion

In the following we present and compare the SFA results from the measurement chain of Sect. 3 with the ones from the prediction model of Sect. 4 in the NIR and SWIR regime. $M_{spectral}$ can be interpreted as the average number of speckle patterns generated by the diffuser per spectral channel at the given wavelengths. $M_{detector}$ can be understood as average number of speckle correlation areas influencing the measurement in a detector pixel. The values of relevant parameters used are depicted in Table 1. They were chosen to represent a proposed instrument for ESA's CO2M mission (Meijer et al. (2019)). In Figure 4 the measured Pearson correlations between speckle patterns $I_n(\lambda_n)$ and $I_m(\lambda_m)$ with respect to their relative spectral distance is shown as blue stars for the NIR band at 776 nm and the SWIR band at 1571 nm, respectively. All wavelength shift combinations up to 0.1 nm are averaged for 120 images. Error bars are omitted, because the standard error of the mean is too small to be displayed. The red line denotes a fit to Eq. (10) leading to the values for l_t depicted in Table 1. They are in agreement with the supplier's given value of $l_t = 56\,\mu m$ at 500 nm wavelength.

Table 1. Sample spectrometer parameters used for the measurement and prediction. They were chosen to represent a proposed instrument for ESA's CO2M mission (Meijer et al. (2019)).

Parameter	Value
Magnification M_x	0.34
Magnification M_y	0.30
Aperture diameter	40.0 mm
Slit dimensions (x, y-direction)	295 μm, 152 μm
Detector dimensions (a, b-direction)	105 μm, 45 μm
Telescope focal length	131 mm
Diffuser thickness d	3 mm
NIR specific	
Spectral resolution λ_{res}	0.128 nm
Spectral tuning range $\lambda_1...\lambda_N$	776.4 nm - 777.7 nm
Tuning step size $\Delta\lambda$	1 pm
Refractive index of diffuser material $n_s(\lambda)$	1.454
Mean free path $l_t(\lambda)$	$(59.3 \pm 0.4)\,\mu m$
SWIR specific	
Spectral resolution λ_{res}	0.4 nm
Spectral tuning range $\lambda_1...\lambda_N$	1571 nm - 1577.5 nm
Tuning step size $\Delta\lambda$	3.1 pm
Refractive index of diffuser material $n_s(\lambda)$	1.444
Mean free path $l_t(\lambda)$	$(67.8 \pm 0.5)\,\mu m$
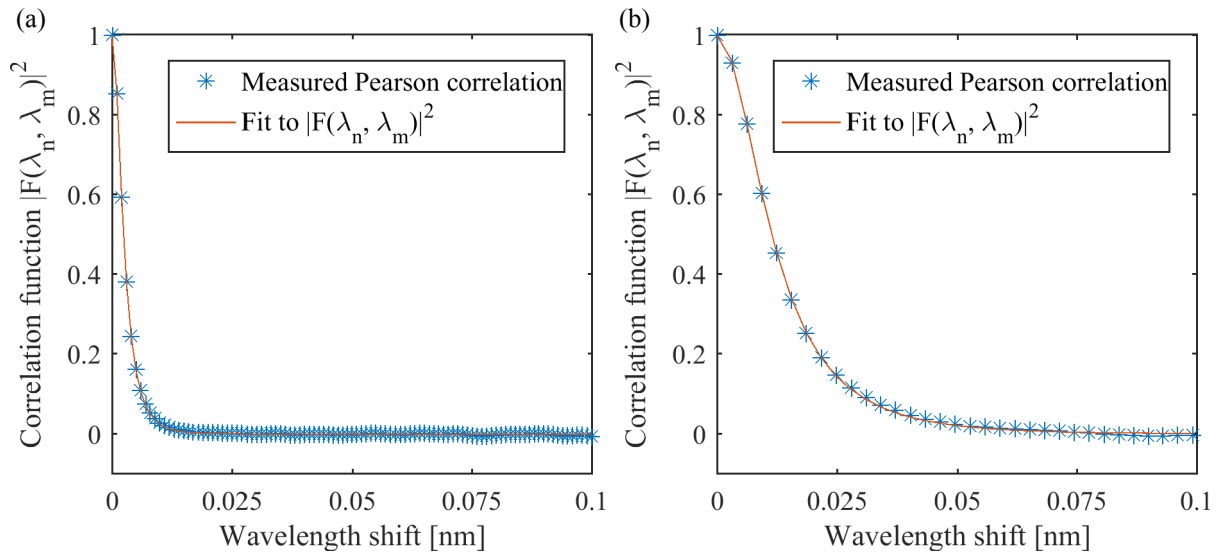
Figure 4. Measurement of the correlation function $|F(\lambda_n, \lambda_m)|^2$ in (a) the NIR (776 nm) and (b) the SWIR band (1571 nm). Blue stars denote the measured Pearson correlations between speckle patterns $I_n(\lambda_n)$ and $I_m(\lambda_m)$. All wavelength shift combinations up to 0.1 nm are averaged for 120 images. Error bars are omitted, because the standard error of the mean value is too small to be displayed. The red graph denotes the fit of the measured data points to Eq. (10) with (a) $l_t = (59.3 \pm 0.4)$ µm and (b) $l_t = (67.8 \pm 0.5)$ µm.

Table 2 shows the SFA values of the measurements and predictions with their corresponding intermediate averaging factors $M_{polarization}$, $M_{spectral}$, and $M_{detector}$ introduced in Sect. 4. Their counterparts from the measurement are deducted by calculating the speckle contrast at intermediate steps in the measurement chain. To verify the factor $M_{polarization} = 2$ a linear polarizer is placed *after* the diffuser and the measured speckle contrast compared to the nominal case rises by a factor of $\sqrt{2}$. Additionally, the polarization axis is rotated to different random positions without changing the result, which confirms the assumption made in Sect. 4, that the light exiting the diffuser is depolarized. With an ideal measurement chain the speckle contrast expected in the slit plane for monochromatic polarized speckles is $C_{slit,ideal} = 1$. This is the numerator in Eq. (8). The measured contrast in the slit is smaller, probably due to detector noise as suggested by Postnov et al. (2019). Webster et al. (2003) attributed the reduced measured contrast straylight from multiple reflections in their setup. The experimental factors $M_{spectral,measured}$ and $M_{detector,measured}$ are calculated in similar way using the relations,

$$C_{spectral,measured} = \frac{\langle C_{slit,measured} \rangle}{\sqrt{M_{spectral,measured}}}, \qquad (21)$$

$$C_{detector,measured} = \frac{C_{spectral,measured}}{\sqrt{M_{detector,measured}}}. \qquad (22)$$

The results show a good agreement of prediction model and measurement and are well within the estimated error margins, which are given in the 1σ interval. While the spectral extent of a speckle correlation area at the detector can be derived from the width of μ_{det}, one can also see from the values of $M_{detector}$, that the extent must be small compared to a pixel. This allows for the treatment of the SFA as white noise on detector level, which can be a goal of the instrument design. The uncertainties

Type	$M_{polarization}$	$M_{spectral}$	$M_{detector}$	SFA [%]
Measurement NIR	2	55.9 ± 0.7	$(6.1 \pm 1.8) \times 10^2$	0.38 ± 0.06
Prediction NIR	2	56.5	5.7×10^2	0.39
Measurement SWIR	2	29.9 ± 0.8	$(1.7 \pm 0.4) \times 10^2$	0.99 ± 0.12
Prediction SWIR	2	30.0	1.8×10^2	0.96

Table 2. Comparison SFA results of the measurement chain with the prediction model. Measurement uncertainties are given in the 1σ interval.

$\Delta M_{spectral,measured}$ and $\Delta M_{detector,measured}$ are estimated by considering fluctuations in the two contributors F and Ψ to the field correlation μ_{mn} in Eq. (12) and μ_{det} in Eq. (20). For F the measured standard deviation of each data point in Figure 4 is taken as a measure, which amounts to 1-2 % in the NIR and 2-3 % in the SWIR band. The average fluctuations of Ψ were determined by measuring its width over 120 speckle patterns by autocorrelation to be 1.3 % and 2.7 % in the NIR and SWIR, respectively. The impact on the $M_{spectral,measured}$ and $M_{detector,measured}$ is estimated by a Monte Carlo propagation of both uncertainties. It yields that those two factors are the major contributors to the uncertainty $\Delta M_{spectral}$. This approach also accounts for detector noise, if it caused the fluctuations. The impact on $M_{detector}$, however, is small. It was deduced from Ahn and Fessler (2003), that $\Delta M_{detector}$ is primarily caused by the uncertainty of the standard deviation estimator for a small amount of samples. In our case the sample number equals the amount of detector pixel (see Figure 2), which is $n = 30$ for the NIR and $n = 48$ for the SWIR measurement. The uncertainties presented also reflect the fact that the SFA is not constant over a wavelength range of a few nanometers as illustrated in Figure 5. It shows plots of the SFA scaling with wavelength for a CO2I like instrument in the NIR and SWIR band calculated with the prediction model. The linear dependency only holds under the assumption of a constant dispersion, which is usually not the case for a realistic instrument. However, this result gives a clear indication to the general scaling of the speckle effect.
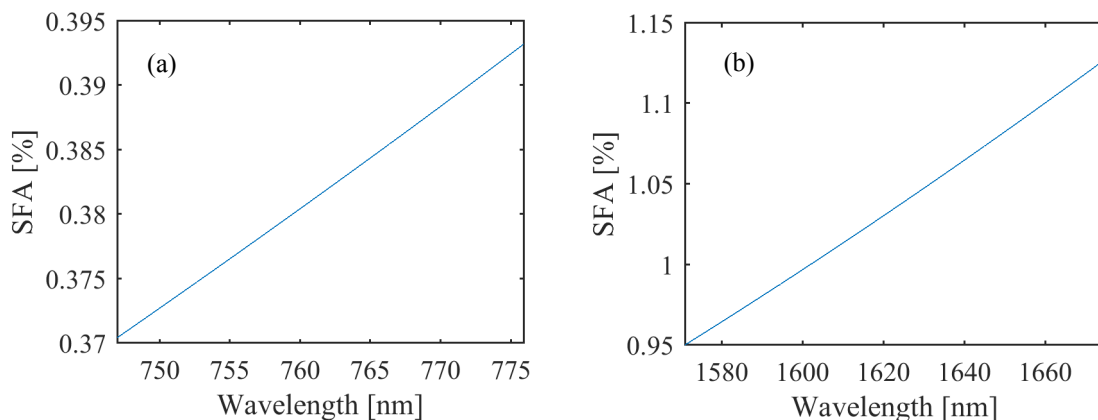
Figure 5. Scaling of the SFA with wavelength of a CO2I like instrument in the (a) NIR and (b) SWIR band using the prediction model.

 ## 6 Conclusions

We demonstrated a comprehensive and numerical approach to quantify diffuser induced spectral features during solar calibration of space imaging spectrometers, which is based on established speckle theory concepts. We compared our prediction results with our current measuring method and observed a good agreement. We also give an indication regarding the wavelength dependency of the effect. The presented speckle averaging mechanisms are not a complete representation of the real

305 in-orbit situation of an instrument. The effect of the Sun's disk, which consists of many incoherent points sources distributed over a 0.5 degree angle, needs to be taken into account as well as the averaging due to the movement of the instrument relative to the Sun. Also, unlike the used laser point sources for the measurements, the Sun's light features an additional orthogonal polarization state, adding two polarization configurations to $M_{polarization}$ in the case of a highly scattering volume diffuser. The presented approach can be used for other diffuser types and optical geometries as well. It provides a solid starting point for

310 future investigations into angular averaging mechanisms, which will complement the description of speckle reduction effects in imaging spectrometers of this type.

Data availability. The datasets generated and/or analyzed for this work are available from the corresponding author on reasonable request, subject to confirmation of Airbus Defence and Space GmbH.

Author contributions. Florian Richter was responsible for the acquisition and the analysis of the measurement data supported by all co-

315 authors. Florian Richter developed the prediction model supported by Jerome Caron with insights about polarization contributions, spectral averaging, and properties of coherency matrices and revised by Corneli Keim and Mark Wenig. Florian Richter prepared the manuscript with contributions and critical revision from all co-authors.

Competing interests. The authors declare, that they have no conflict of interest.

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